Dark Matter Searches with IceCube / Neutrinos

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KIT Seminar

Coma Cluster

Outline

Motivation

Neutrino Telescopes and IceCube

Selected Searches

- Self-annihilating Dark Matter in the Galactic Halo
- Astrophysical Neutrinos and Dark Matter Decay
- Dark Matter Captured in the Sun

Conclusions and Outlook

Coma Cluster

Motivation

The Dark Matter Mystery

Coma Cluster

The Dark Matter Mystery

Coma Cluster

- Since Zwicky observed the Coma cluster evidence has hardened
  - Structure formations - Cosmological simulations
  - Gravitational lensing
  - Rotation curves
  - Cosmic microwave background
  - ...

The Dark Matter Mystery

Coma Cluster

- Since Zwicky observed the Coma cluster evidence has hardened
  - Structure formations - Cosmological simulations
  - Gravitational lensing
  - Rotation curves
  - Cosmic microwave background
  - ...

- Dark Matter already gravitationally "observed", but ...
  - What is it ?
  - What are it's properties ?

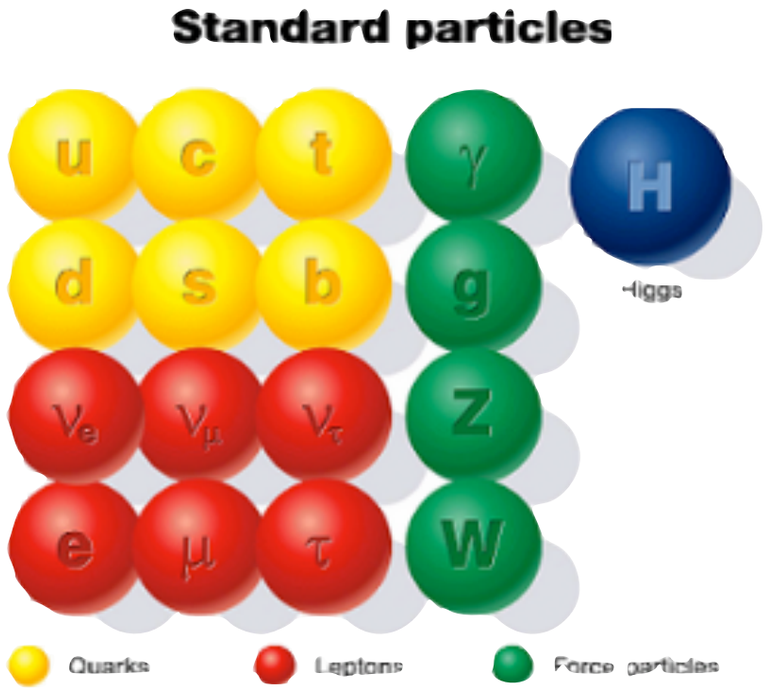
Weakly Interacting Massive Particle (χ)

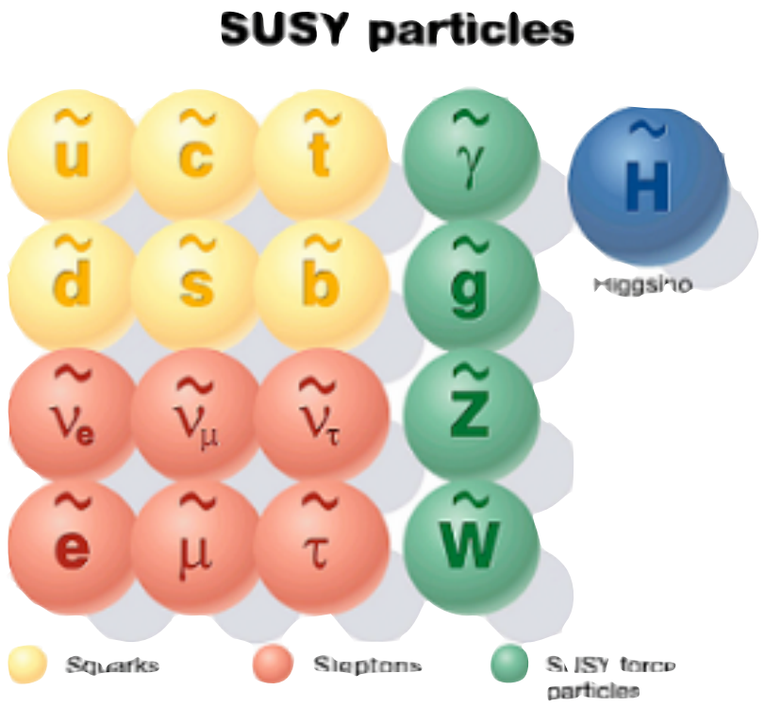
- **Observational Evidence for Dark Matter points to**
  - Non-baryonic
  - Cold massive
  - Not strongly interacting

  } WIMP

  - Stable (long lived)

- **WIMPs often arise naturally in extensions to the Standard Model of Particle Physics: Supersymmetry, ...**

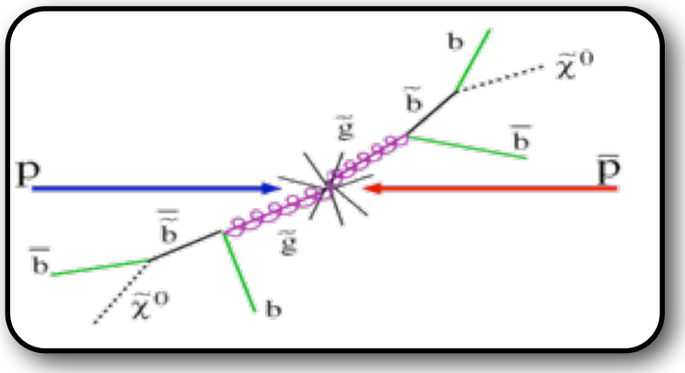




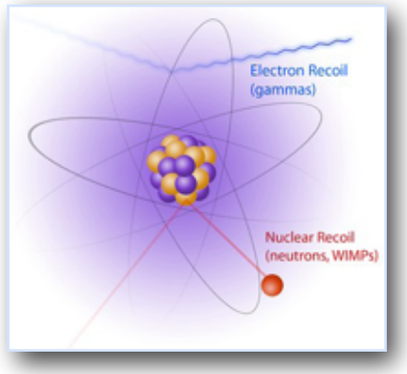
Severe constraints on the WIMP hypothesis from collider searches, but still a good assumption as a generic dark matter candidate particle

Searches for WIMPs

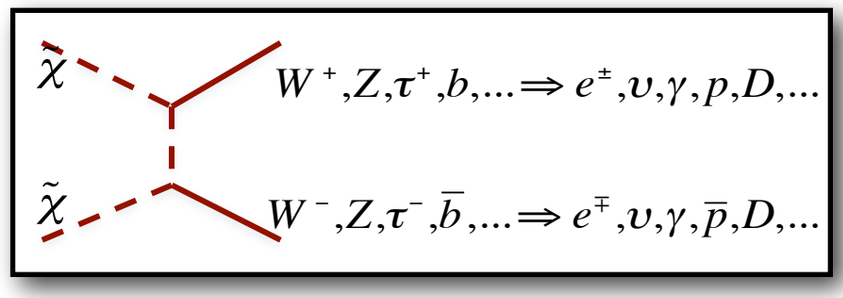
Production



Scattering



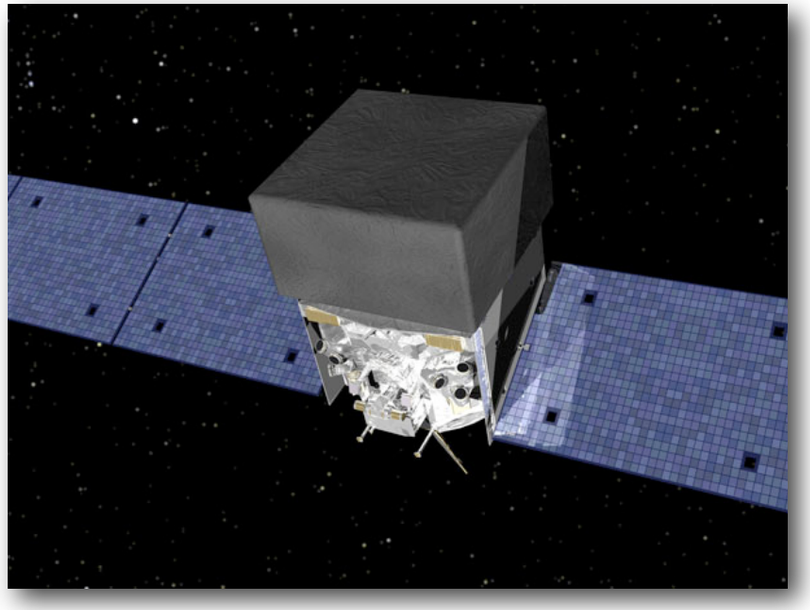
Annihilation



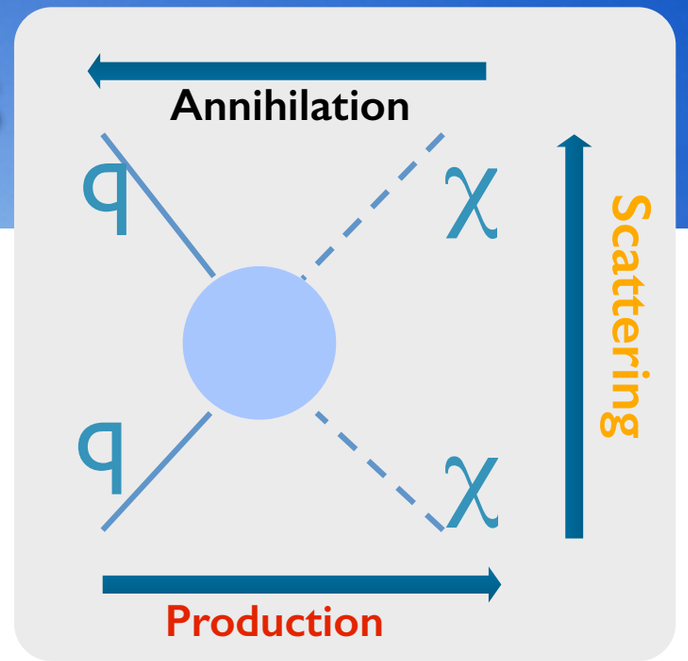
Colliders

Direct

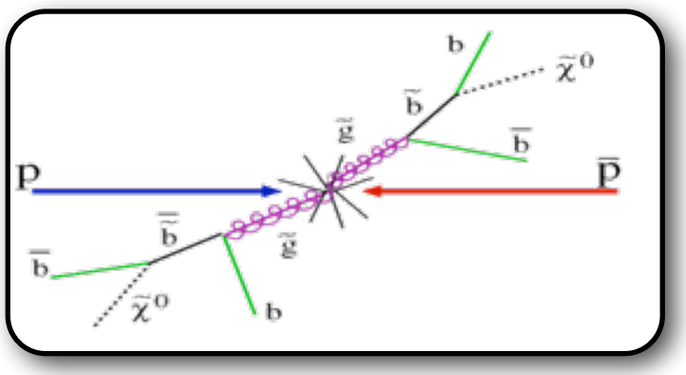
Indirect



Role of Neutrinos

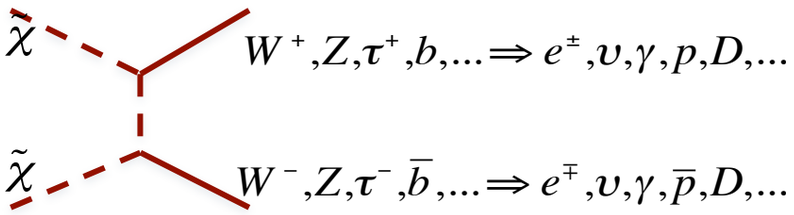


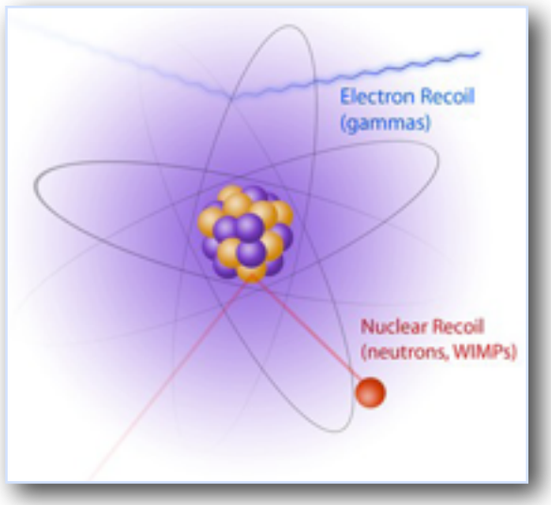
WIMP - Weakly Interacting Massive Particle



- **Production**
  - Colliders
- **Indirect Searches**
  - Dark Matter Decay

Dark Matter Lifetime





- Annihilation of Dark Matter in Galactic Halo, ...
  - Gamma-rays, electrons, neutrinos, anti-matter, ...
- Annihilation signals from WIMPs captured in the Sun (or Earth)
  - Neutrinos
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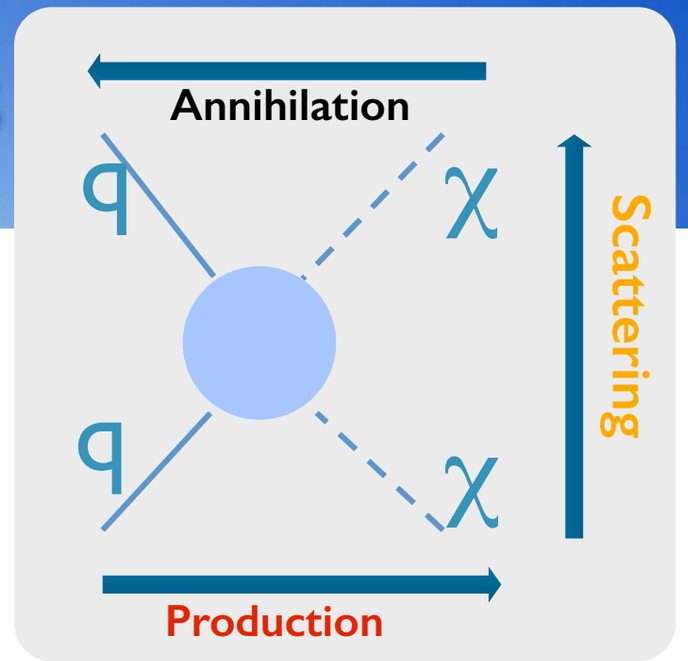
  - WIMP scattering of nucleons → Nuclear recoils

Dark Matter
Self-annihilation cross section

DM -Nucleon
Scattering cross section

Role of Neutrinos



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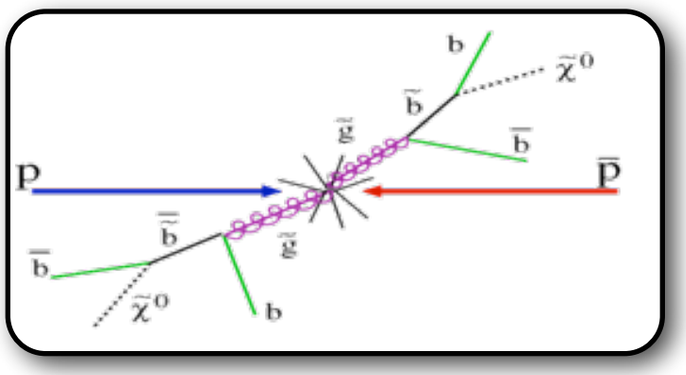

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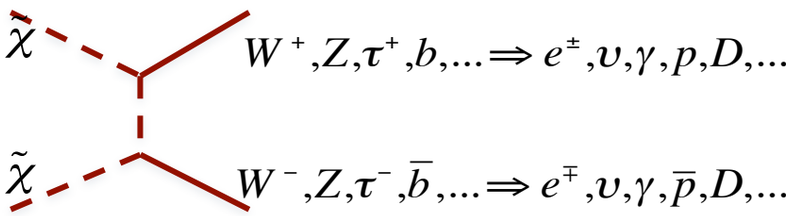

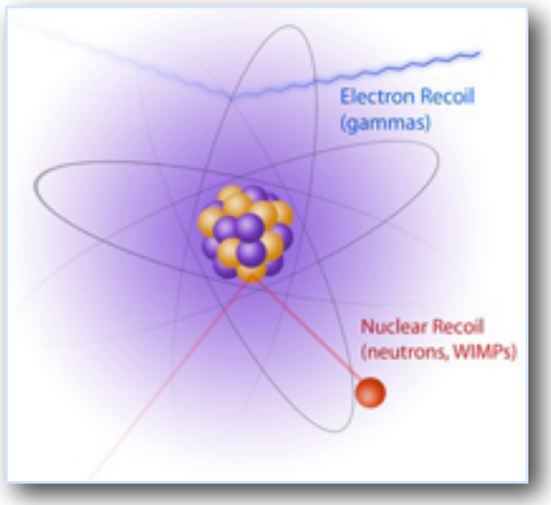

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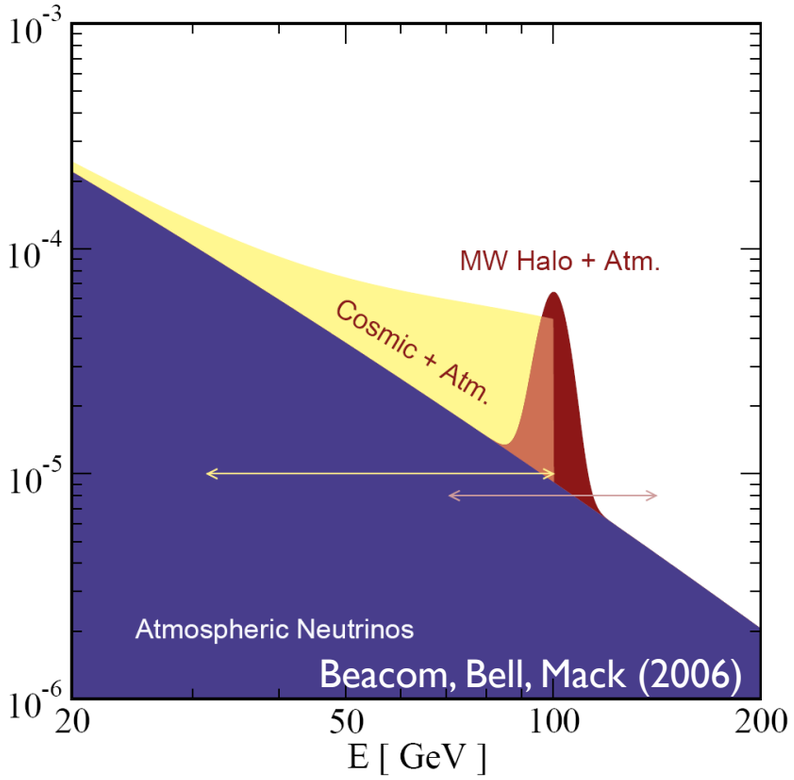
Self-annihilation cross section
DM -Nucleon Scattering cross section

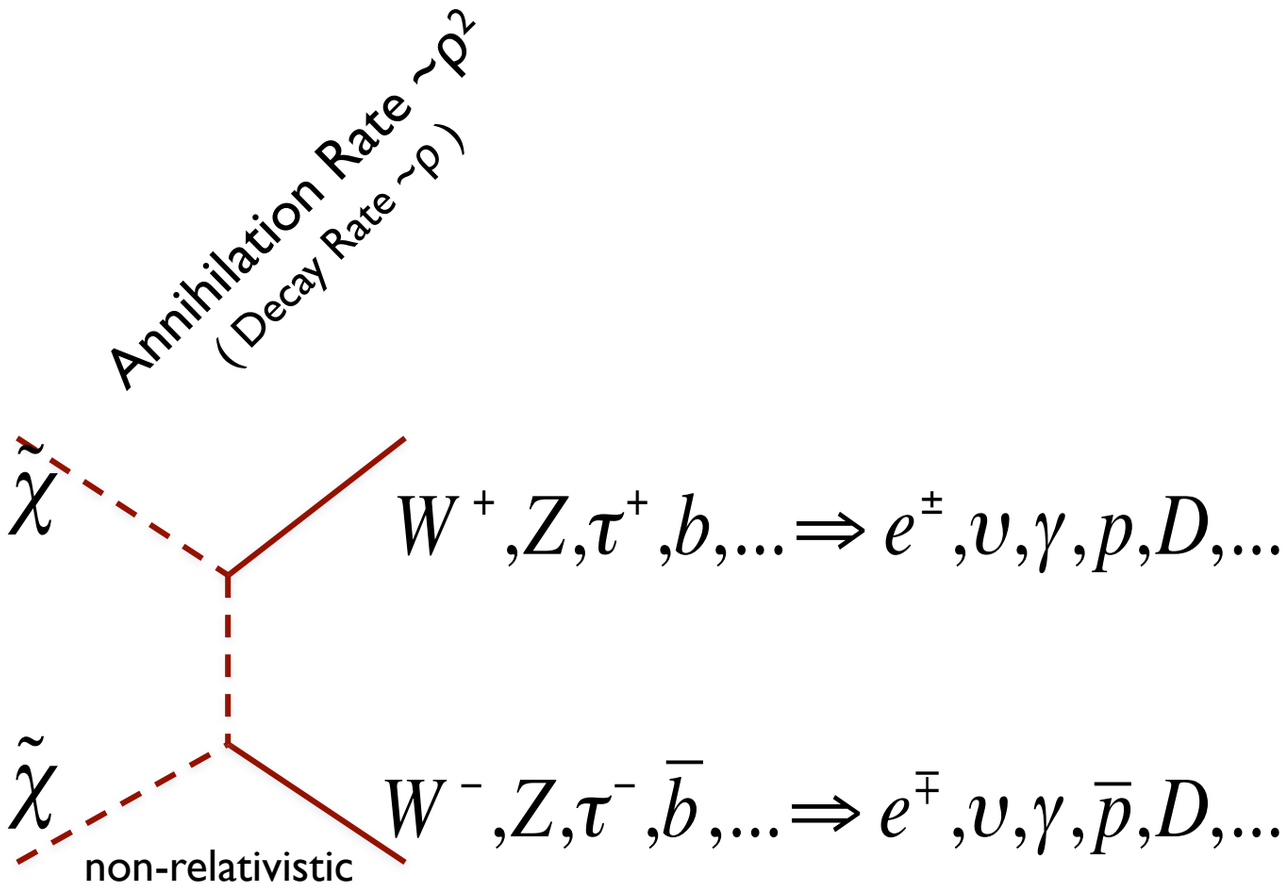
Dark Matter Signals

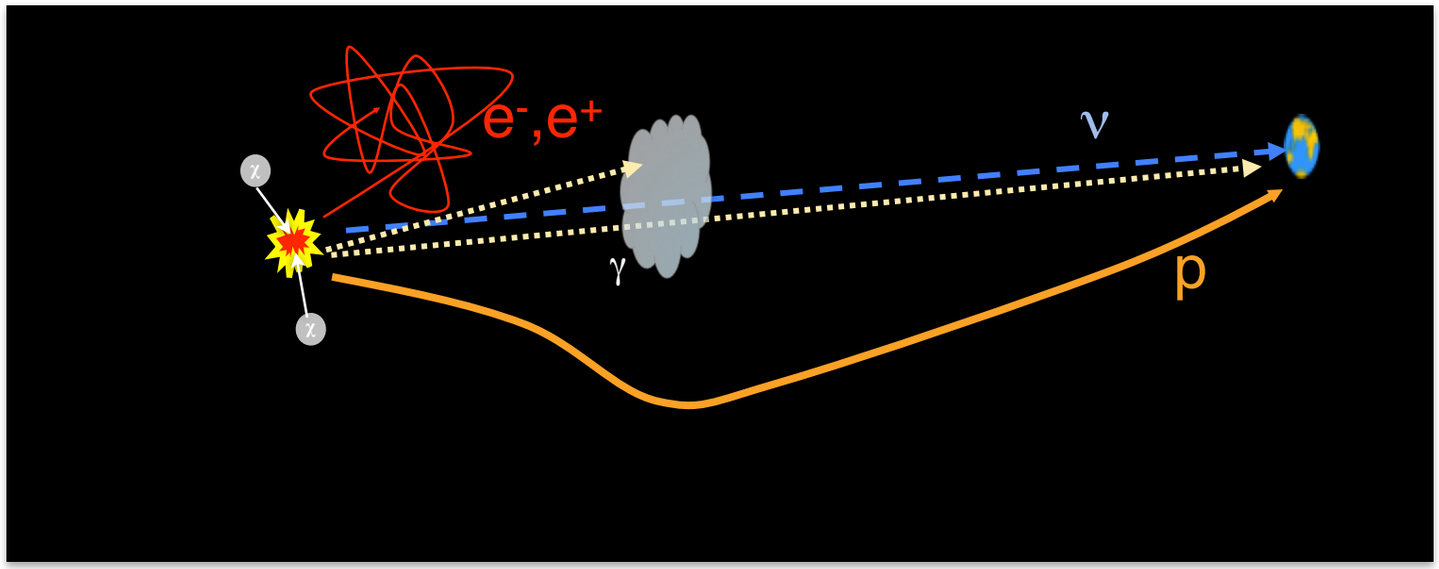
- Identify overdense regions of dark matter

  ⇒self-annihilation can occur at significant rates

- Pick prominent Dark Matter target
- Understand / predict backgrounds
- Exploit features in the signal to better distinguish against backgrounds







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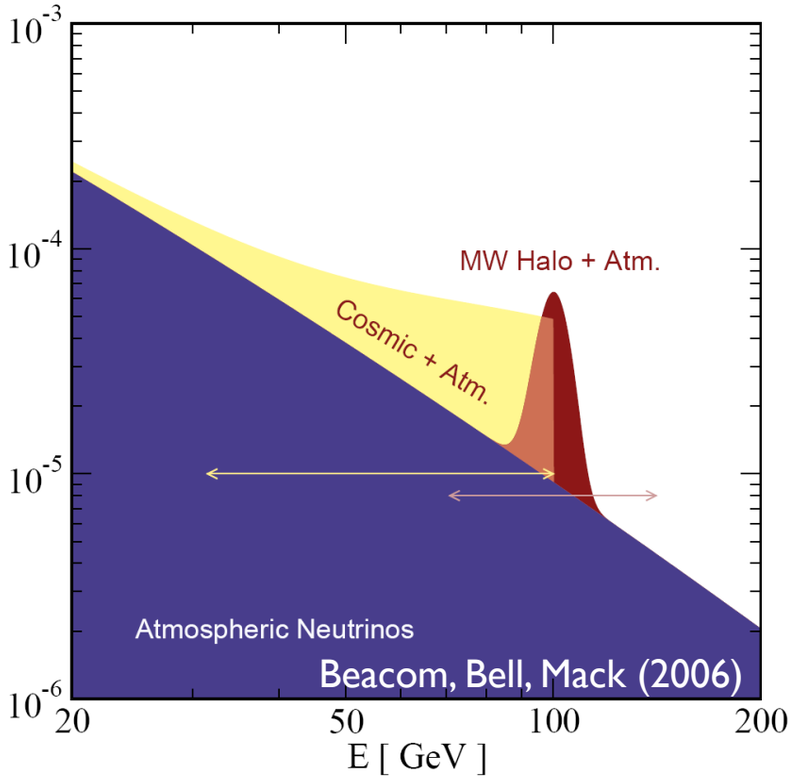
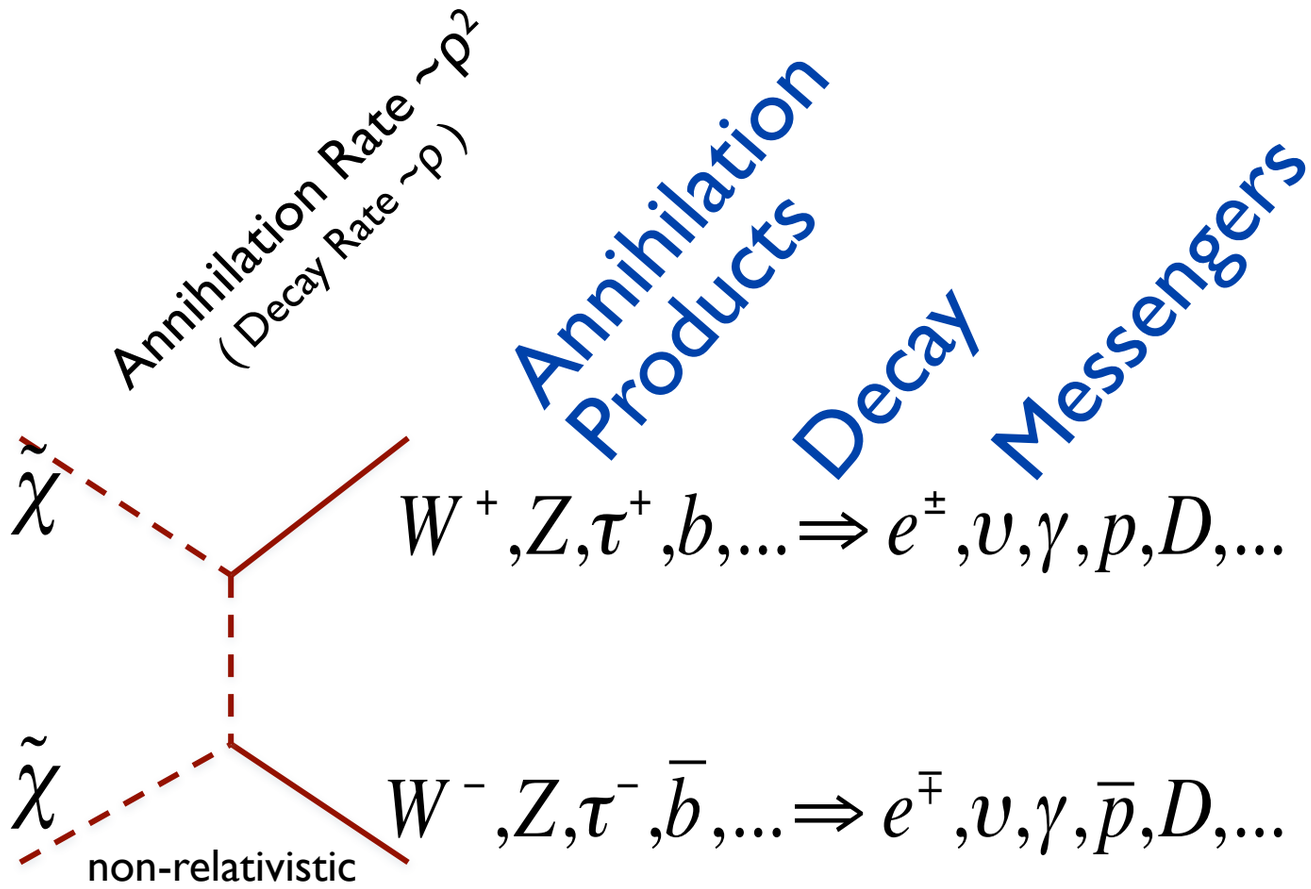
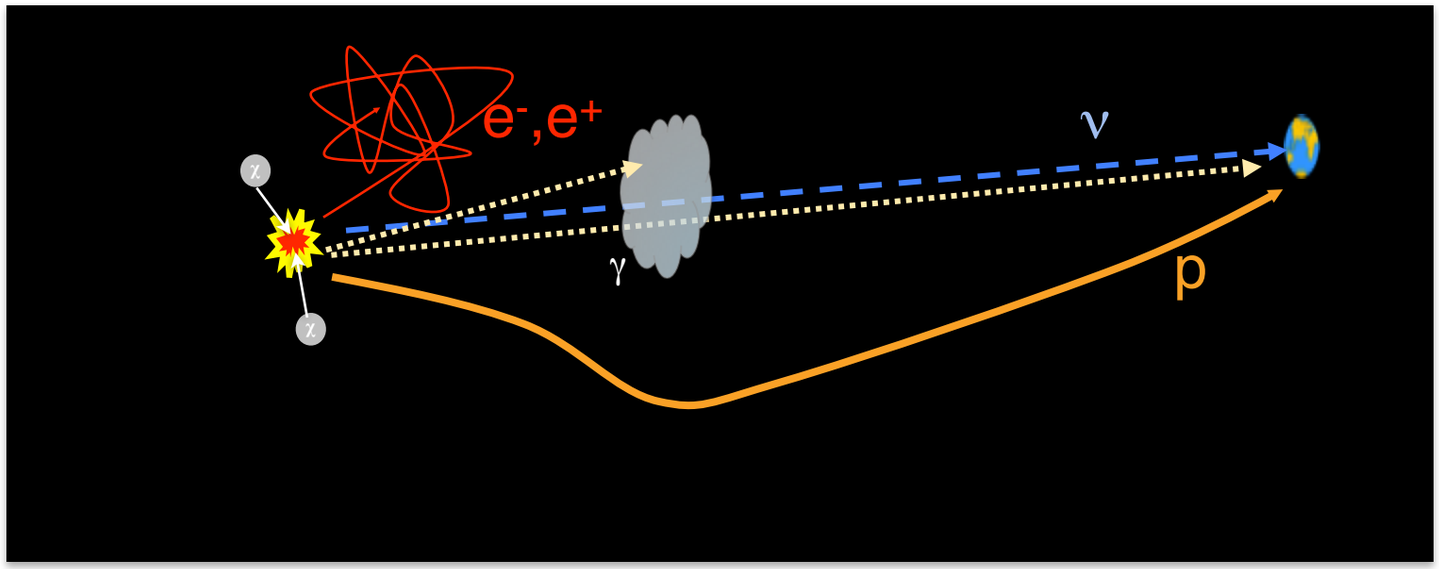
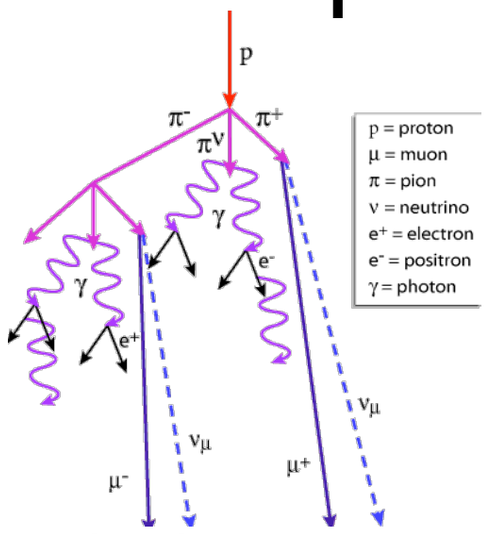






Sources of High Energy Neutrinos

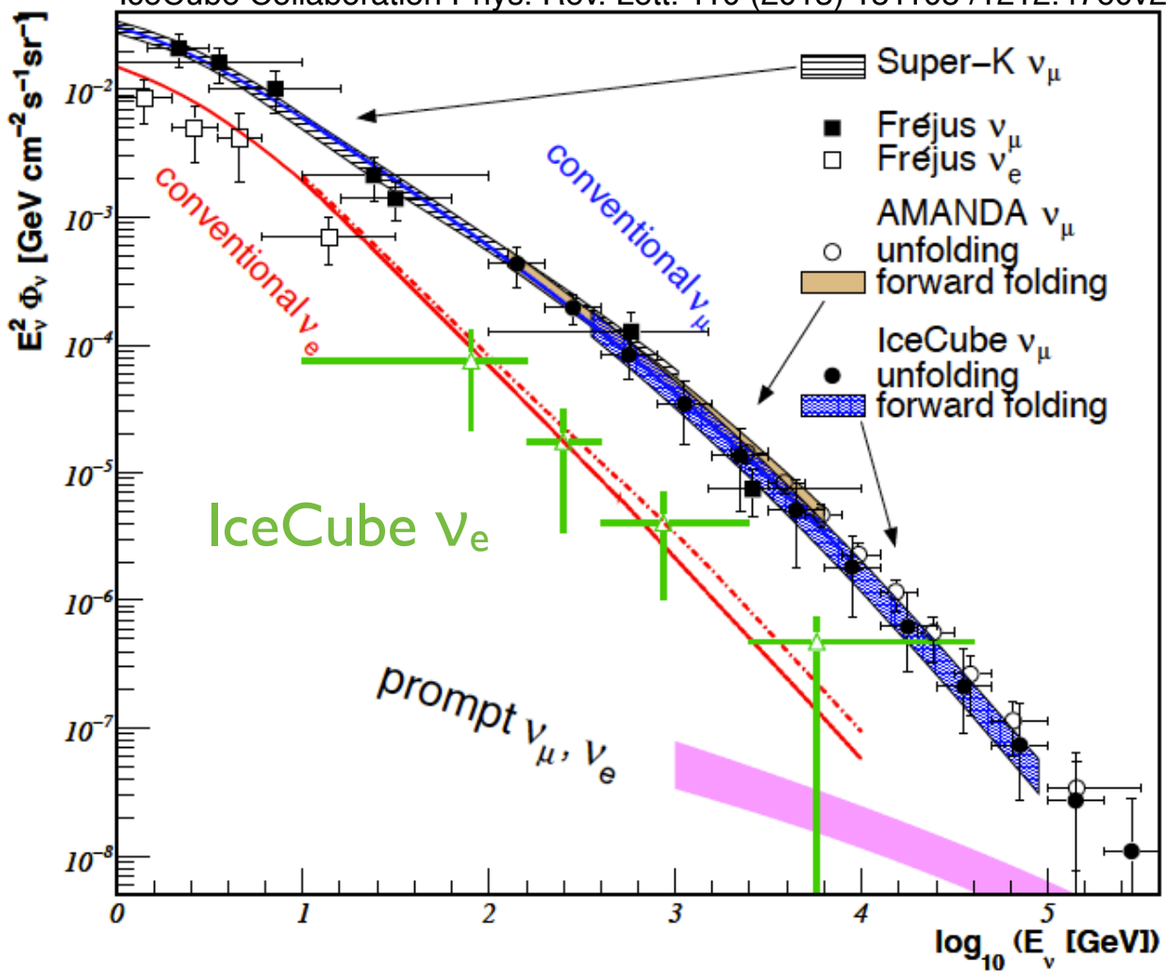
Atmospheric Neutrinos



Cosmic rays interact in the upper atmosphere:

$p + A \rightarrow \pi^{\pm} (K^{\pm}) +$ other hadrons ... $\pi^{+} \rightarrow \mu^{+}\nu_{\mu} \rightarrow e^{+}\nu_{e}\nu_{\mu}\nu_{\mu}$

IceCube Collaboration Phys. Rev. Lett. 110 (2013) 151105 /1212.4760v2

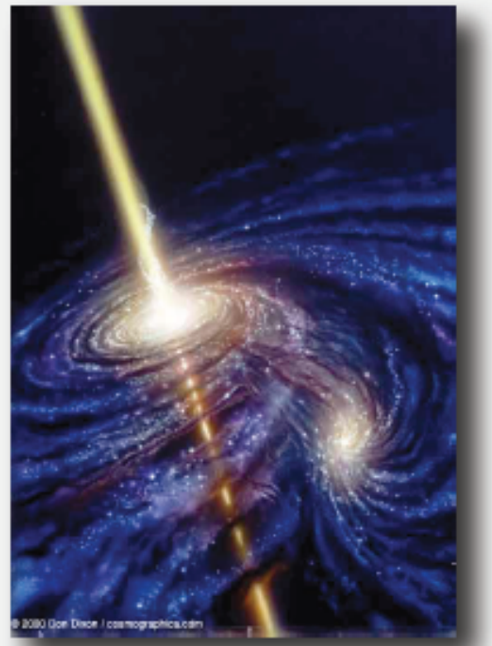


Astrophysical

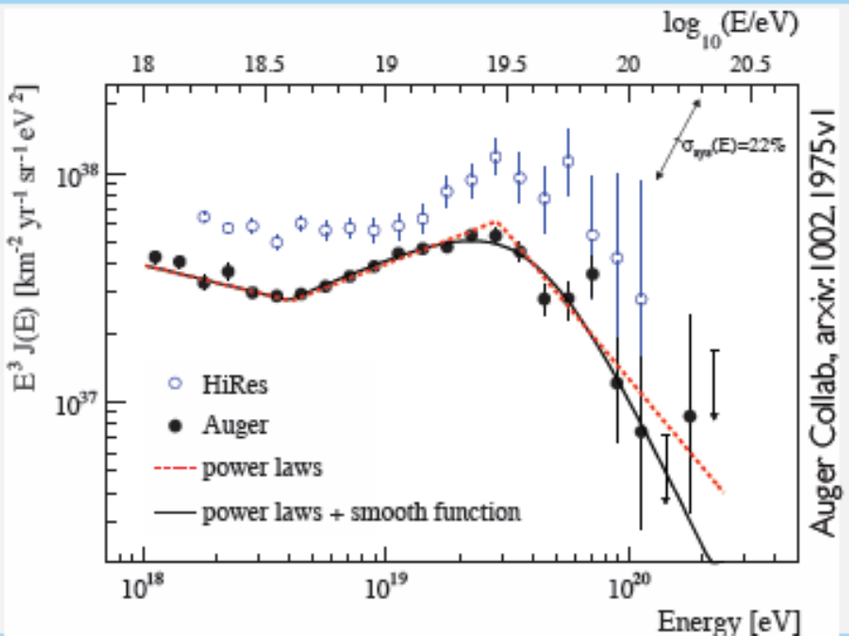
$p + (p,\gamma) \rightarrow \pi^{\pm} \rightarrow \nu$

Gamma-ray Bursts

Active Galactic Nuclei



$p+\gamma_{3K} \rightarrow \Delta^{+}$ (GZK)



Principle of an optical Neutrino Telescope

Array of optical sensors capture the light

$\gamma_{č}$

Cherenkov Radiation

41°

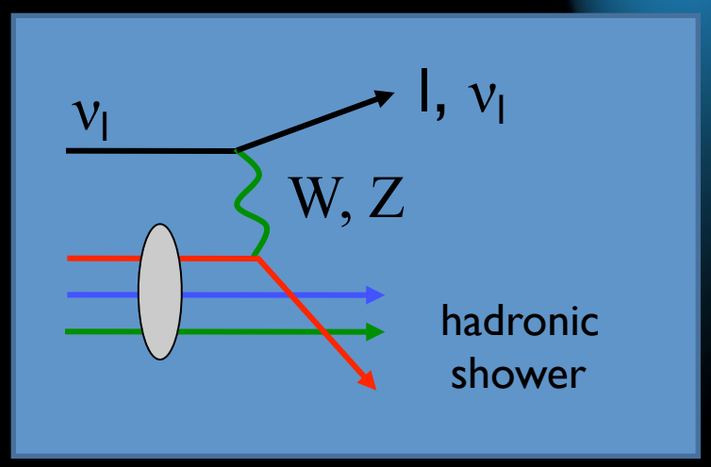
Muon

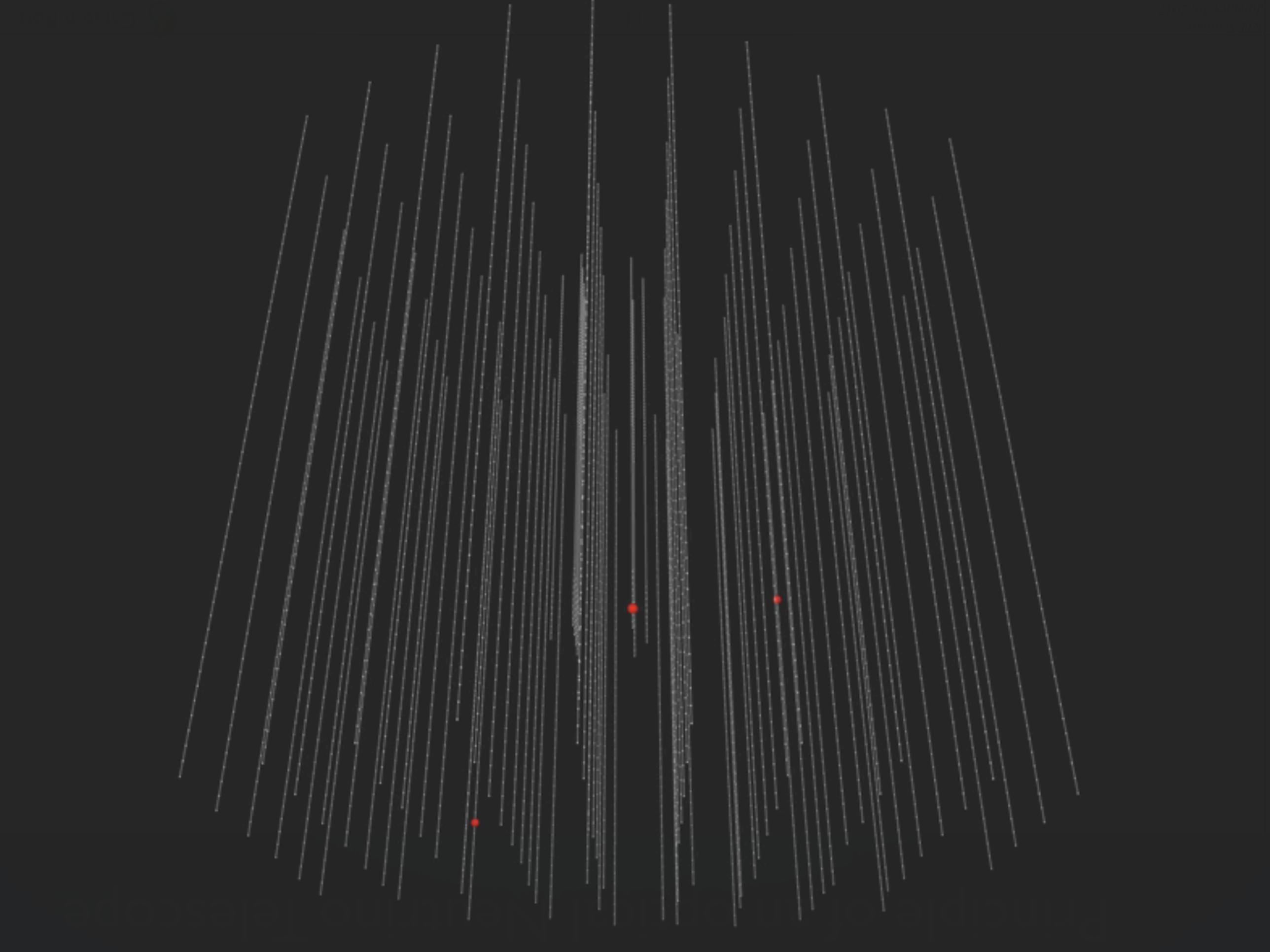
μ

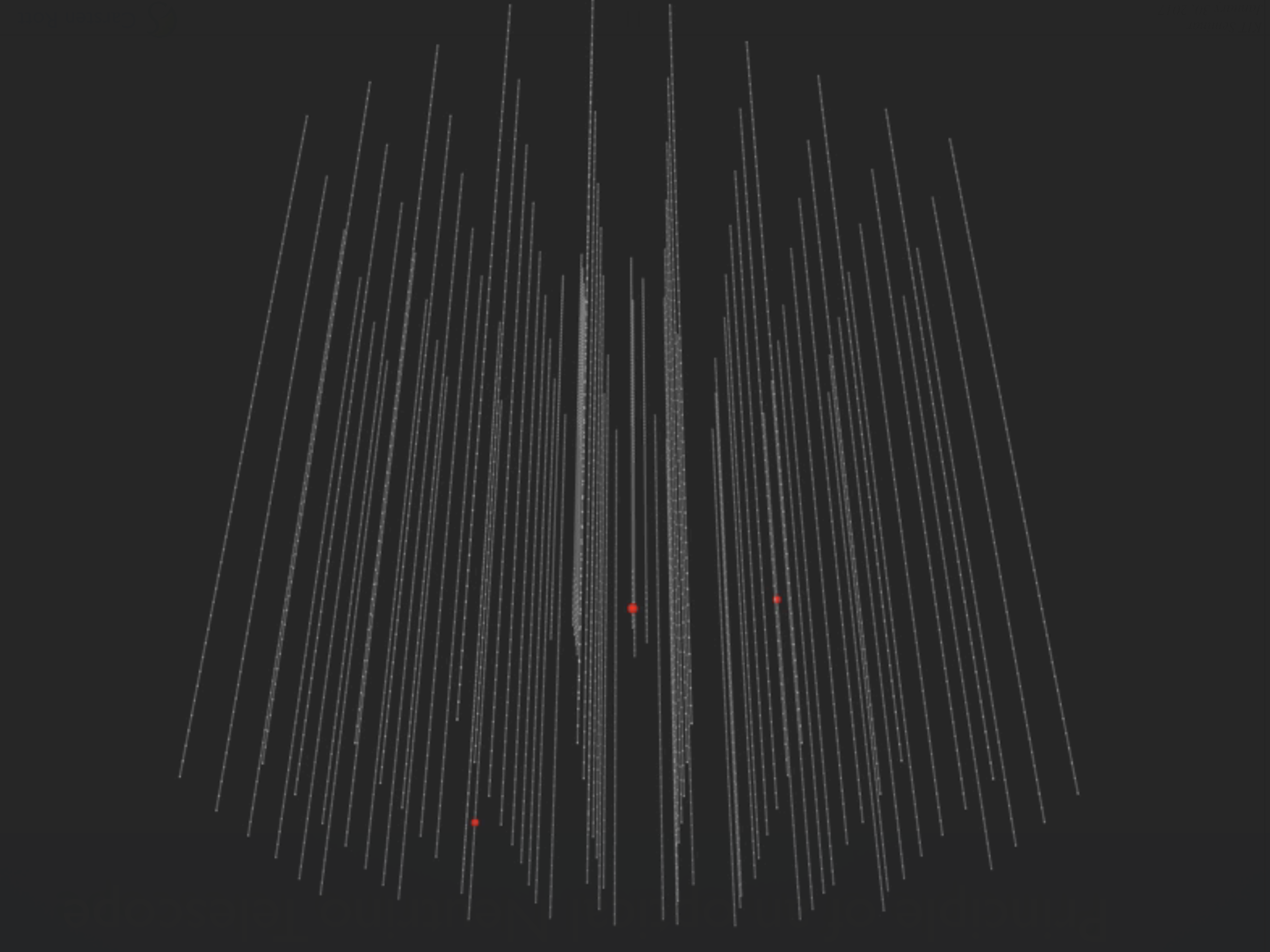
Muon Neutrino

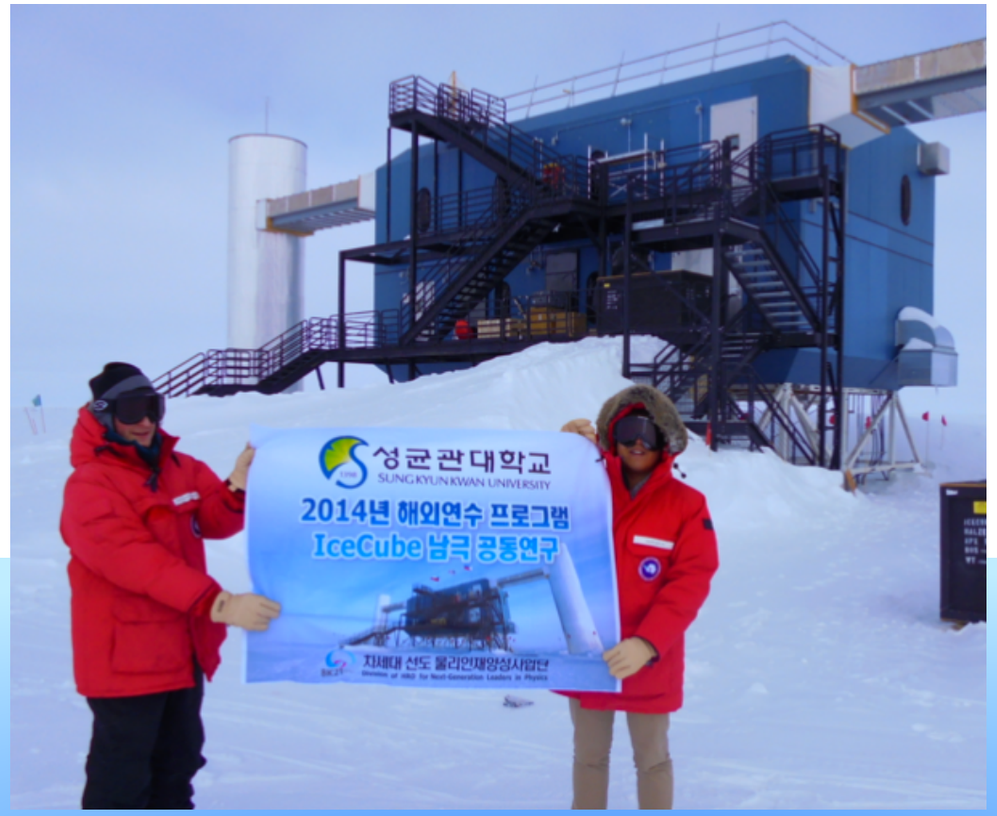
ν

interaction



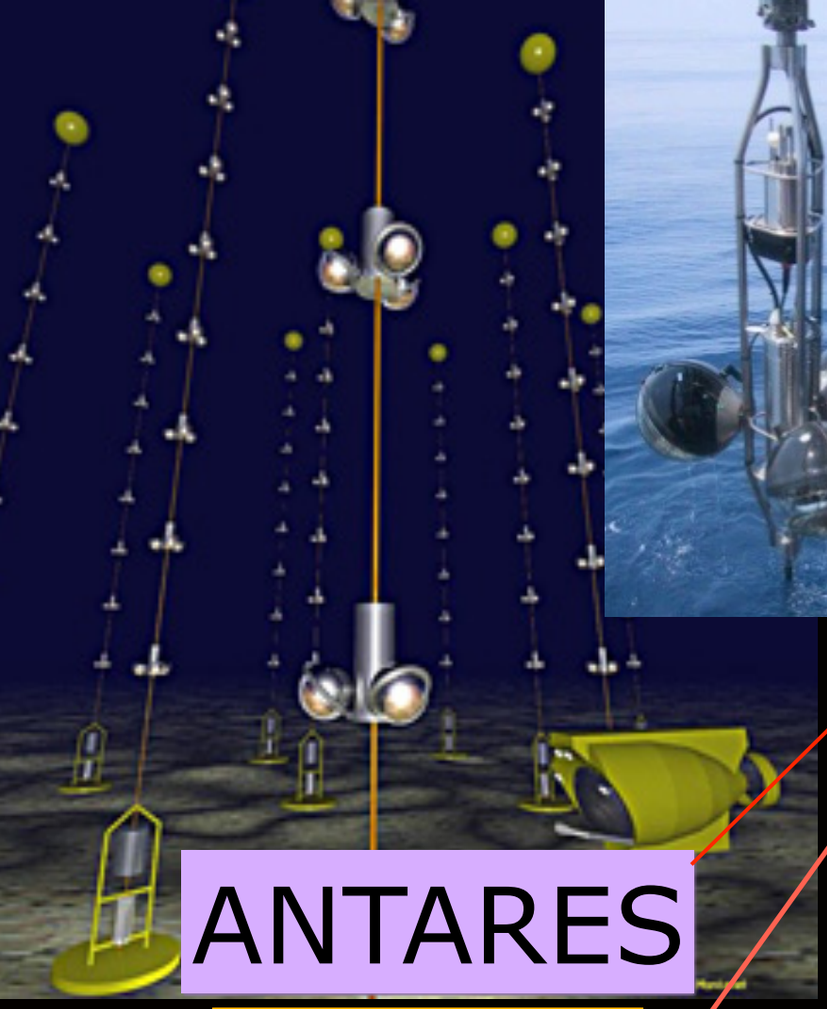






Neutrino Telescopes

Large Water Cherenkov Neutrino Detectors



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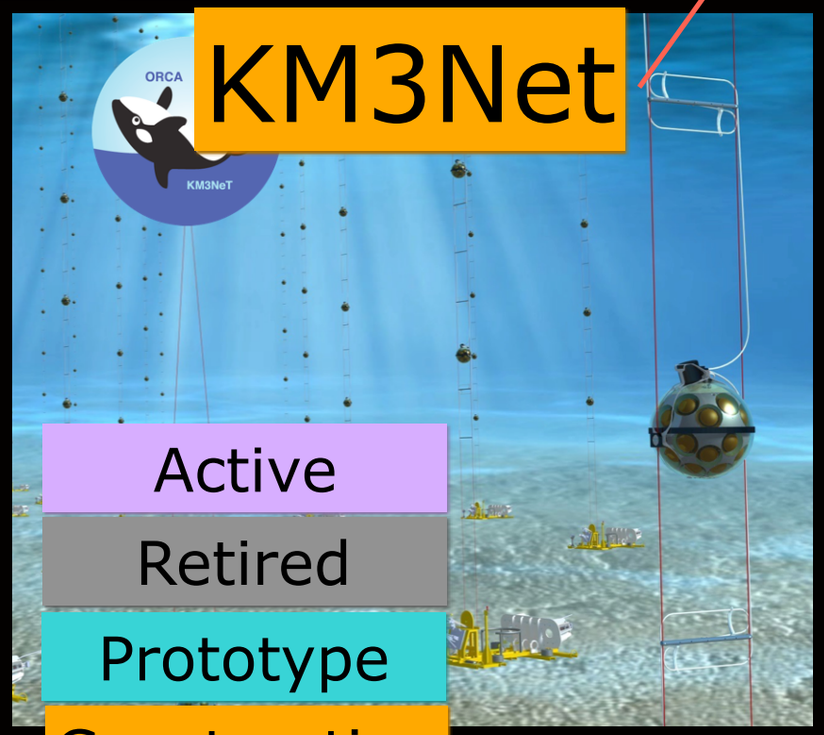


KM3Net

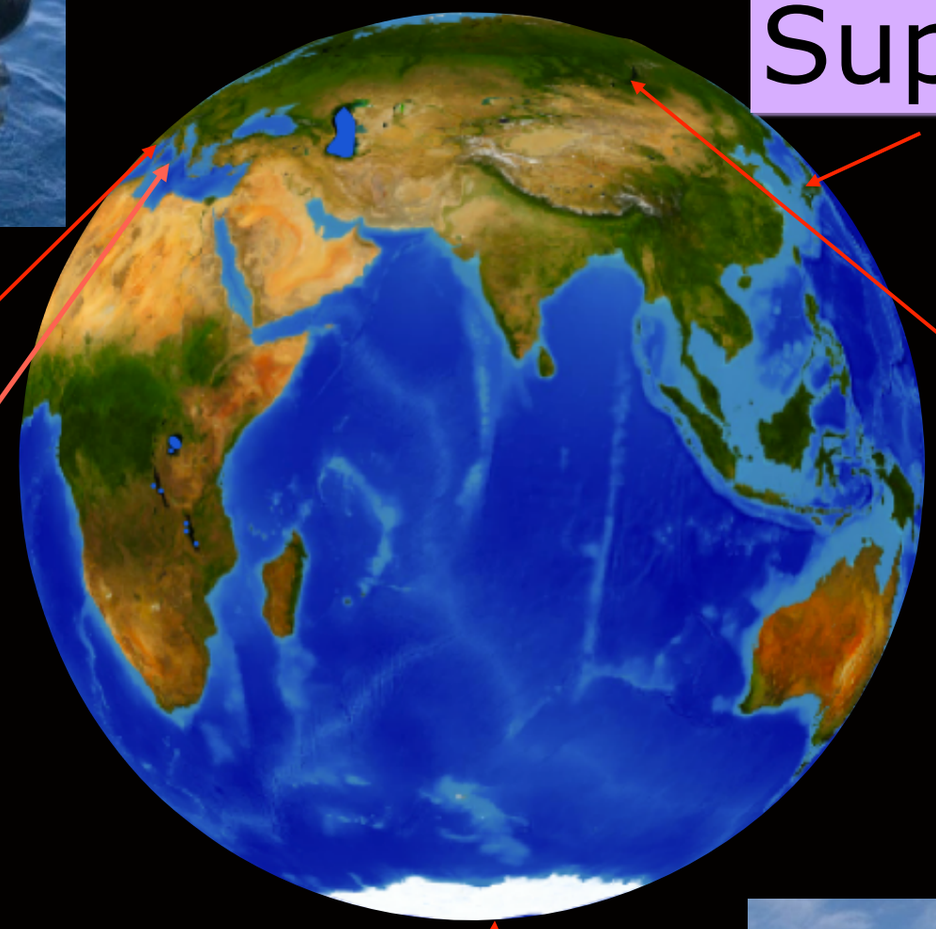
Active

Retired

Prototype

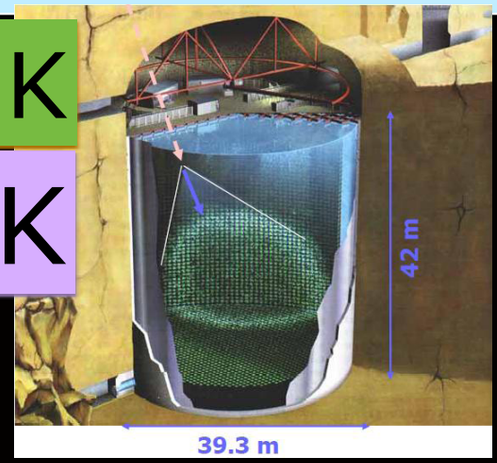
Construction

Planned



Hyper-K

Super-K



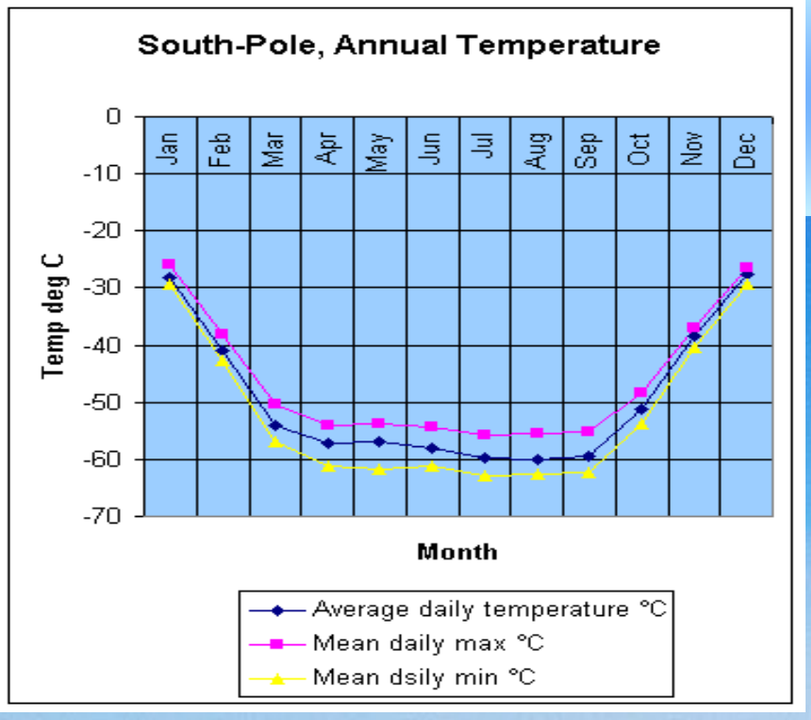
Lake Baikal

GVD

IceCube

Gen2/PINGU

Laboratory at the South Pole



Geographic South Pole

Amundsen Scott South Pole Station

Road to work

Skiway

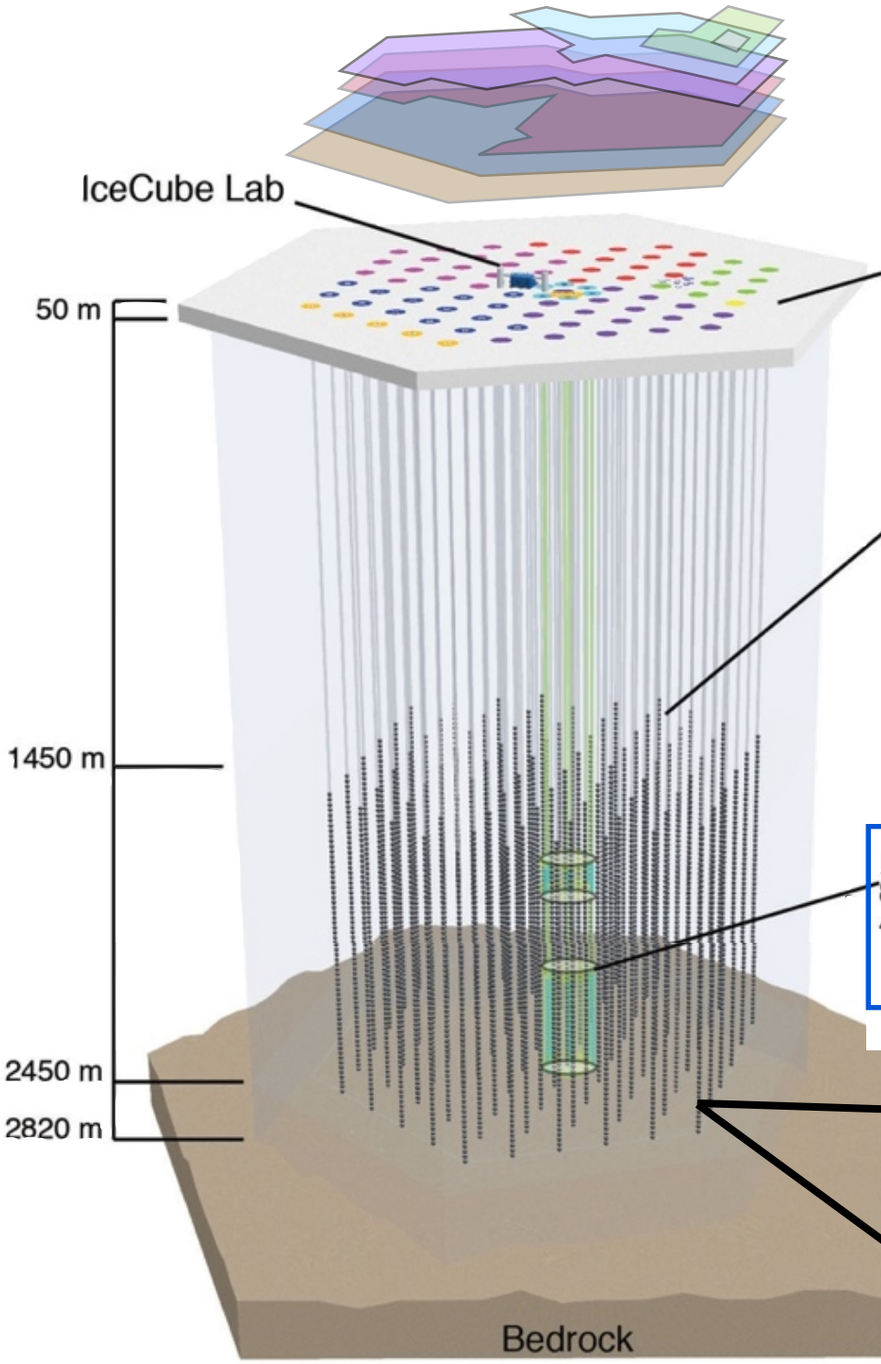
1 km

IceCube

The IceCube Neutrino Telescope

- Gigaton Neutrino Detector at the Geographic South Pole
- 5160 Digital optical modules distributed over 86 strings
- Completed in December 2010, start of data taking with full detector May 2011
- Data acquired during the construction phase has been analyzed
- Neutrinos are identified through Cherenkov light emission from secondary particles produced in the neutrino interaction with the ice

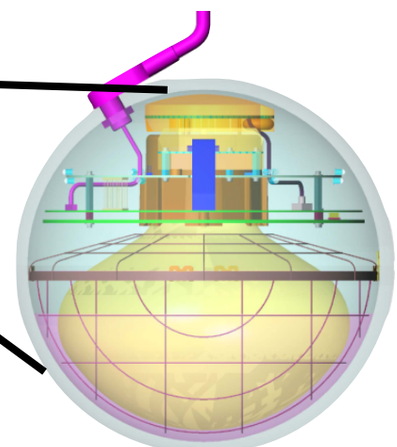
Strings	Dataset
1	2005-2006
9	2006-2007
22	2007-2008
40	2008-2009
59	2009-2010
73 + 6	2010-2011
78 + 8	2011 - .....



IceTop
81 Stations, each with
2 IceTop Cherenkov detector tanks
2 optical sensors per tank
324 optical sensors

IceCube Array
86 strings including 8 DeepCore strings
60 optical sensors on each string
5160 optical sensors
December, 2010: Project completed, 86 strings
E_{Thr} ~ 100 GeV

DeepCore
8 strings-spacing optimized for lower energies
480 optical sensors
E_{Thr} ~ 10 GeV



Dark Matter Searches

- **Galactic Center is 29° above the horizon**
- **Sun is at +/- 23°**

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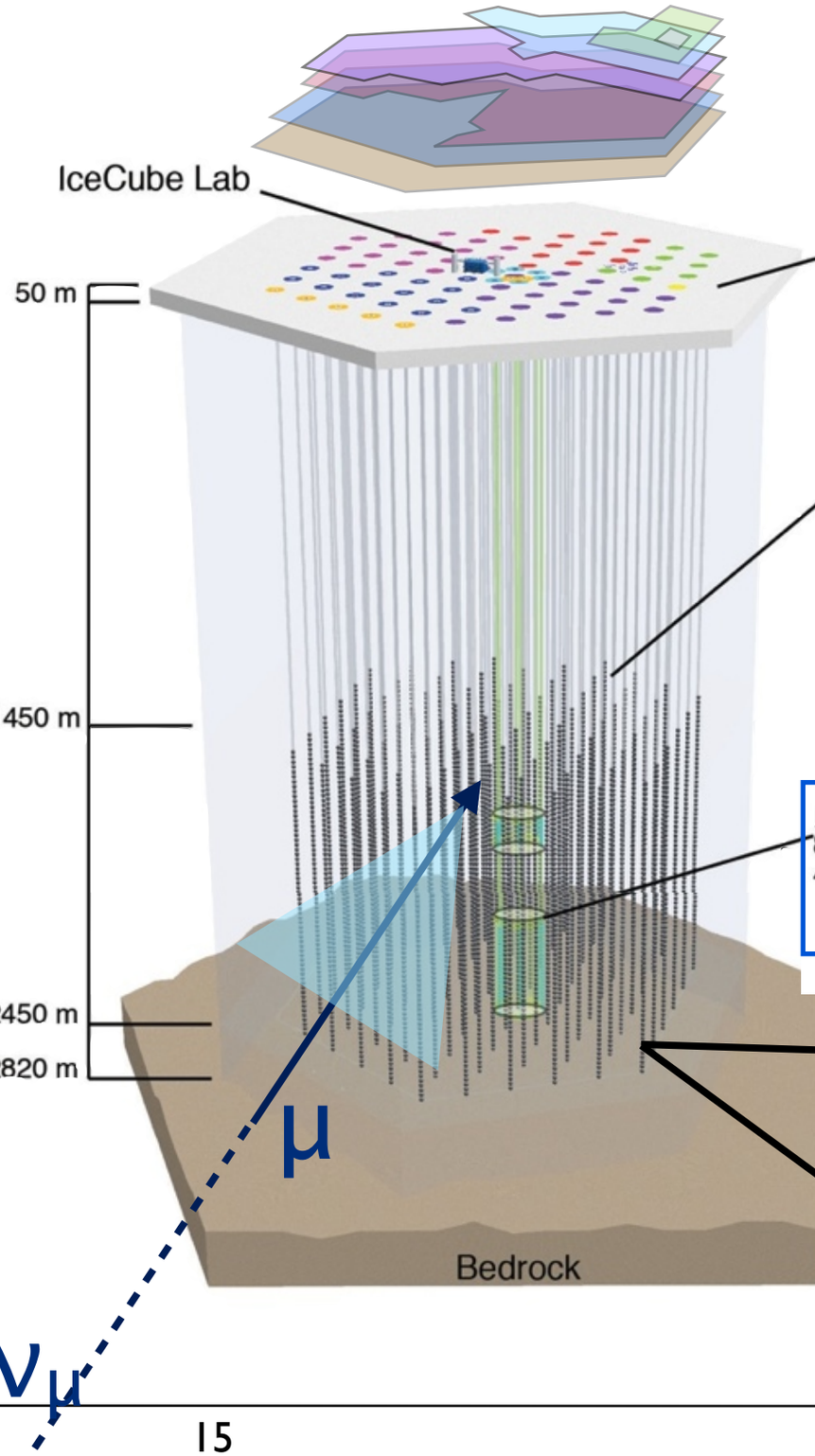

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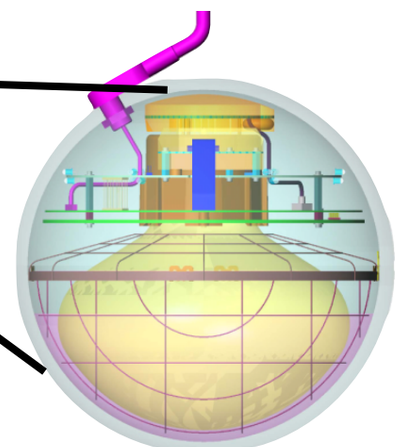

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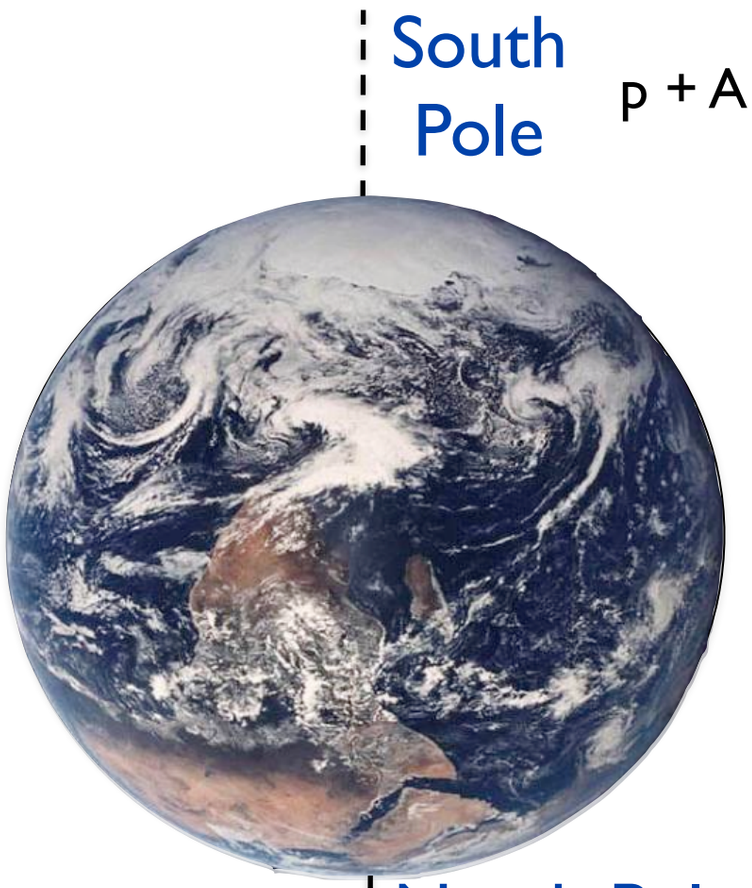
Signals in IceCube



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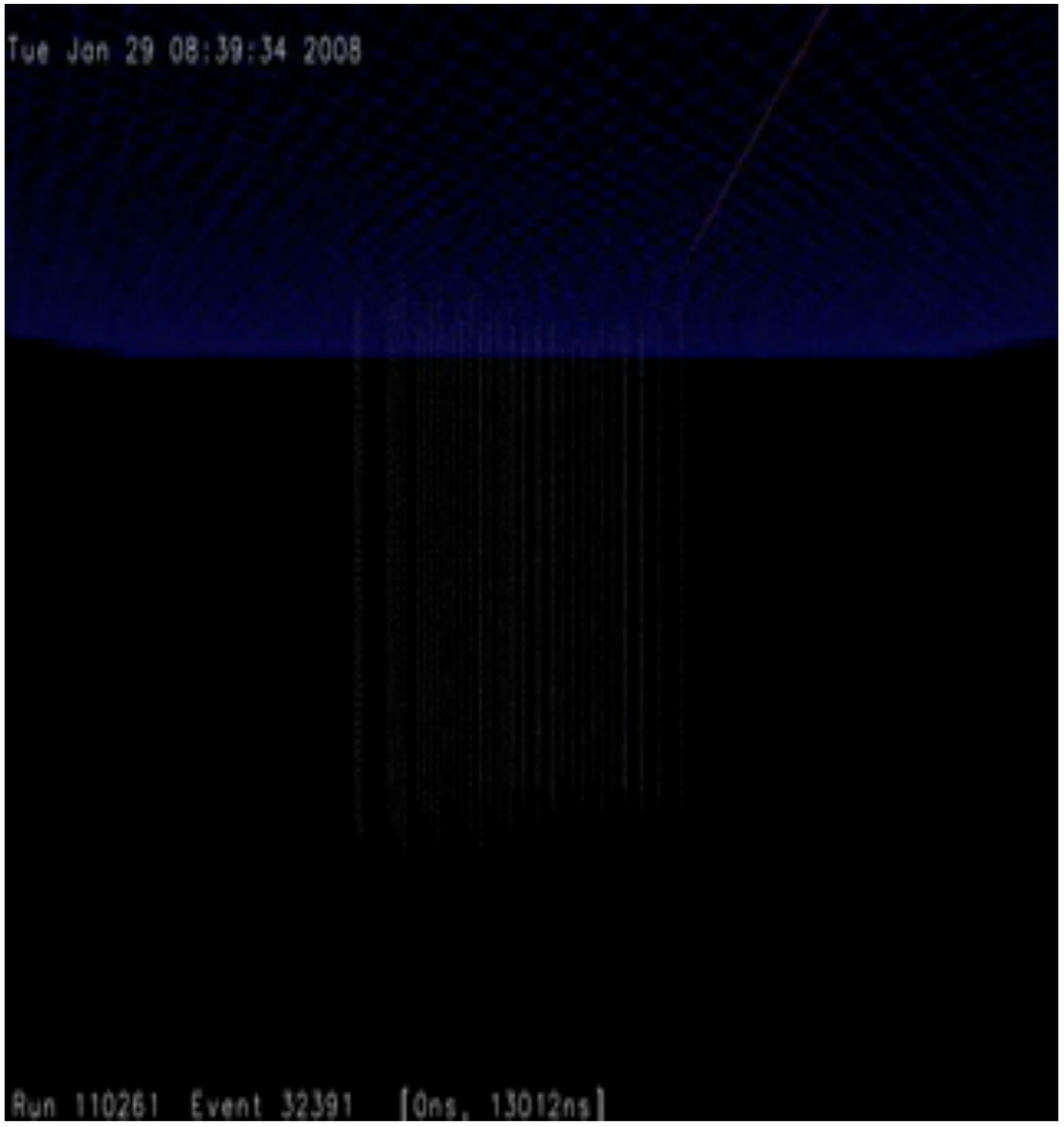

- Up-going events can be used to obtain "clean" neutrino sample
  - Earth is used as muon filter
- Atmospheric neutrinos create irreducible neutrino background to extra terrestrial neutrino fluxes

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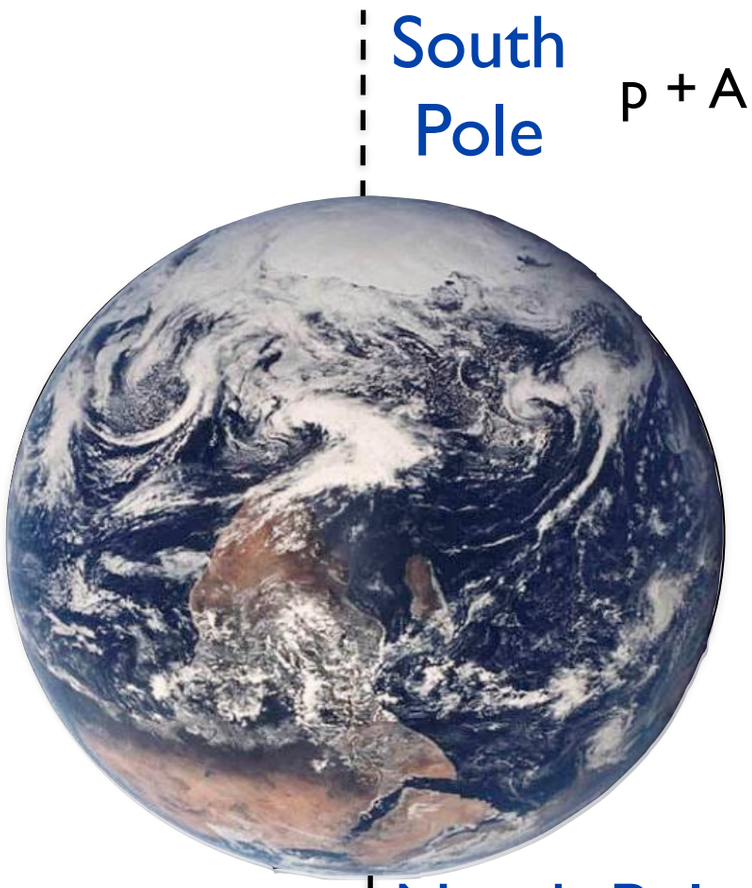

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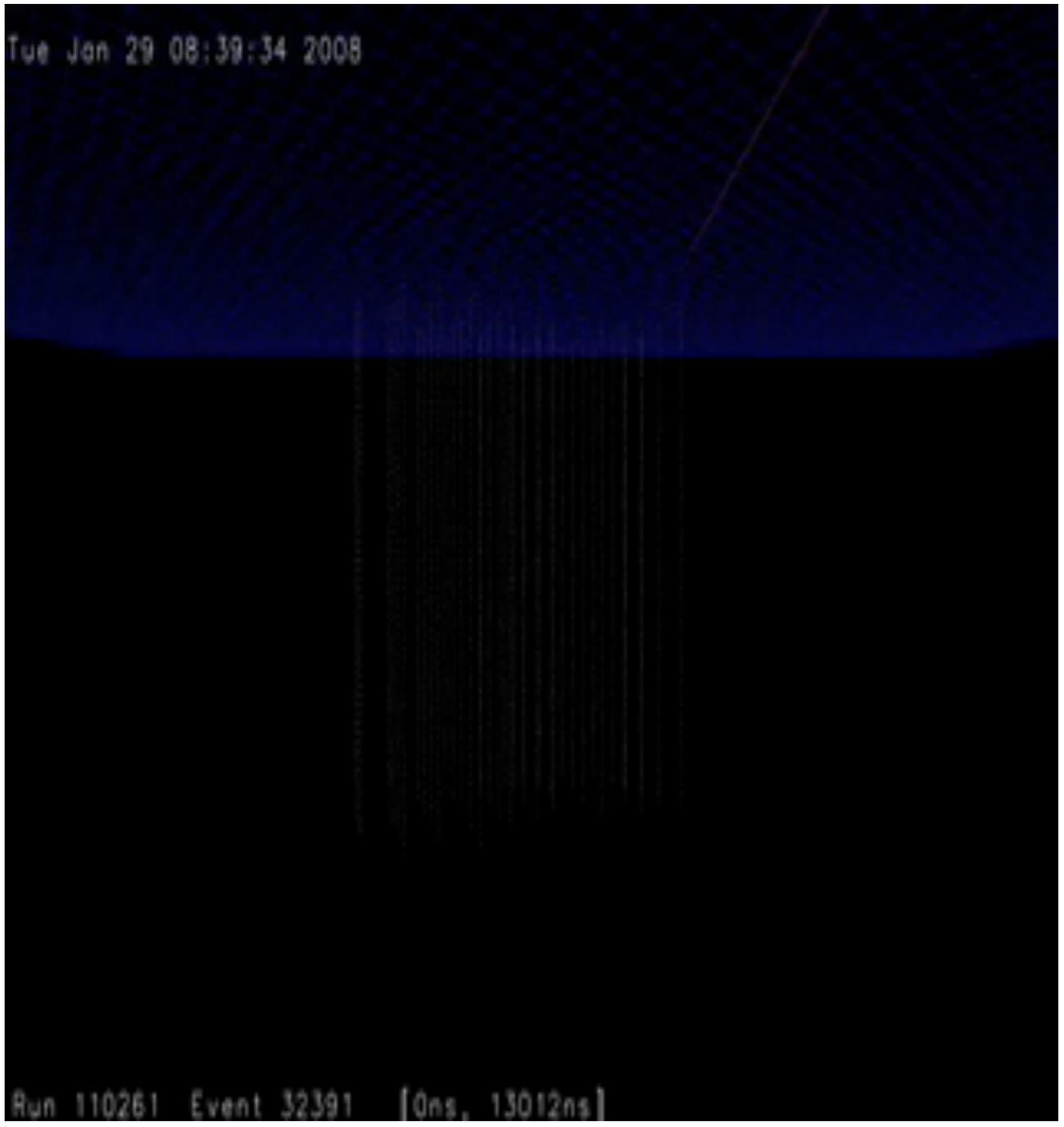

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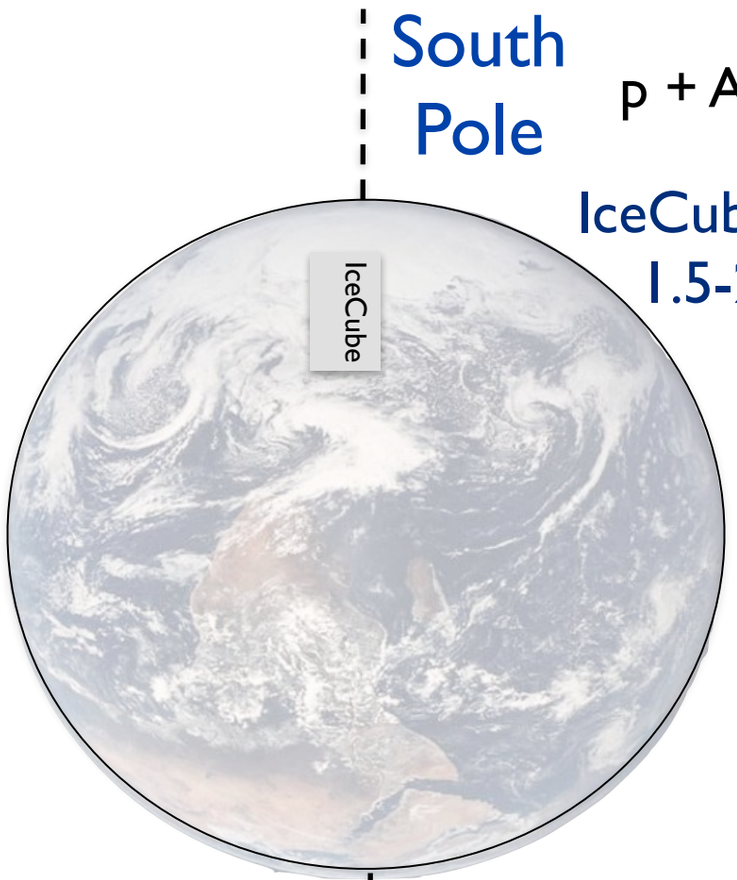

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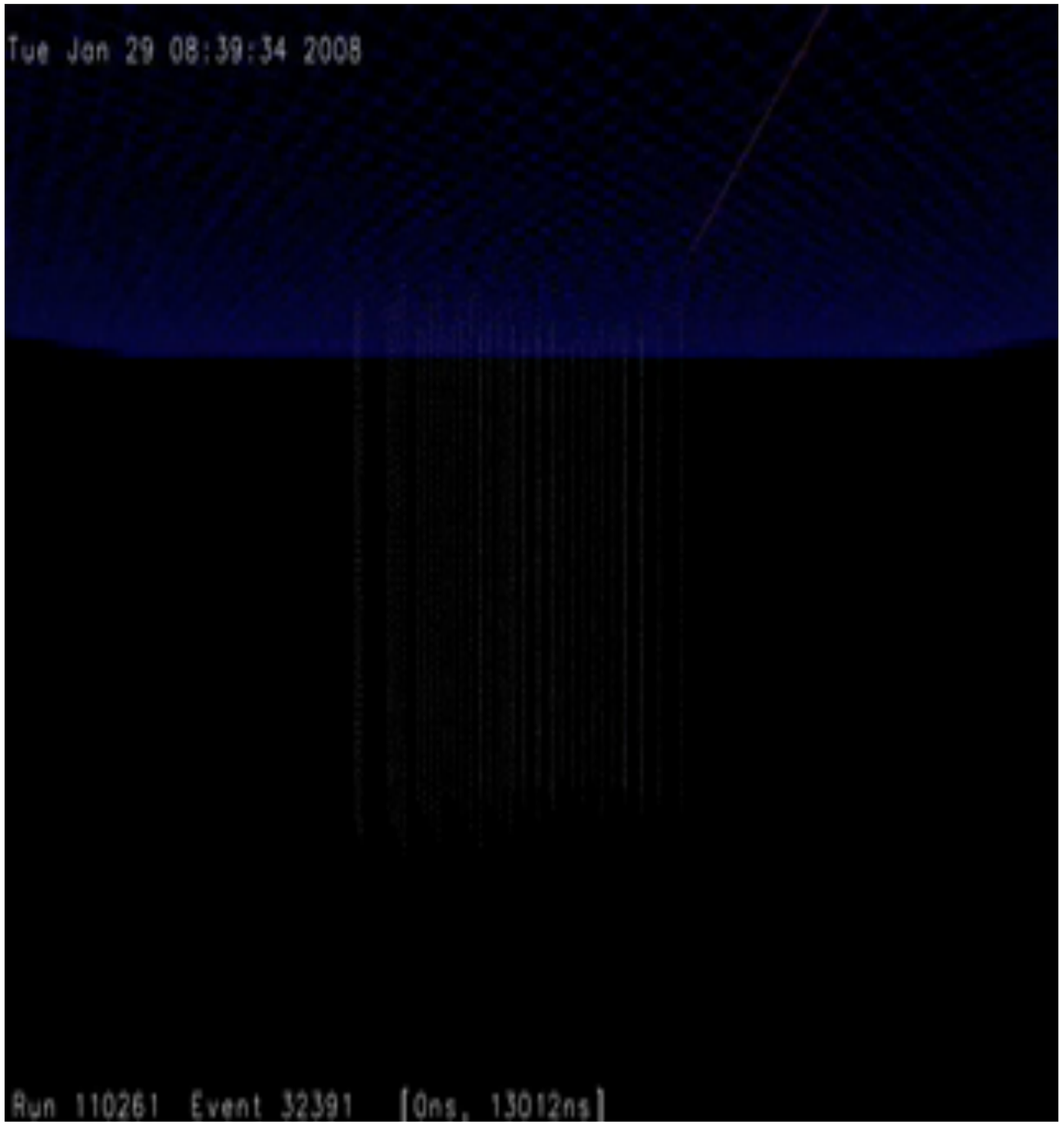
IceCube Depth: 1.5-2.5 km



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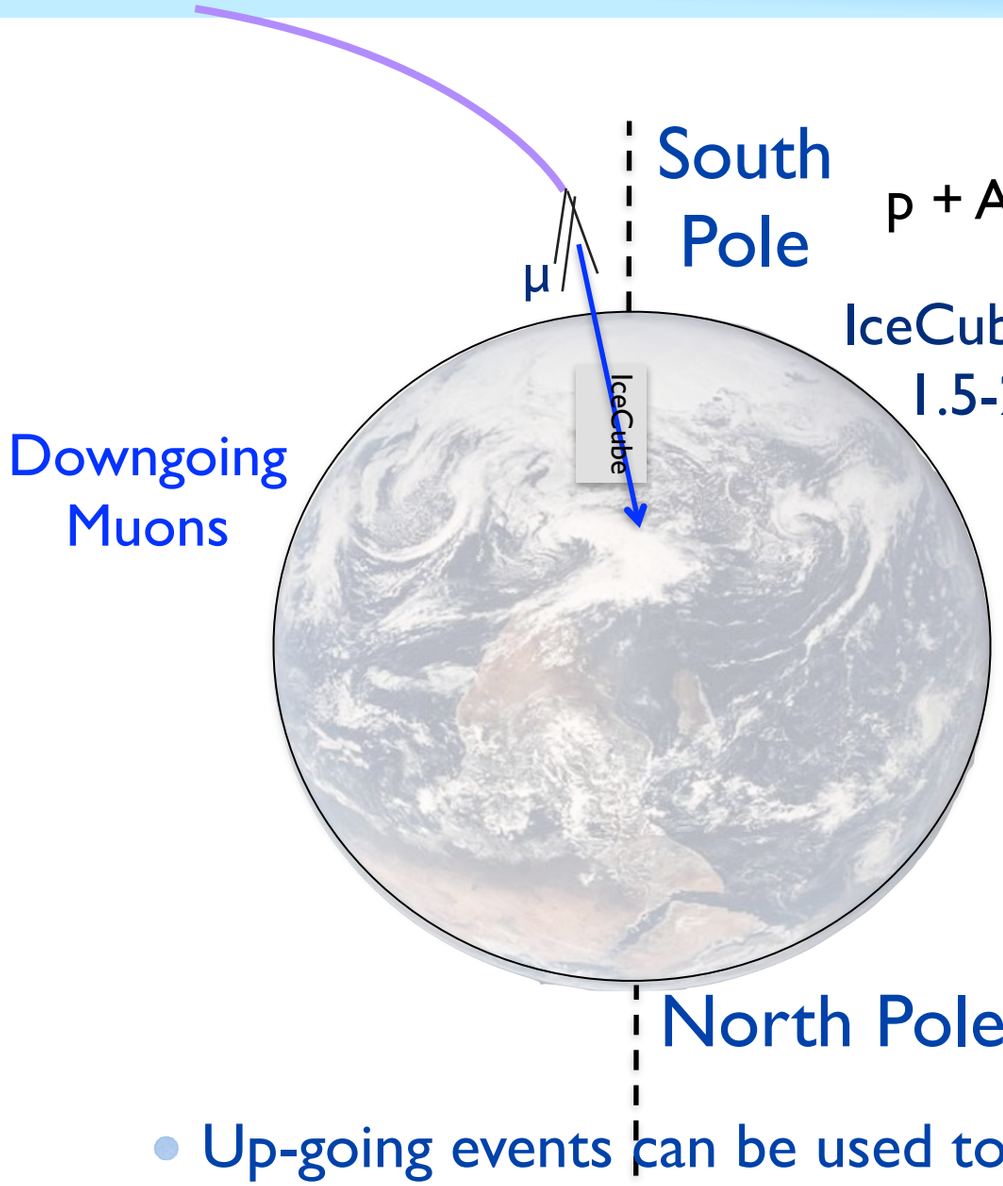

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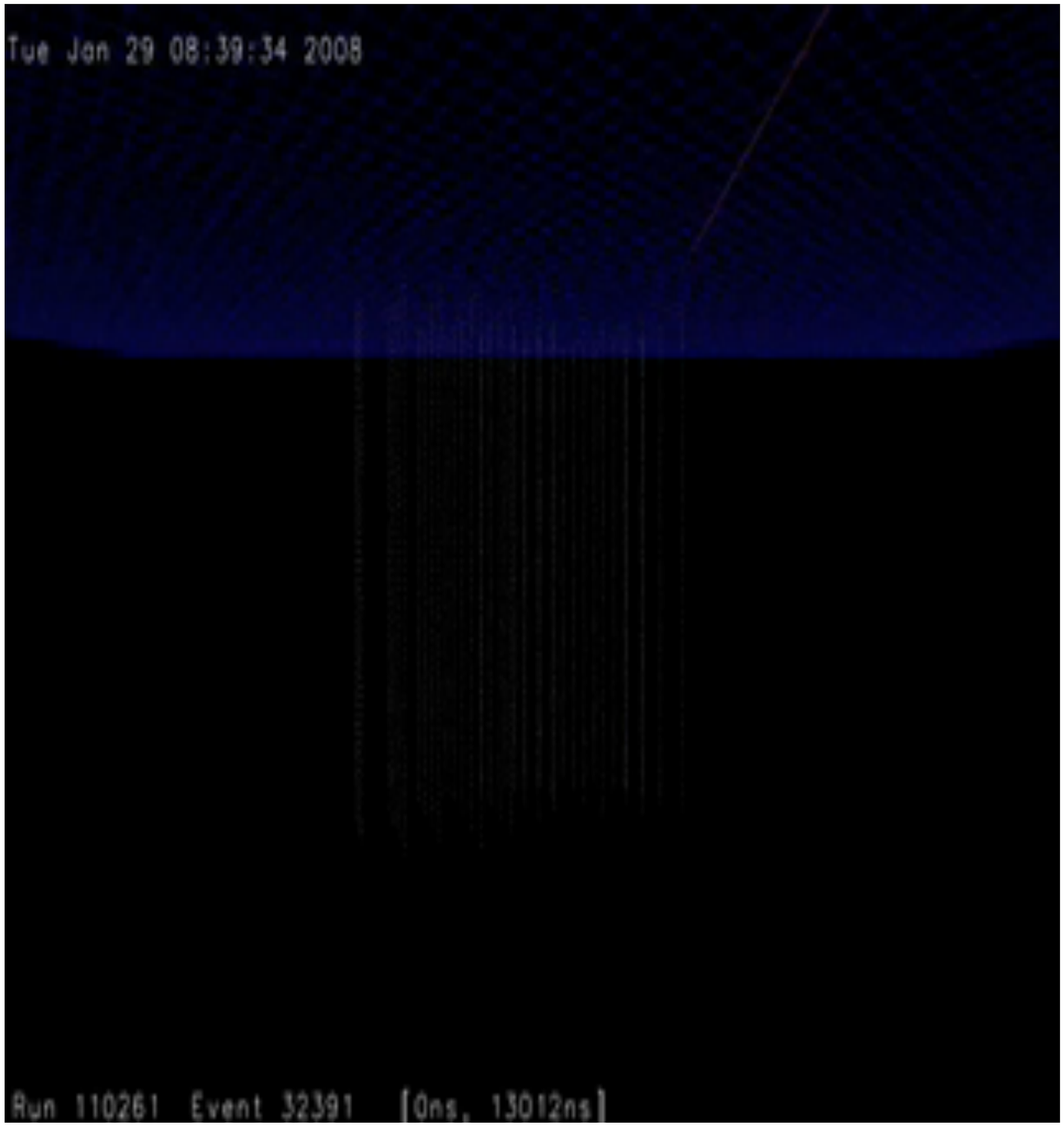

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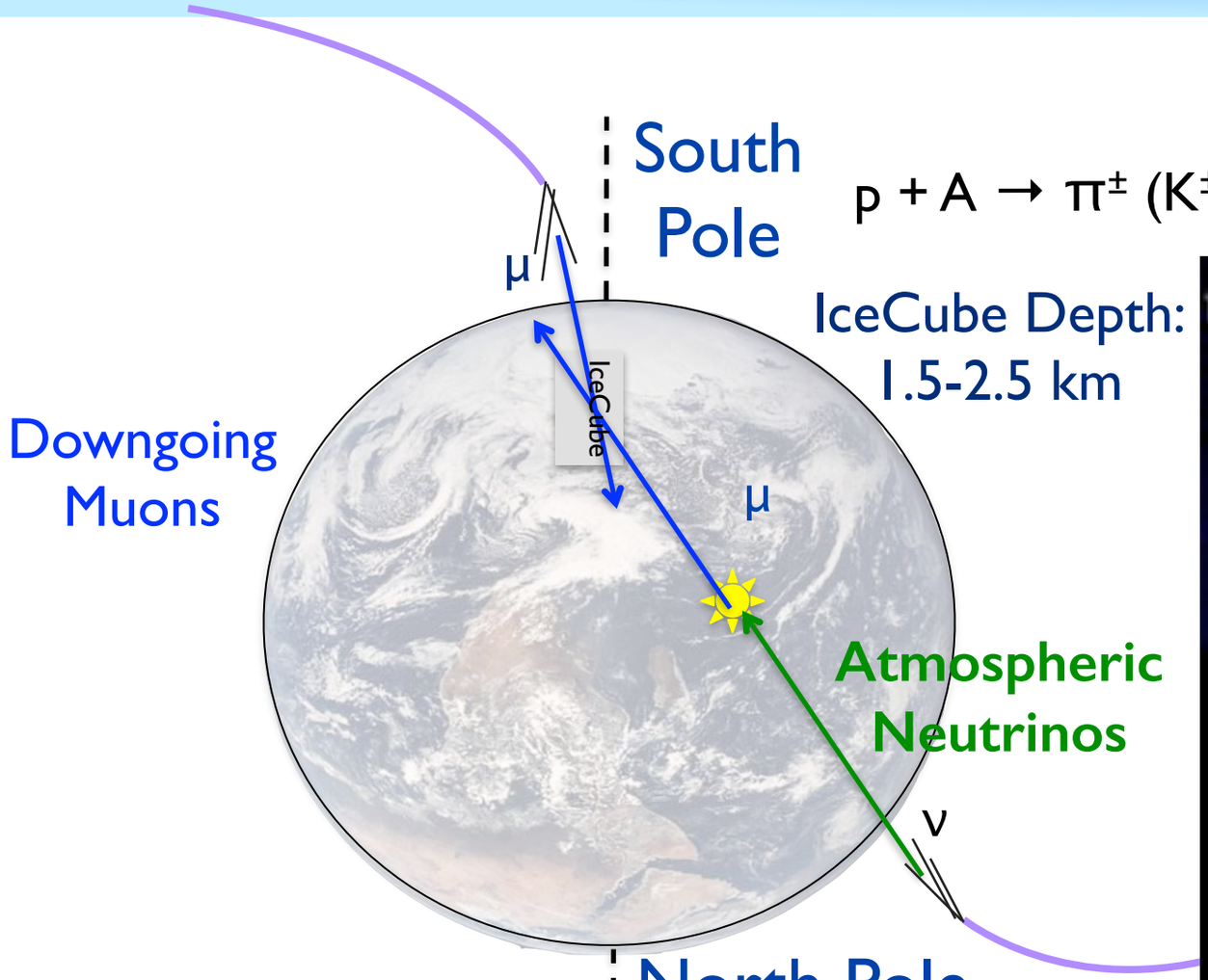
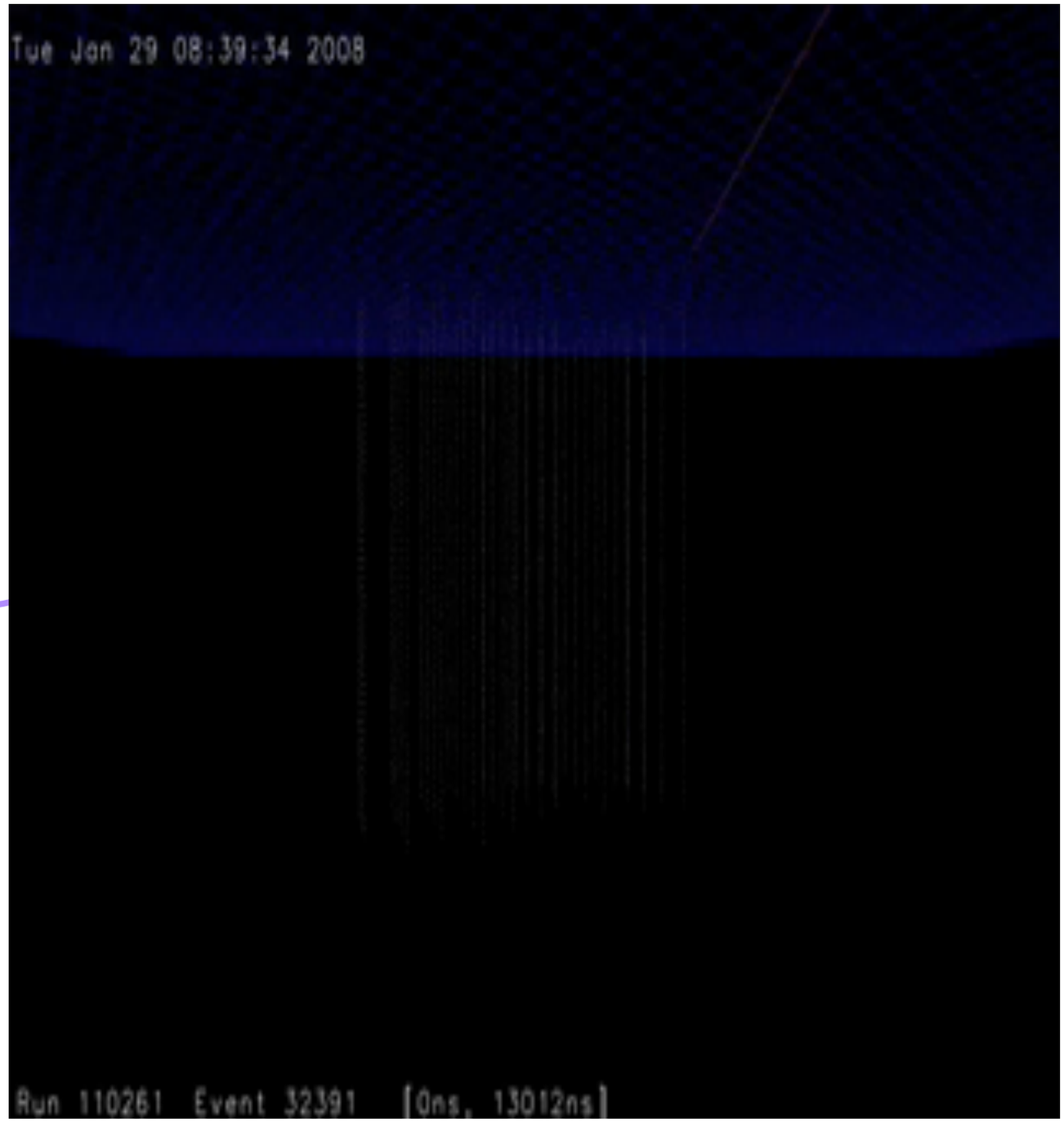


$p + A \rightarrow \pi^{\pm}\,(K^{\pm}) + \text{other hadrons} \;\ldots\; \pi^{+} \rightarrow \mu^{+}\nu_{\mu} \rightarrow e^{+}\nu_{e}\nu_{\mu}\nu_{\mu}$

IceCube Depth: 1.5-2.5 km



- Up-going events can be used to obtain "clean" neutrino sample
  - Earth is used as muon filter
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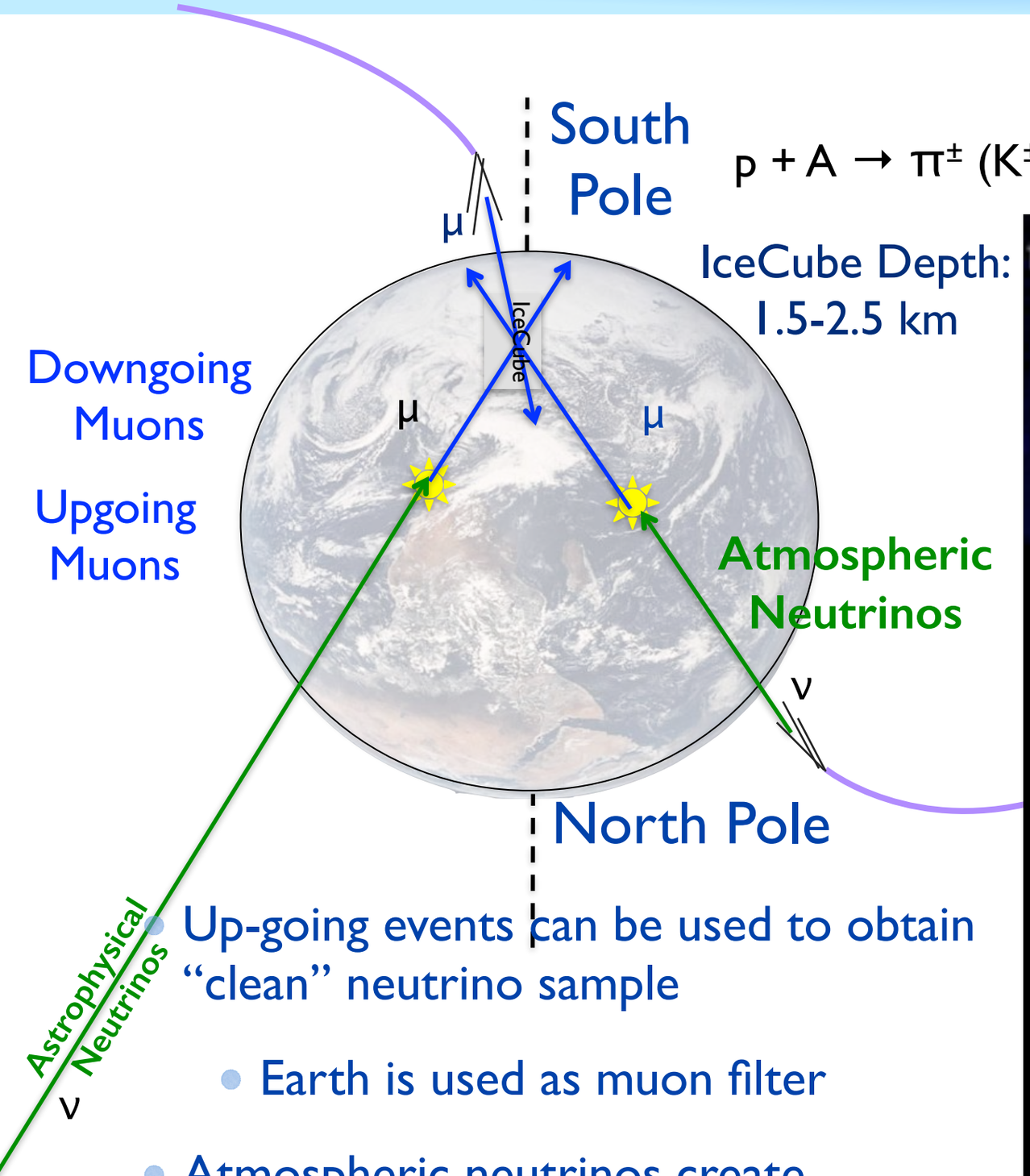
Signals in IceCube



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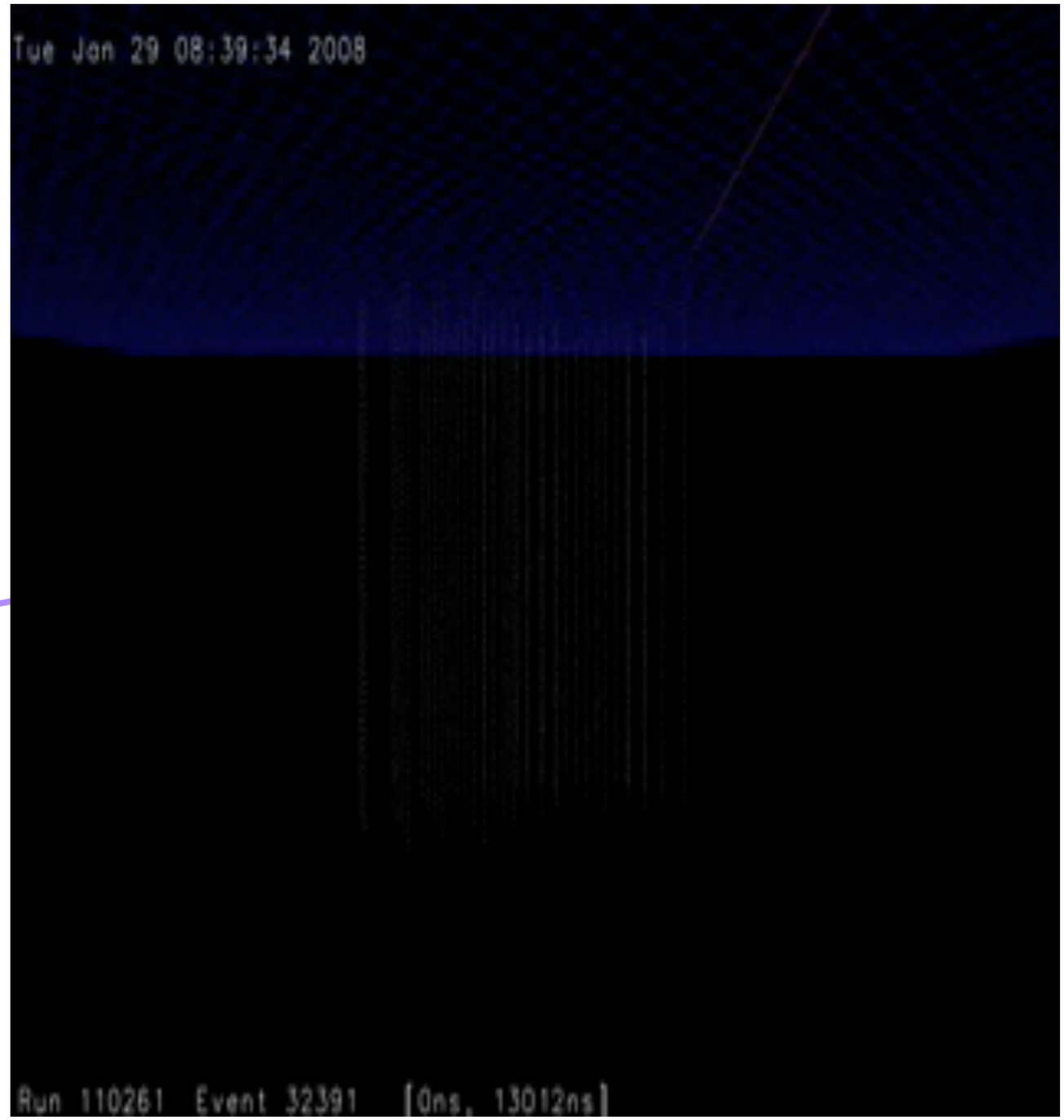

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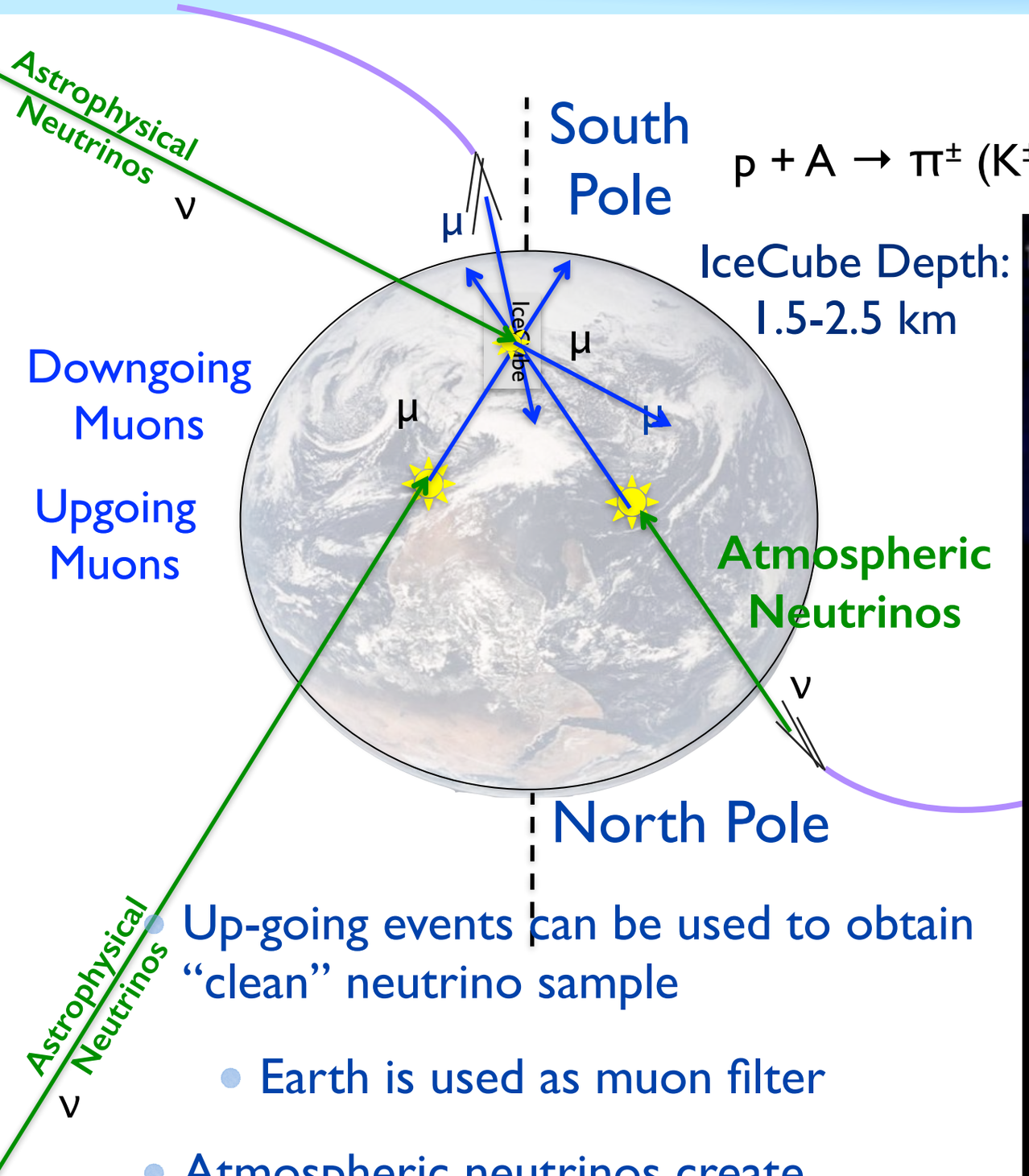

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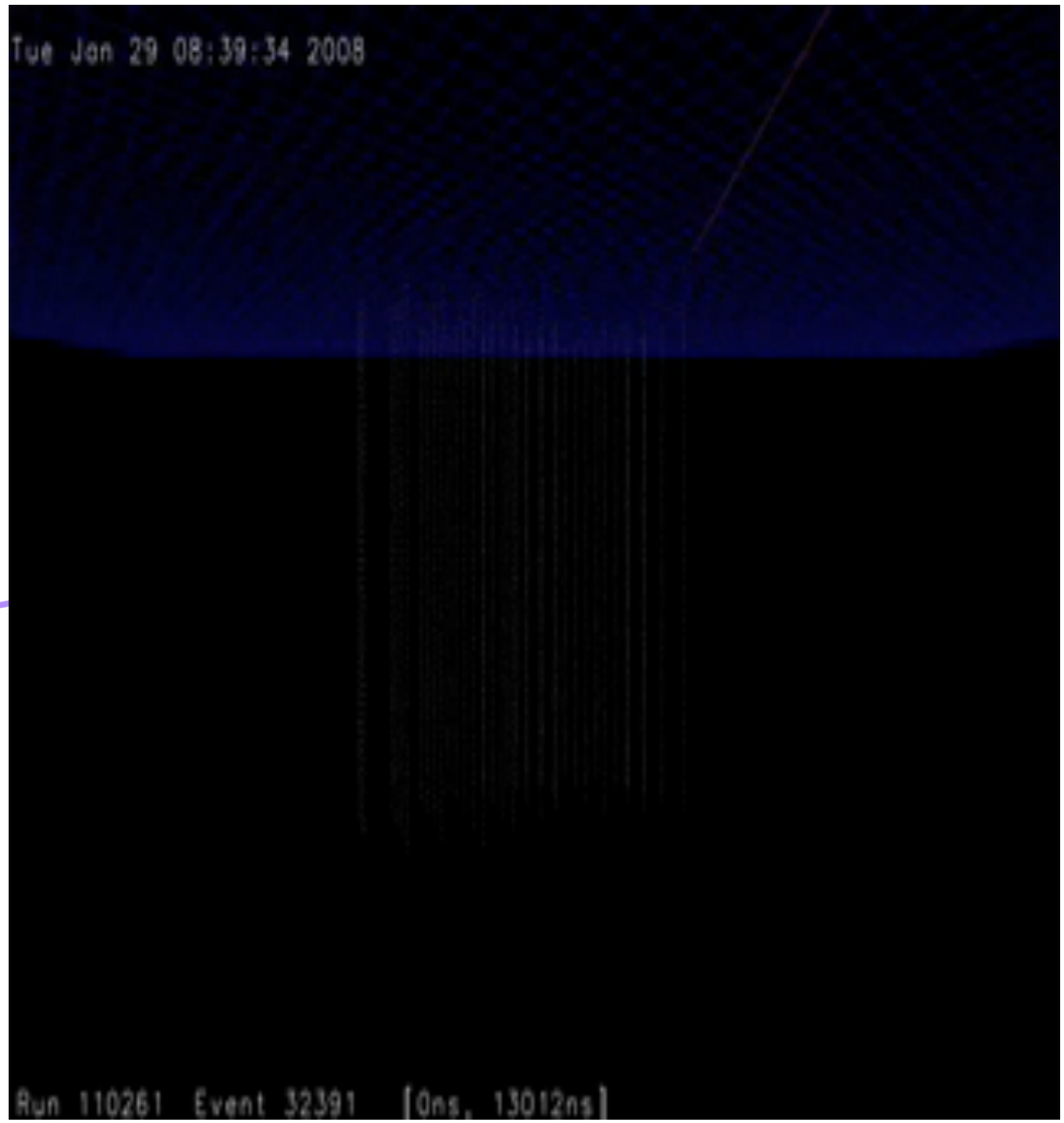
IceCube Depth: 1.5-2.5 km



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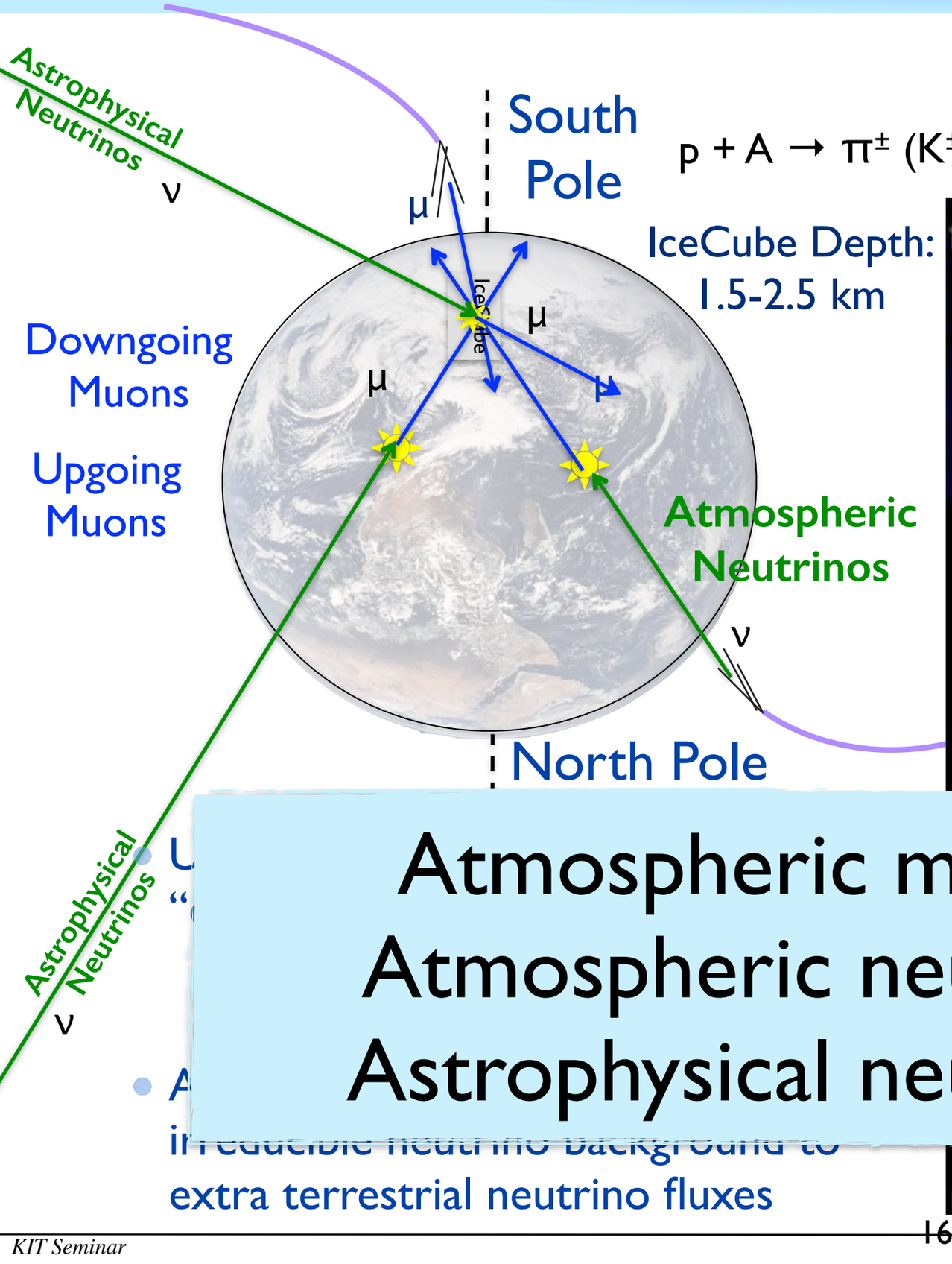
Signals in IceCube



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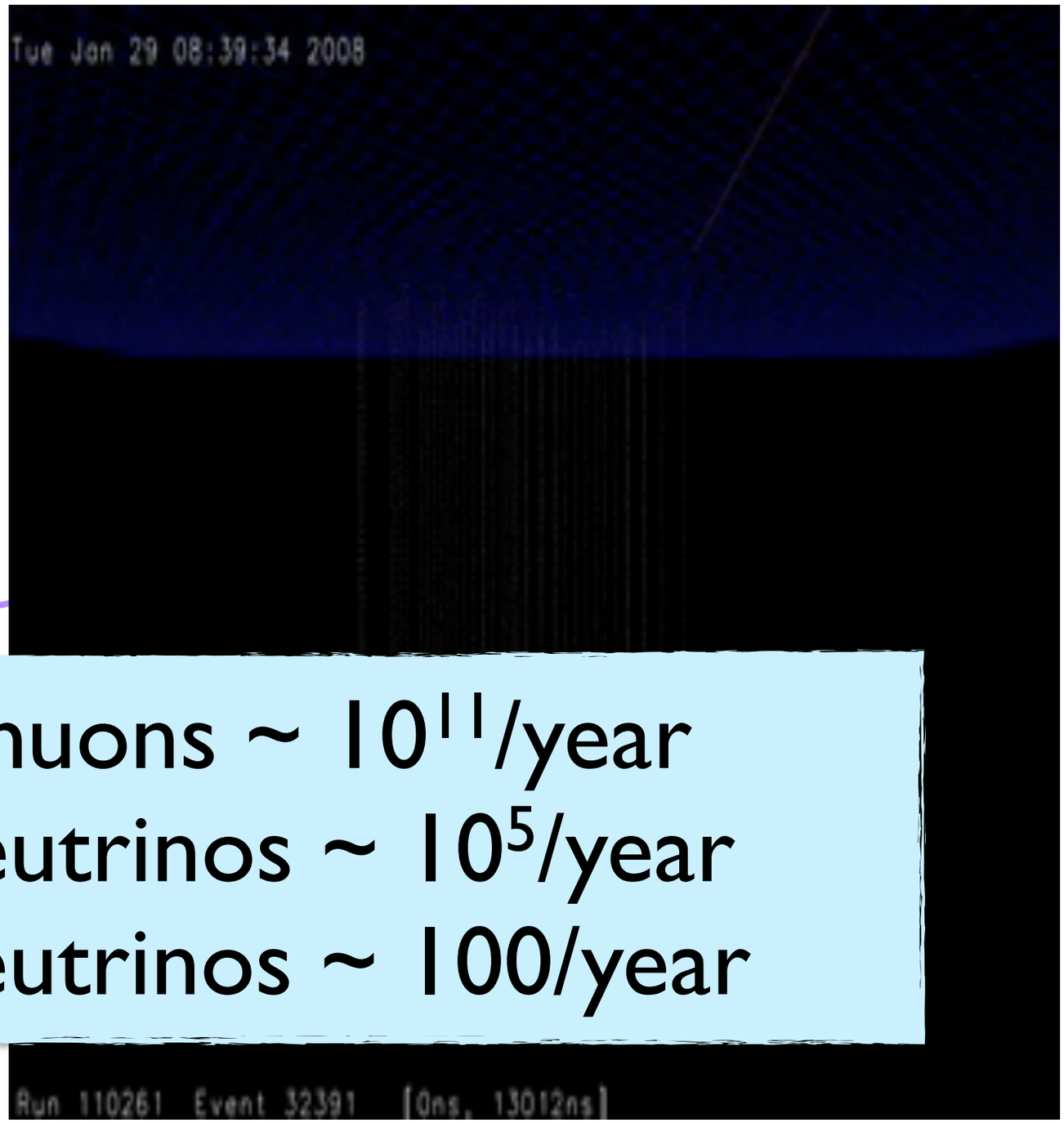
IceCube Depth: 1.5-2.5 km



- U
  "
- A
  irreducible neutrino background to extra terrestrial neutrino fluxes

Atmospheric muons ~ 10^{11}/year
Atmospheric neutrinos ~ 10^{5}/year
Astrophysical neutrinos ~ 100/year

Event Topologies in IceCube

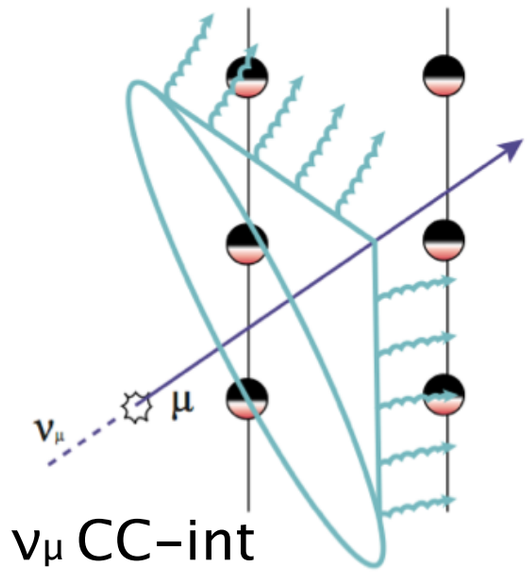
Track topology

(e.g. induced by muon neutrino)

CC: ν_μ

Good pointing,
0.2° - 1°

Lower bound on energy for through-going events



ν_μ CC–int

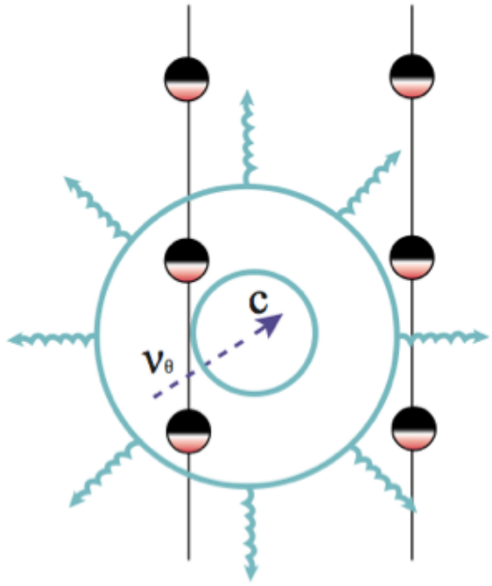
Cascade topology

(e.g. induced by electron neutrino)

CC: ν_e ν_τ
NC: ν_e ν_μ ν_τ

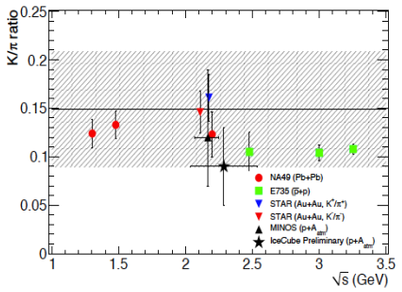
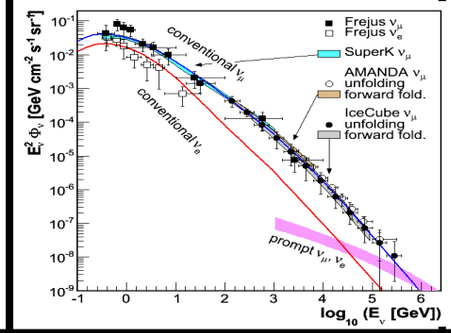
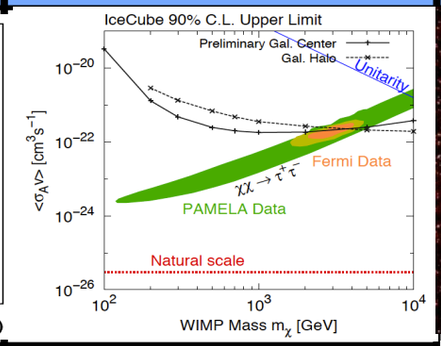
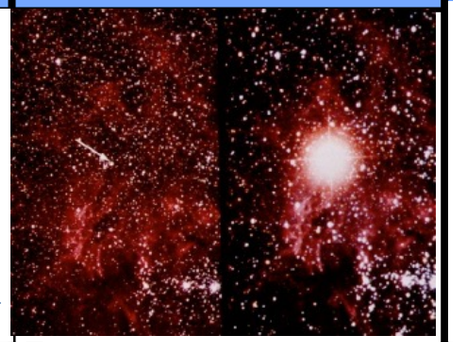
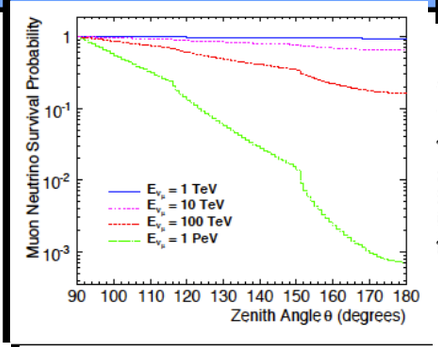
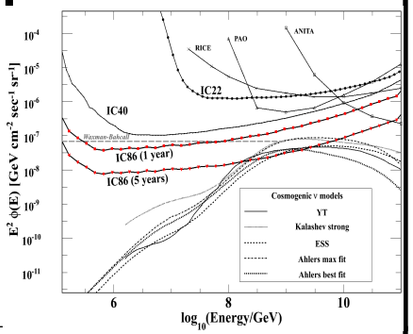
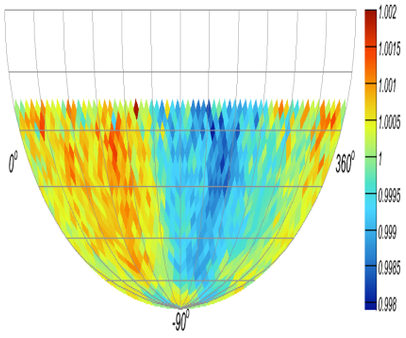
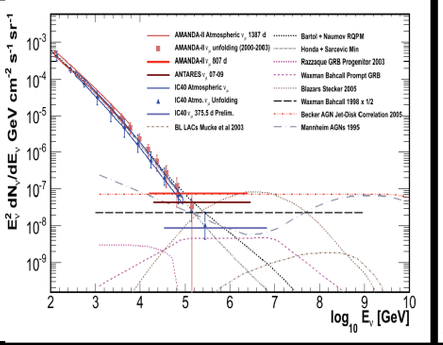
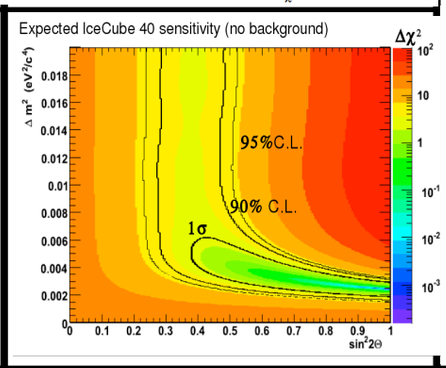
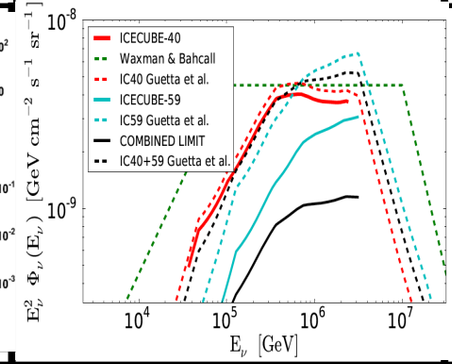
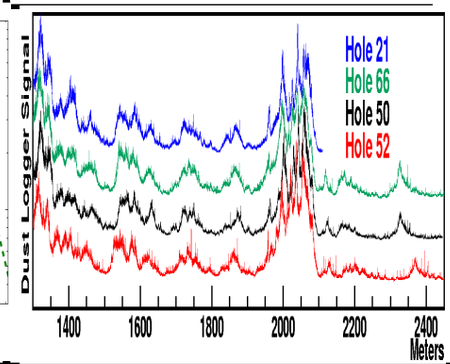
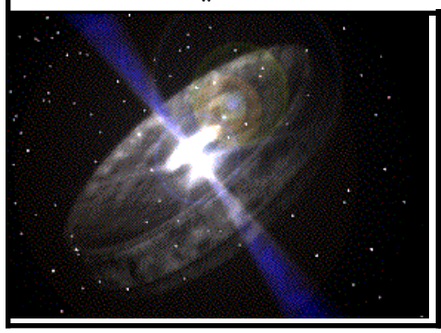
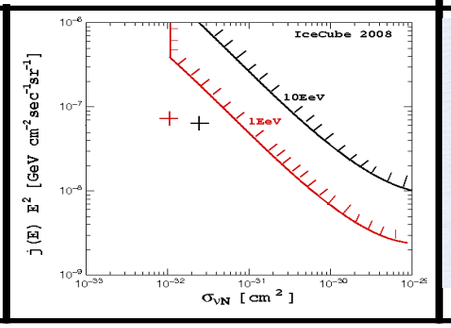
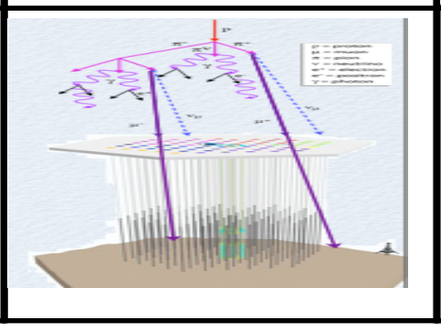
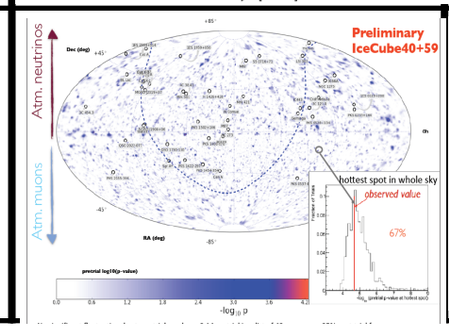
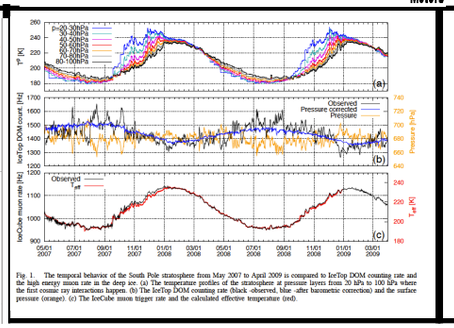
Good energy resolution, 15%
Some pointing,
10° - 15°

time delay
vs. direct light
"on time" → delayed



ν_e ν_τ CC–int & ν_i NC–int

IceCube Science

Cosmic Rays	Atmospheric neutrinos	Particle Physics	Astronomy	Applied science	Cosmology
					
					
					

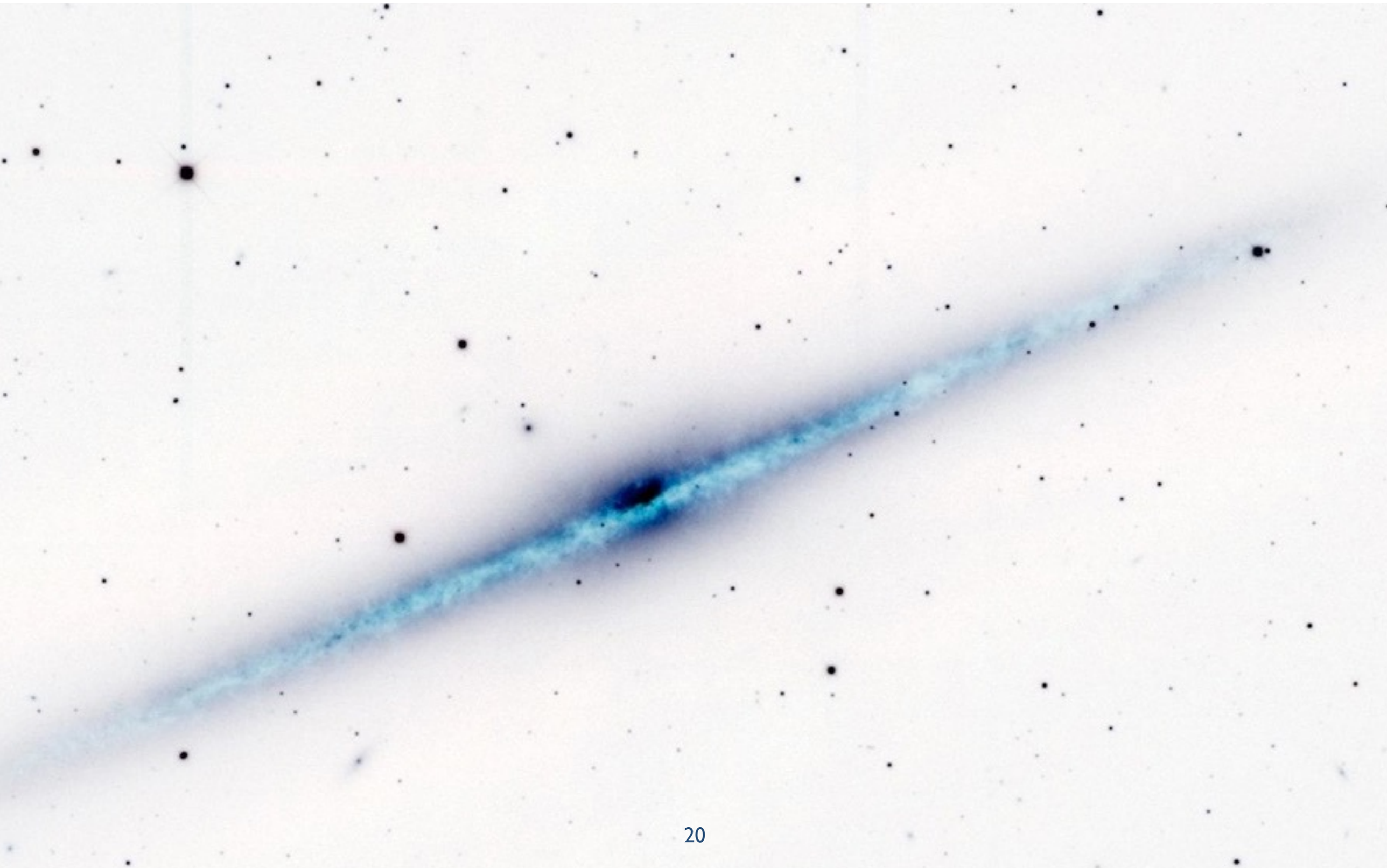
Very diverse science program, with neutrinos from 10GeV to EeV, and MeV burst neutrinos

Dark Matter Self-annihilations

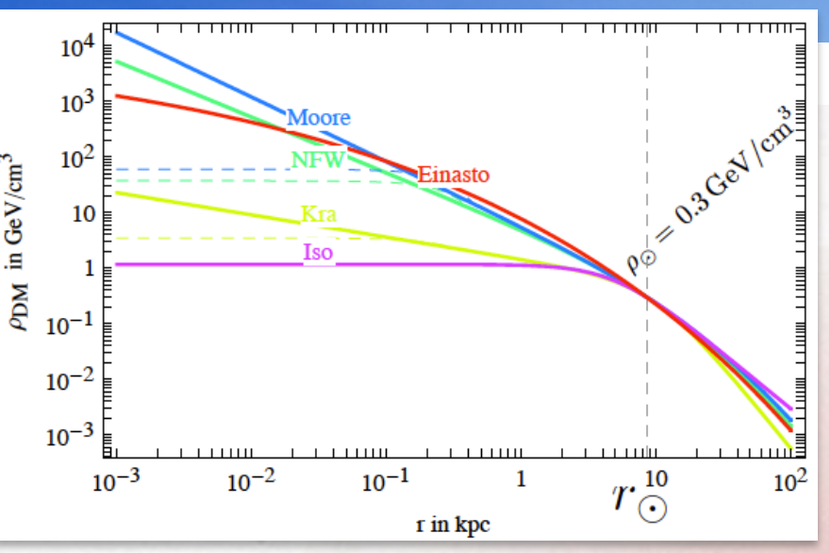
$\langle \sigma_A v \rangle$

Dark Matter in the Milky Way

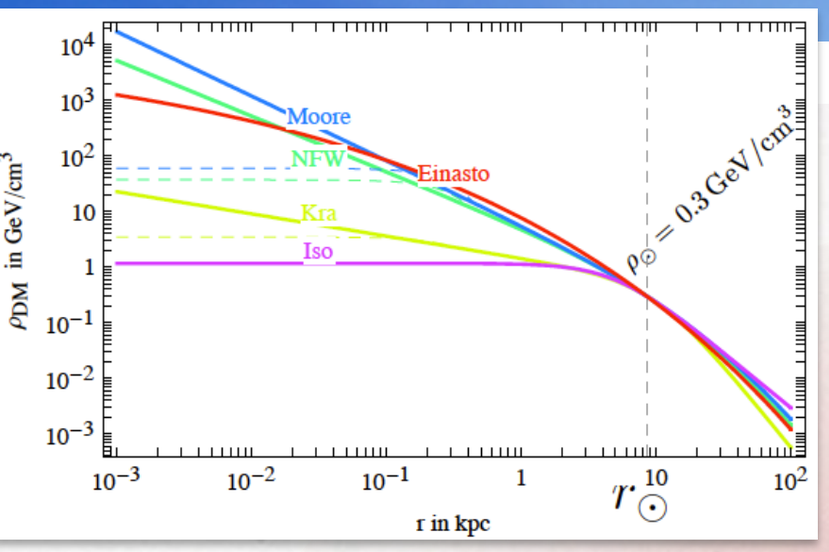
Dark Matter in the Milky Way



Dark Matter in the Milky Way

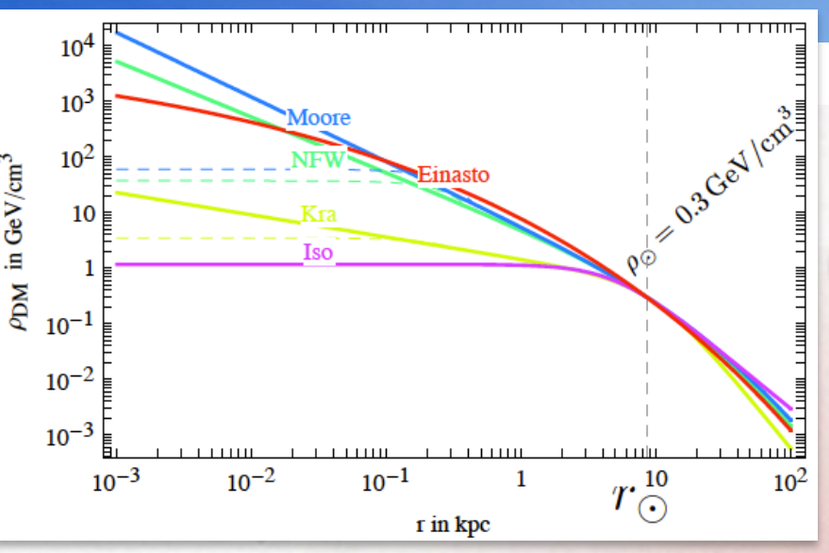


Dark Matter in the Milky Way

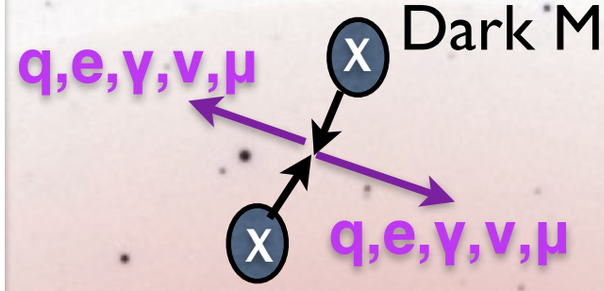


~8kpc

Dark Matter in the Milky Way

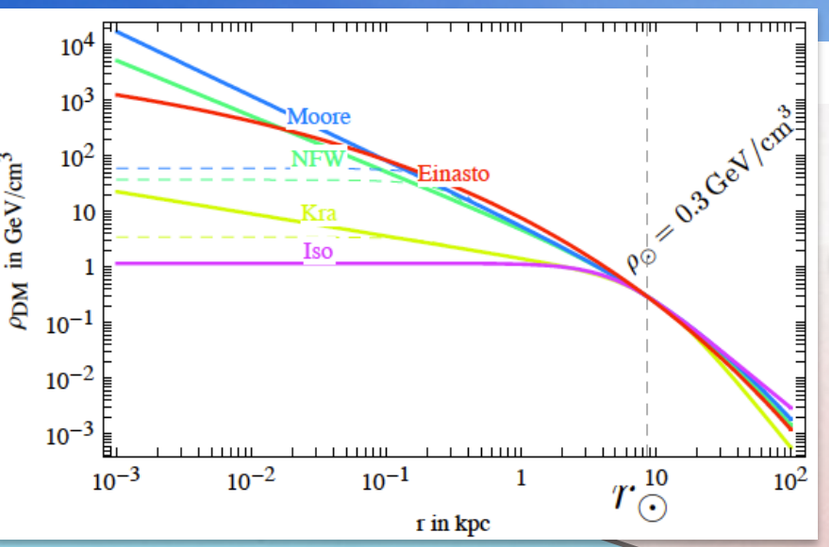


Dark Matter self-annihilation ~ρ^2



~8kpc

Dark Matter in the Milky Way



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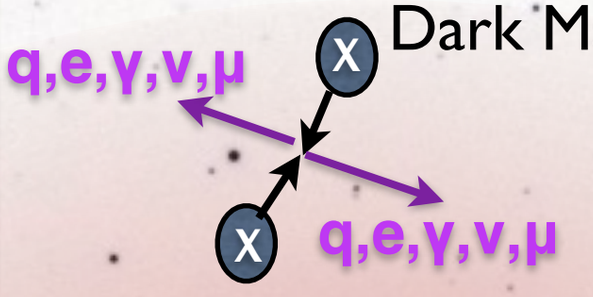
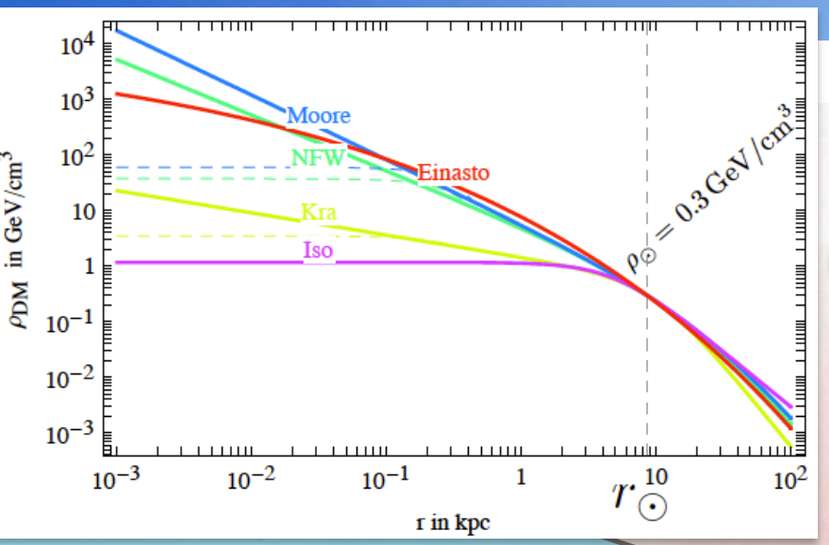


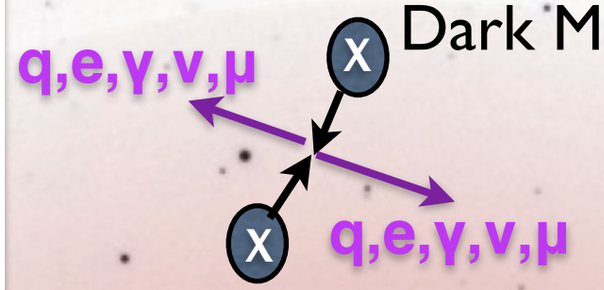
~8kpc

Dark Matter in the Milky Way



1

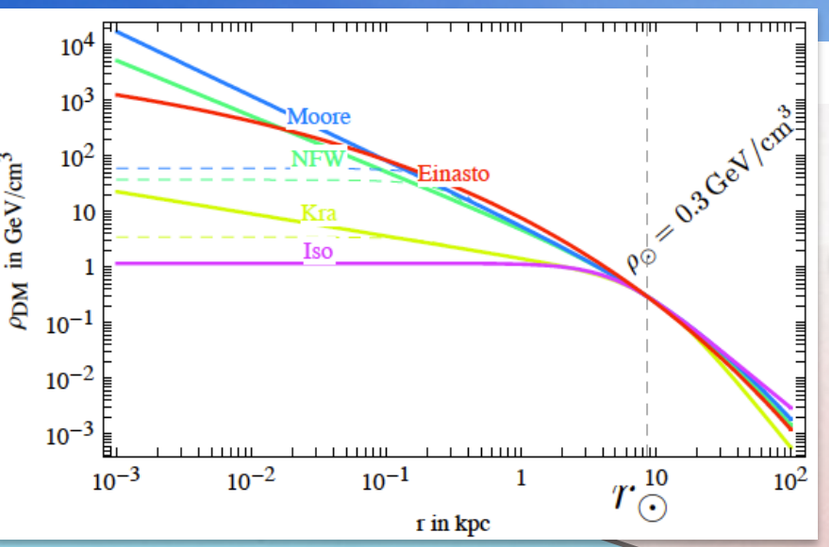
Dark Matter self-annihilation ~ρ^2



~8kpc

1

Dark Matter in the Milky Way



Dark Matter self-annihilation ~ρ^2

q,e,γ,ν,μ

q,e,γ,ν,μ

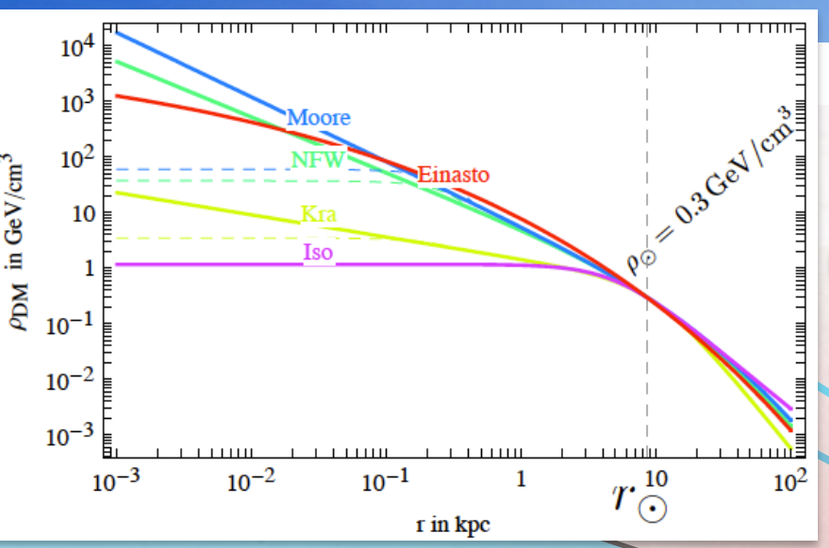
~8kpc

1

1

2

Dark Matter in the Milky Way



1

3

Dark Matter self-annihilation ~ρ^2

X X

q,e,γ,ν,μ

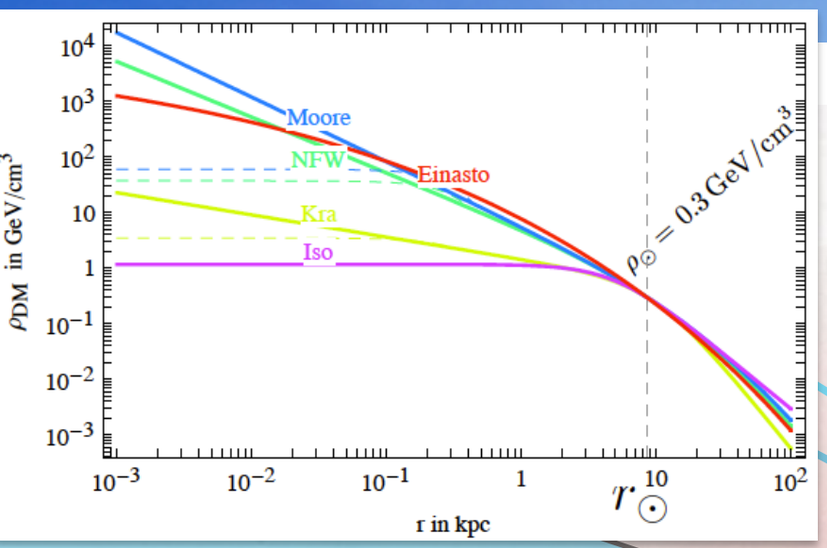
q,e,γ,ν,μ

~8kpc

1

2

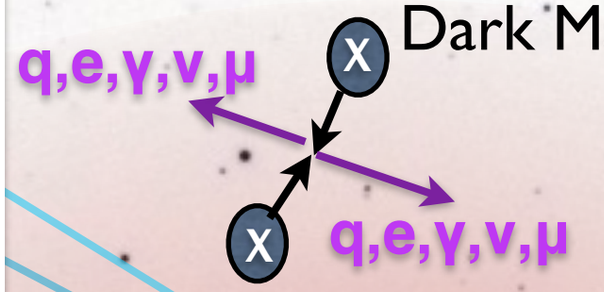
Dark Matter in the Milky Way



1

3

Dark Matter self-annihilation ~ρ^2



1

~8kpc

2

<u>Three targets:</u>
1) Search for a neutrino anisotropy (outer halo)
2) Galactic Center (down-going events)
3) Dwarf Spheriodals, Galaxy Clusters

Analyses follow theoretical discussions in Beacom et al., Phys. Rev. Lett. 99, 231301 (2007) and Yuksel et al., Phys. Rev. D 76, 123506 (2007)

Dark Matter Annihilation

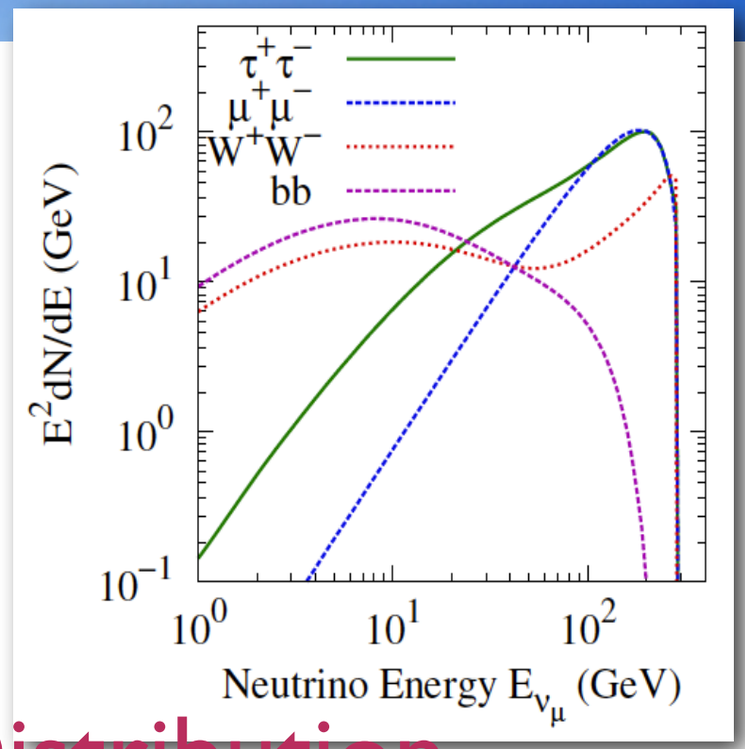
Measure Flux

$$\frac{d\Phi}{dE}(E,\phi,\theta)$$

=

Particle Physics

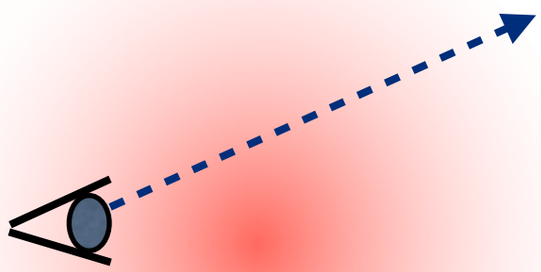
$$\frac{1}{4\pi}\frac{\langle\sigma_A v\rangle}{2m_\chi^2}\Sigma_f\frac{dN}{dE}B_f$$



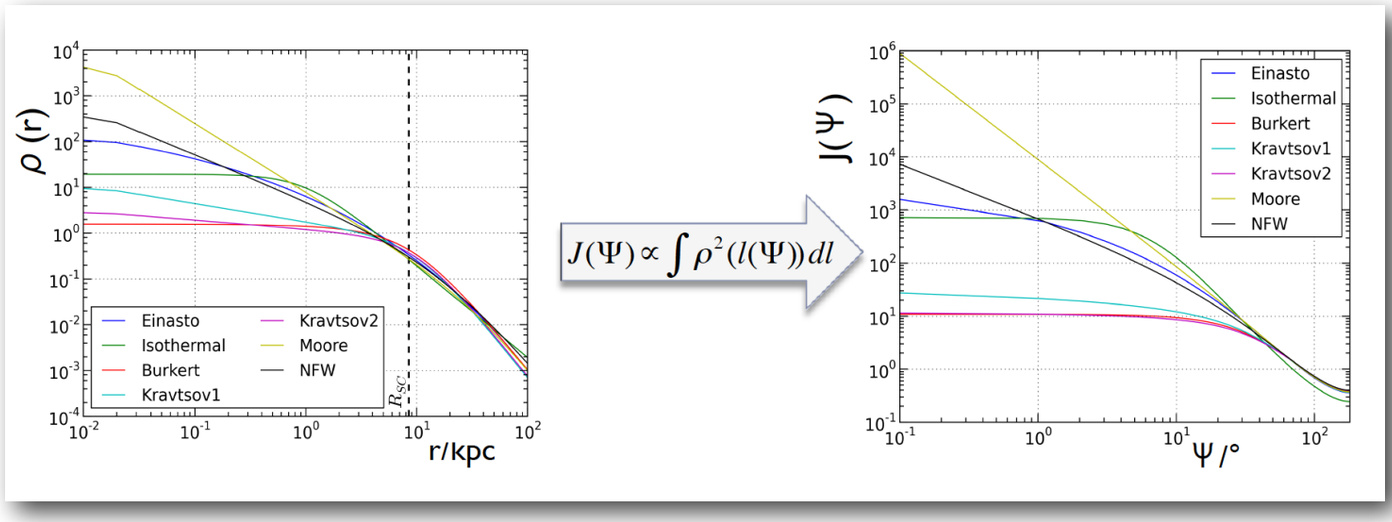
x

Dark Matter Distribution

line of sight (los) integral



$$\int_{\Delta\Omega(\phi,\theta)} d\Omega' \int_{\text{los}} \rho^2(r(l,\phi'))dl(r,\phi')$$



Dark Matter Annihilation

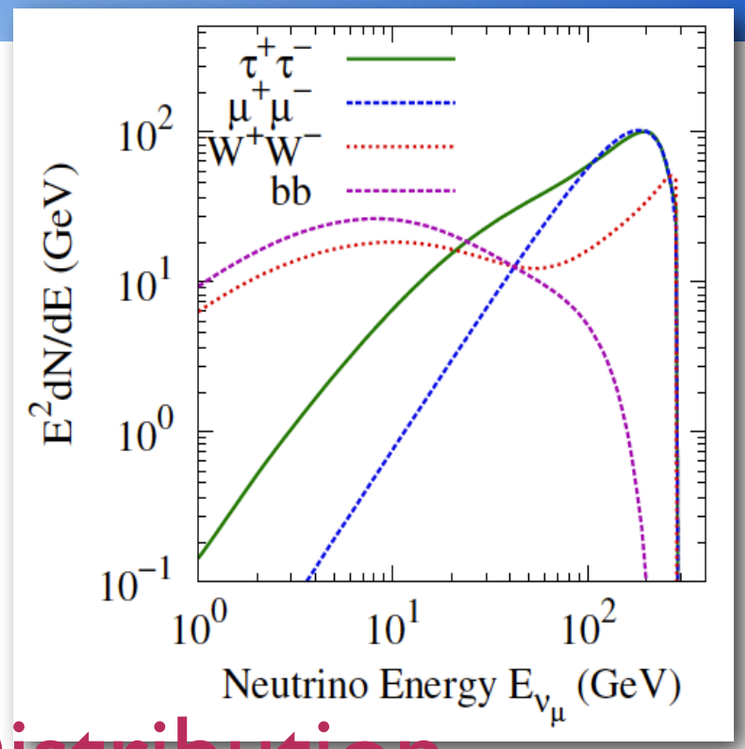
Measure Flux

$$\frac{d\Phi}{dE}(E,\phi,\theta)$$

=

Particle Physics

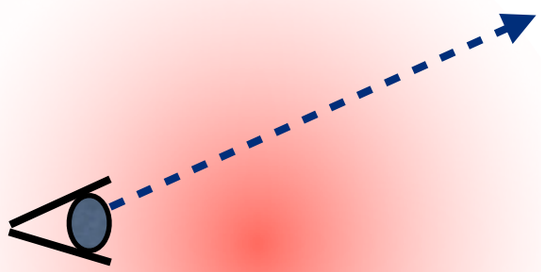
$$\frac{1}{4\pi}\frac{\langle\sigma_A v\rangle}{2m_\chi^2}\Sigma_f \frac{dN}{dE}B_f$$



x

Dark Matter Distribution

line of sight (los) integral



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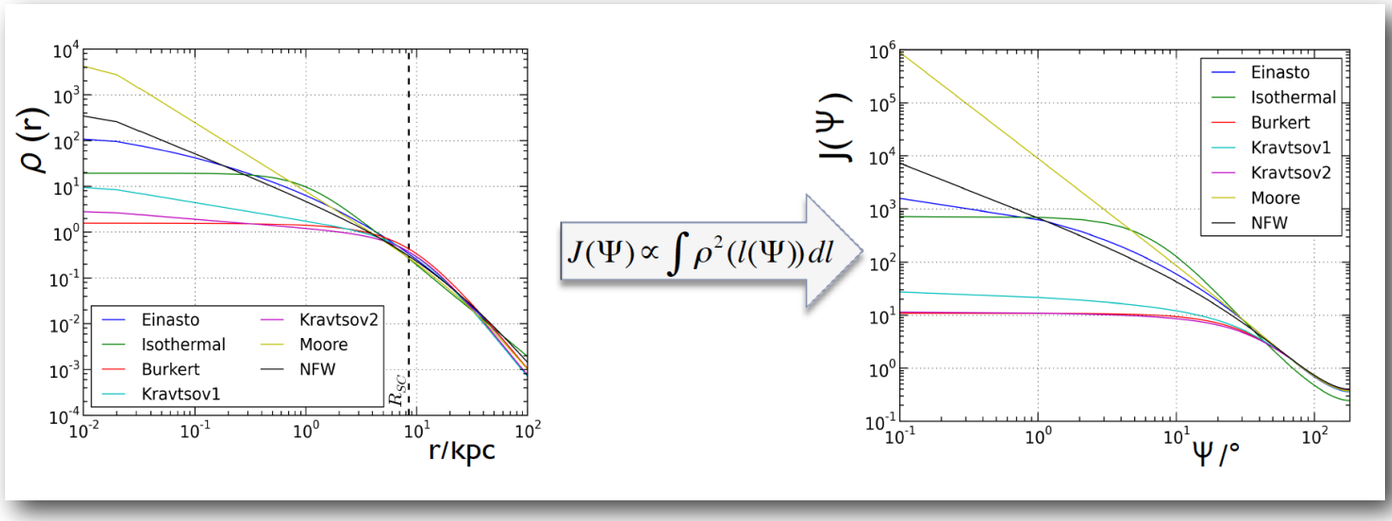
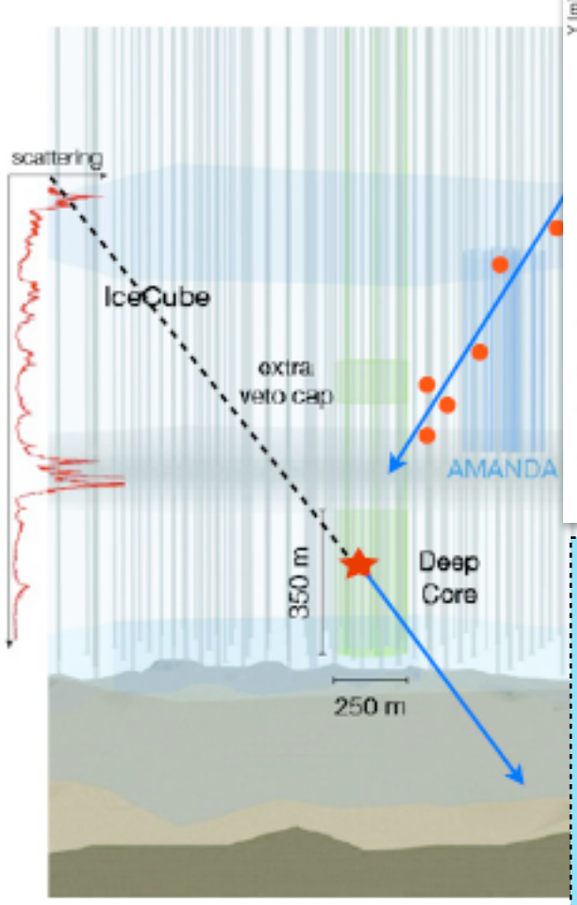
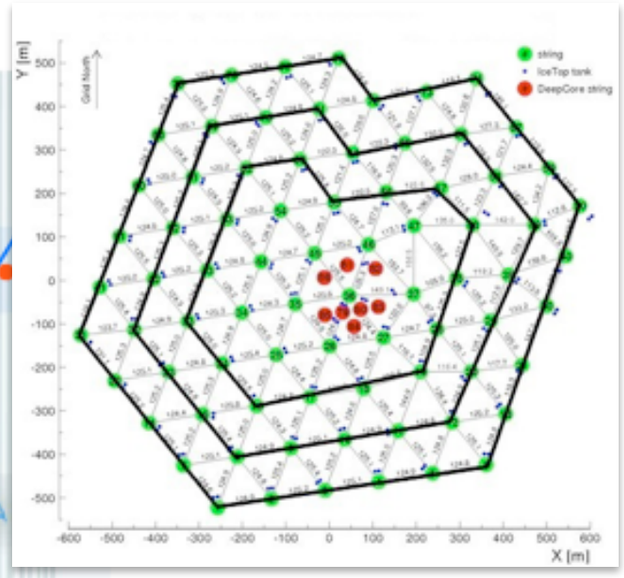


IceCube Collaboration
arXiv1505.07259
Eur.Phys.J. C75 (2015) 10, 492

Galactic Center

Use IceCube external strings as a veto:
- 3 complete layers around DeepCore (~ 375m)
- Full sky sensitivity: access to southern hemisphere





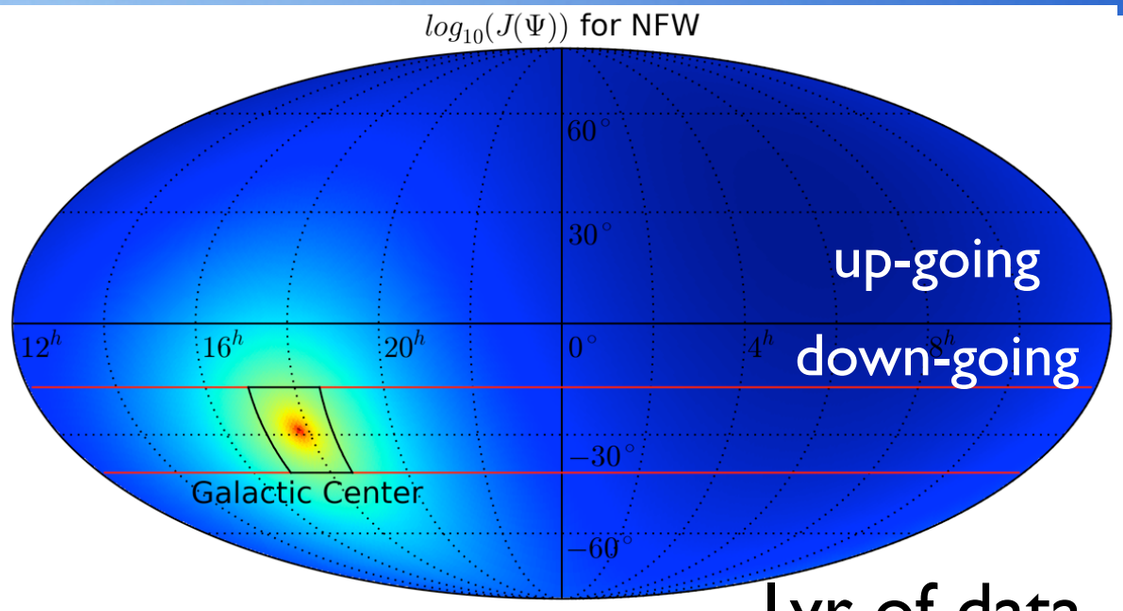
Two analyses, combined based on best sensitivity (dotted).

Both analyses observed underfluctuations

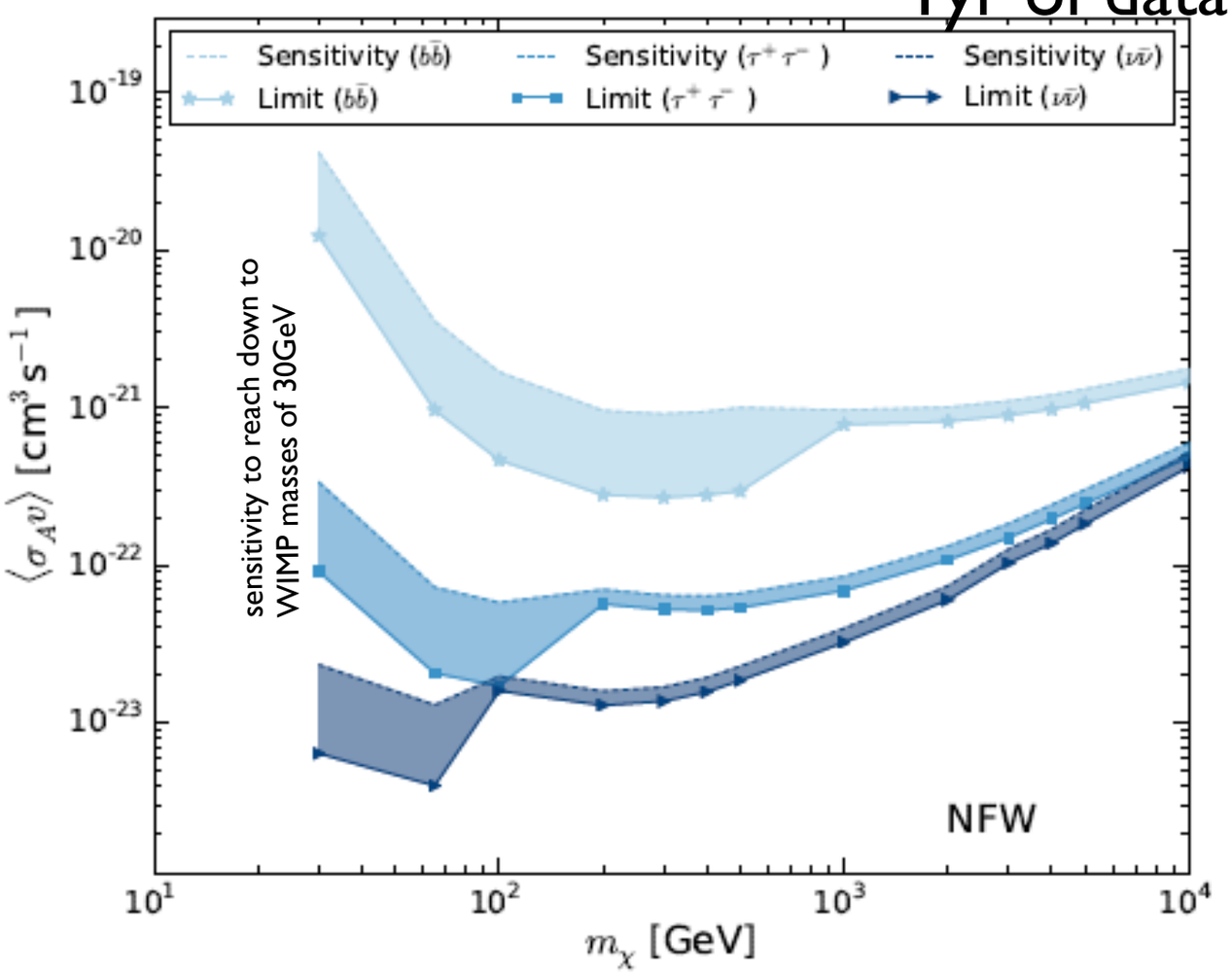
"High energy" analysis: 0.5 σ

"Low energy" analysis: 2σ

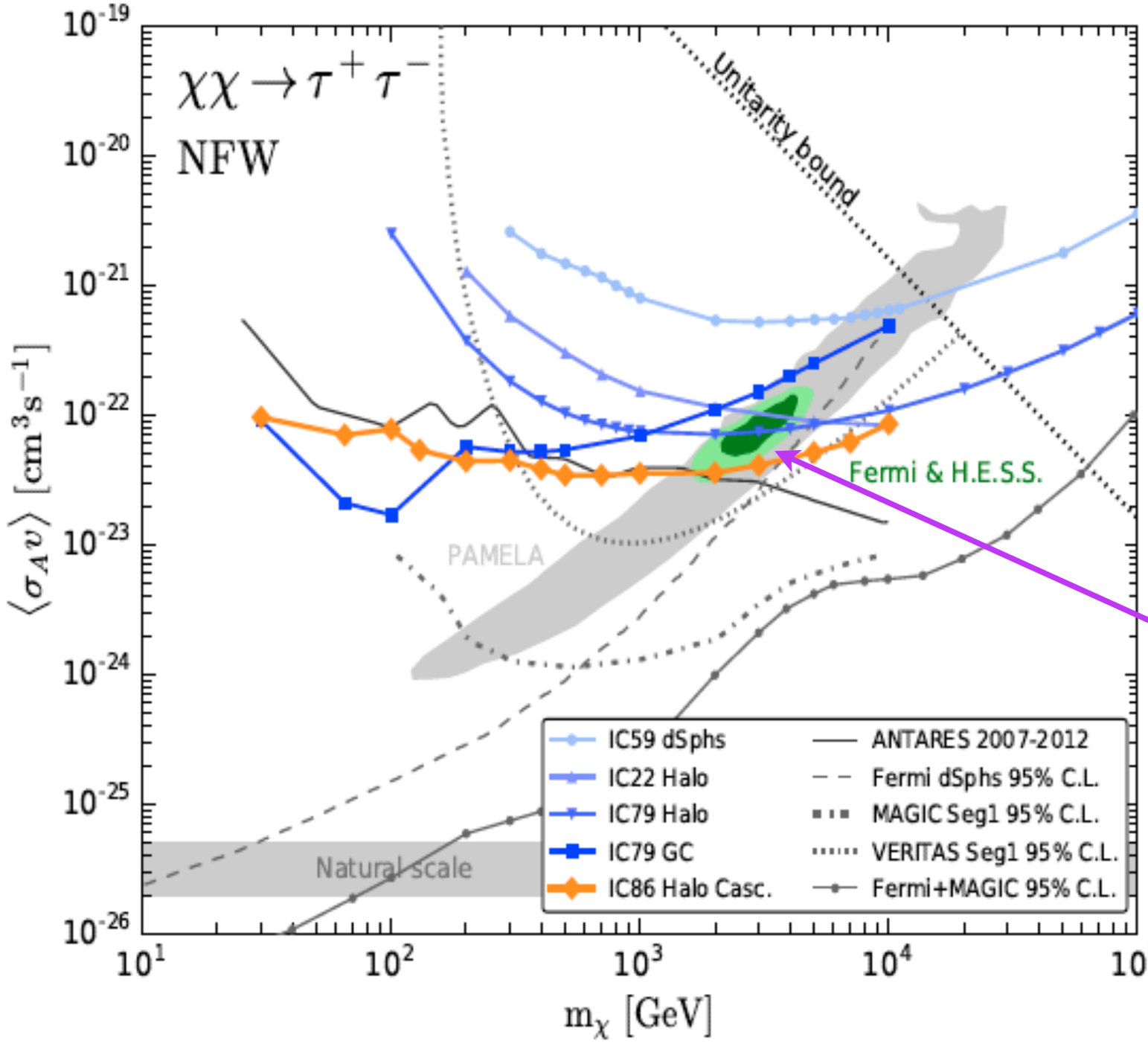
Separate Low energy and High energy optimizations:
GC is above the horizon
→ Fiducial volume in central strings
→ refined muon veto from surrounding layers
Use scrambled data for background estimation



1yr of data



Neutrinos test lepton anomalies



$e^+/(e^+ + e^-)$

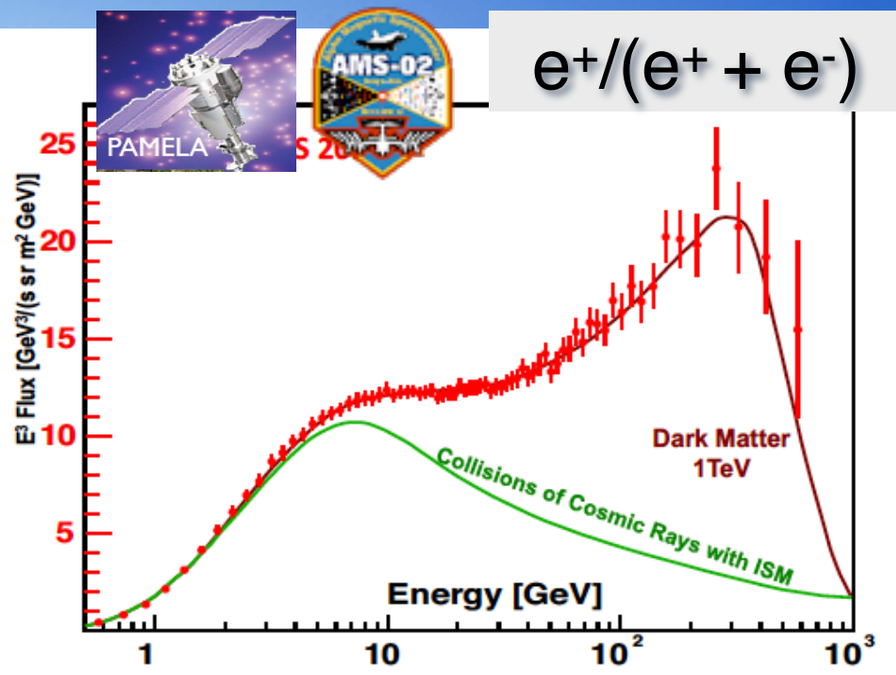
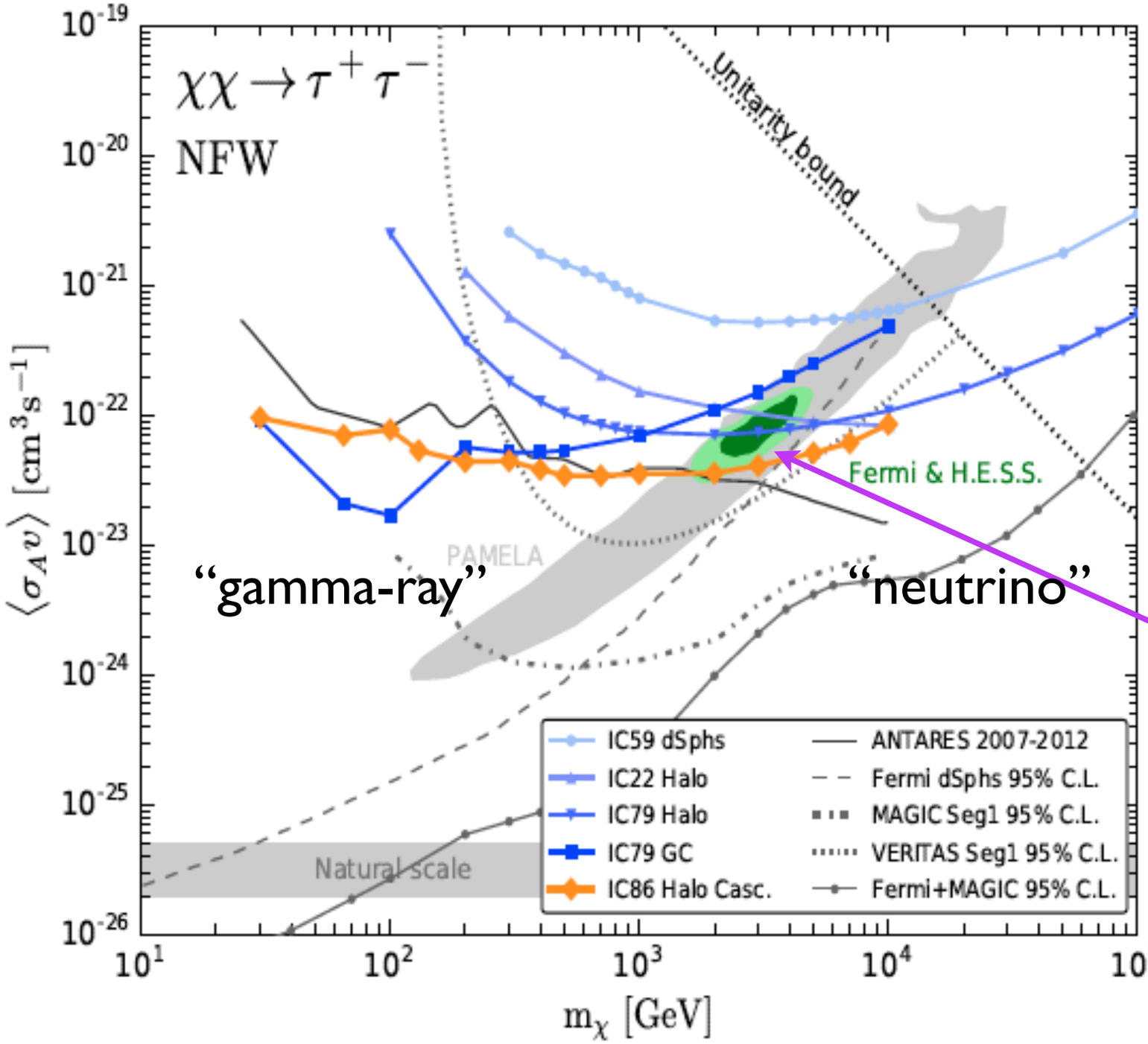


Figure 4. The current AMS positron flux measurement compared with theoretical models.

$(e^+ + e^-)$

Neutrino Telescopes can probe models motivated by the observed lepton anomalies

Neutrinos test lepton anomalies



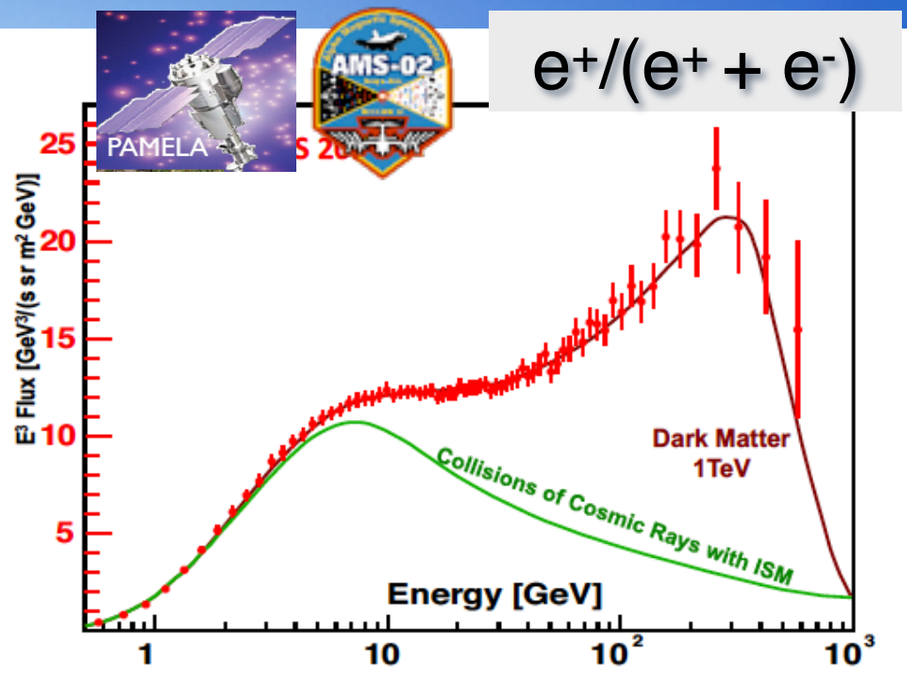
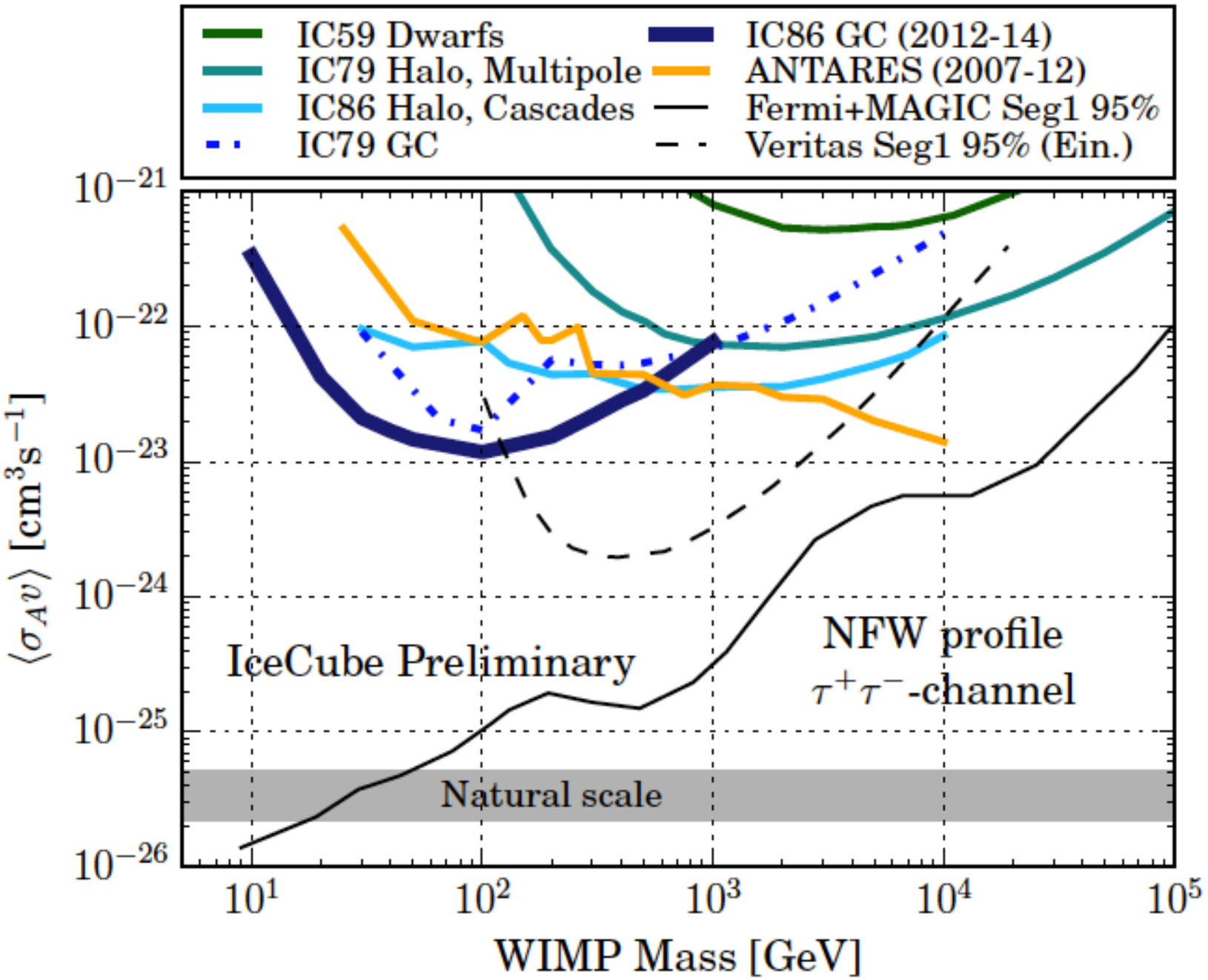


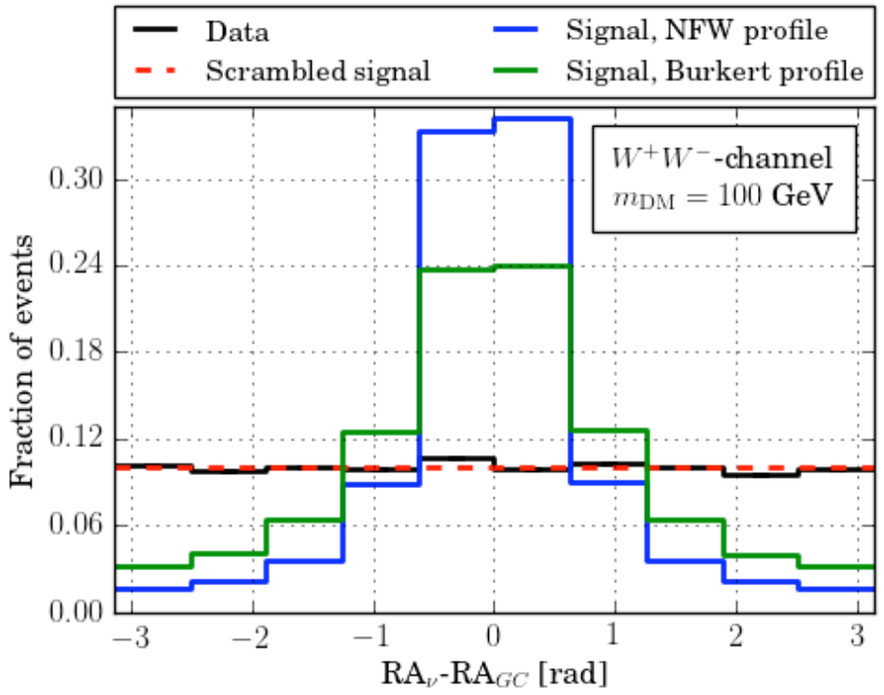
Figure 4. The current AMS positron flux measurement compared with theoretical models.

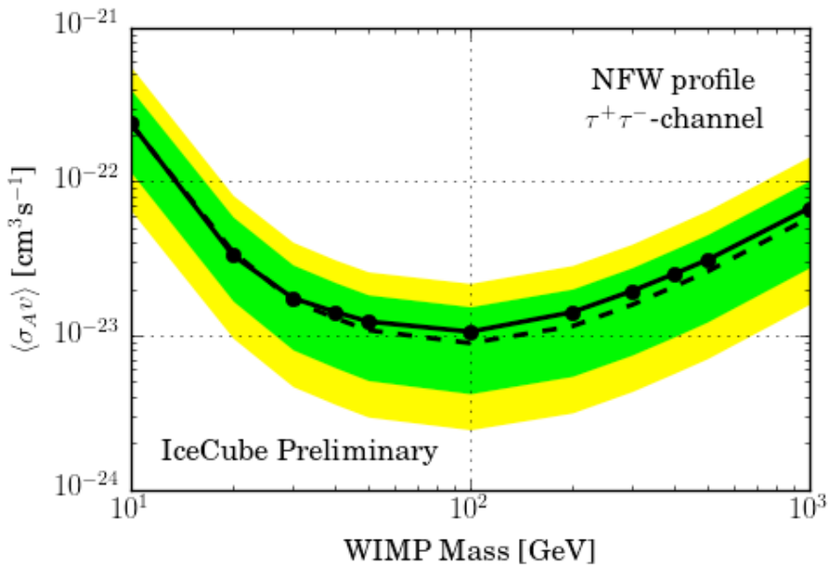
Neutrino Telescopes can probe models motivated by the observed lepton anomalies

DeepCore Analysis



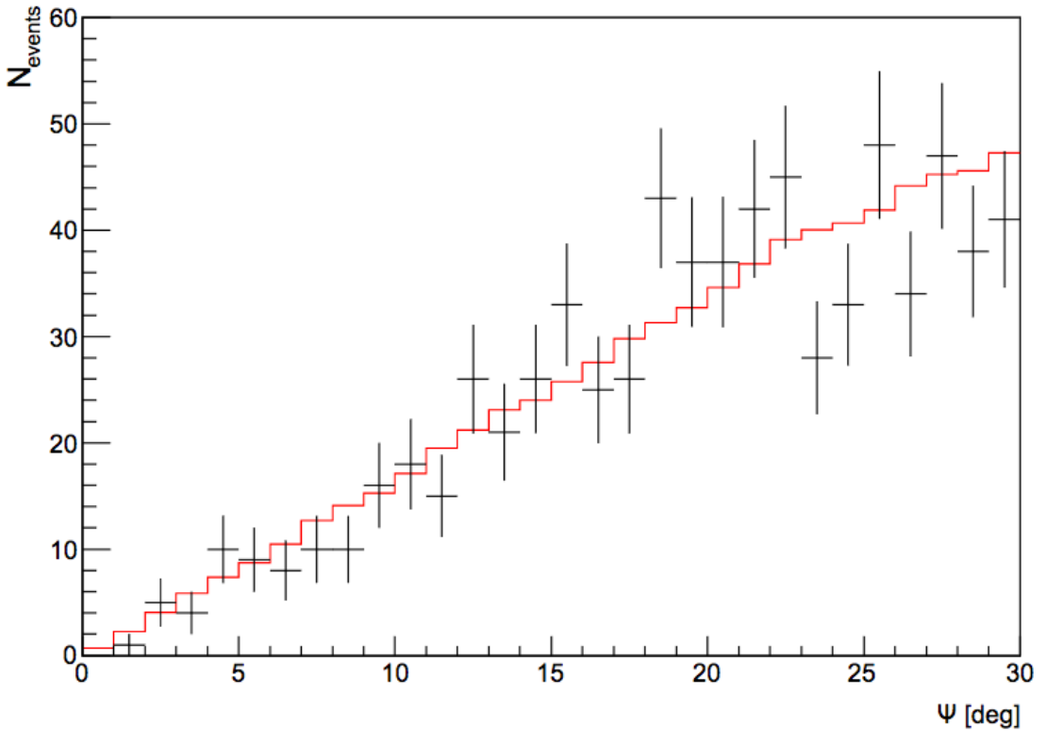
3 years of IceCube-DeepCore data





arXiv:1612.04595

ANTARES 9yrs GC analysis



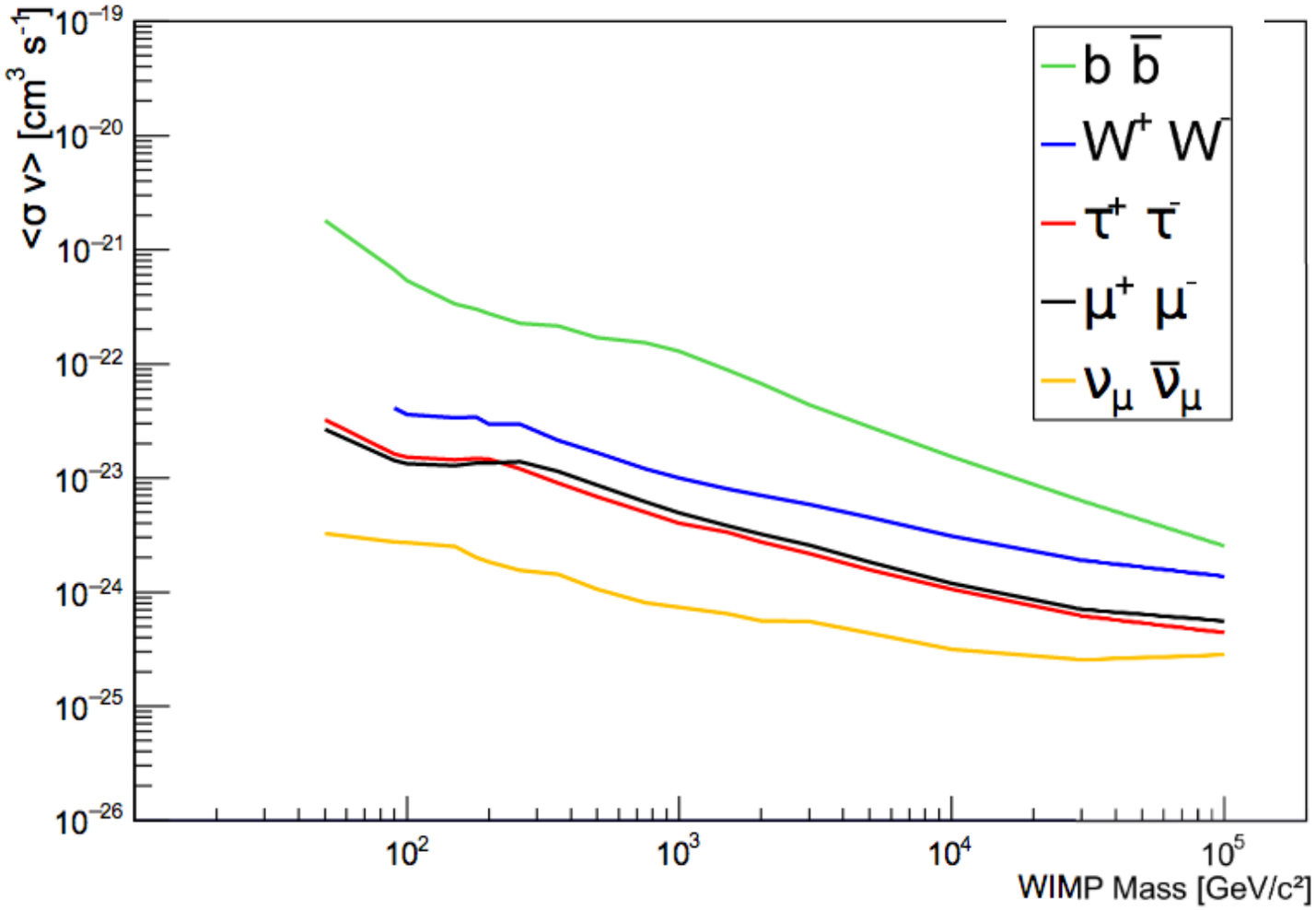
Galactic Center observable in up-going muon neutrino sample for ANTARES

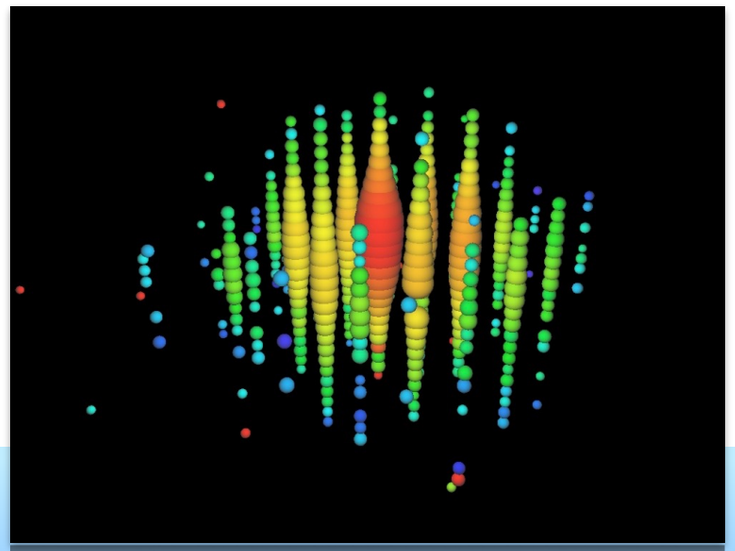
Impressive new limits

Combined analysis of IceCube + ANTARES has started, results for ICRC expected

Combining 9yrs of ANTARES Data

$$\mathcal{A}^j(M_{\mathrm{WIMP}}) = \int A^j_{\mathrm{eff}}(E_{\nu_\mu}) \left.\frac{d\Phi_{\nu_\mu}}{dE_{\nu_\mu}}\right|_{ch,M_{\mathrm{WIMP}}} dE_{\nu_\mu} \cdot T^j_{\mathrm{eff}} + \int A^j_{\mathrm{eff}}(E_{\bar{\nu}_\mu}) \left.\frac{d\Phi_{\bar{\nu}_\mu}}{dE_{\bar{\nu}_\mu}}\right|_{ch,M_{\mathrm{WIMP}}} dE_{\bar{\nu}_\mu} \cdot T^j_{\mathrm{eff}},$$

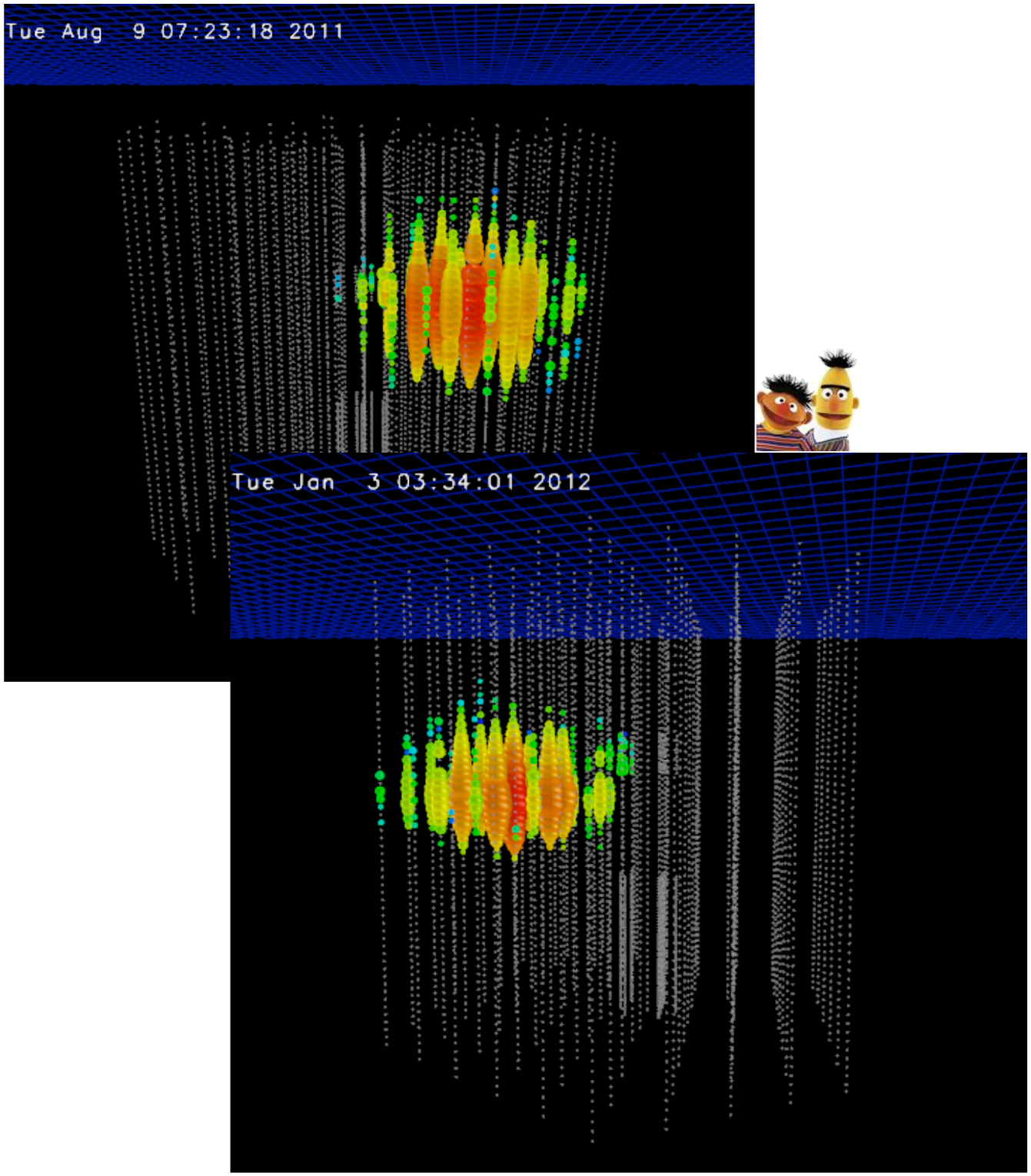




Dark Matter Decay / Astro-physical Neutrinos

Search for highest energy neutrinos

IceCube Coll. Phys.Rev.Lett. 111 (2013) 021103 / arXiv 1304.5356



Dataset / Results
(670days of IC79/IC86 data)
expected 0.08 events
observed 2 events (→ 2.7σ)

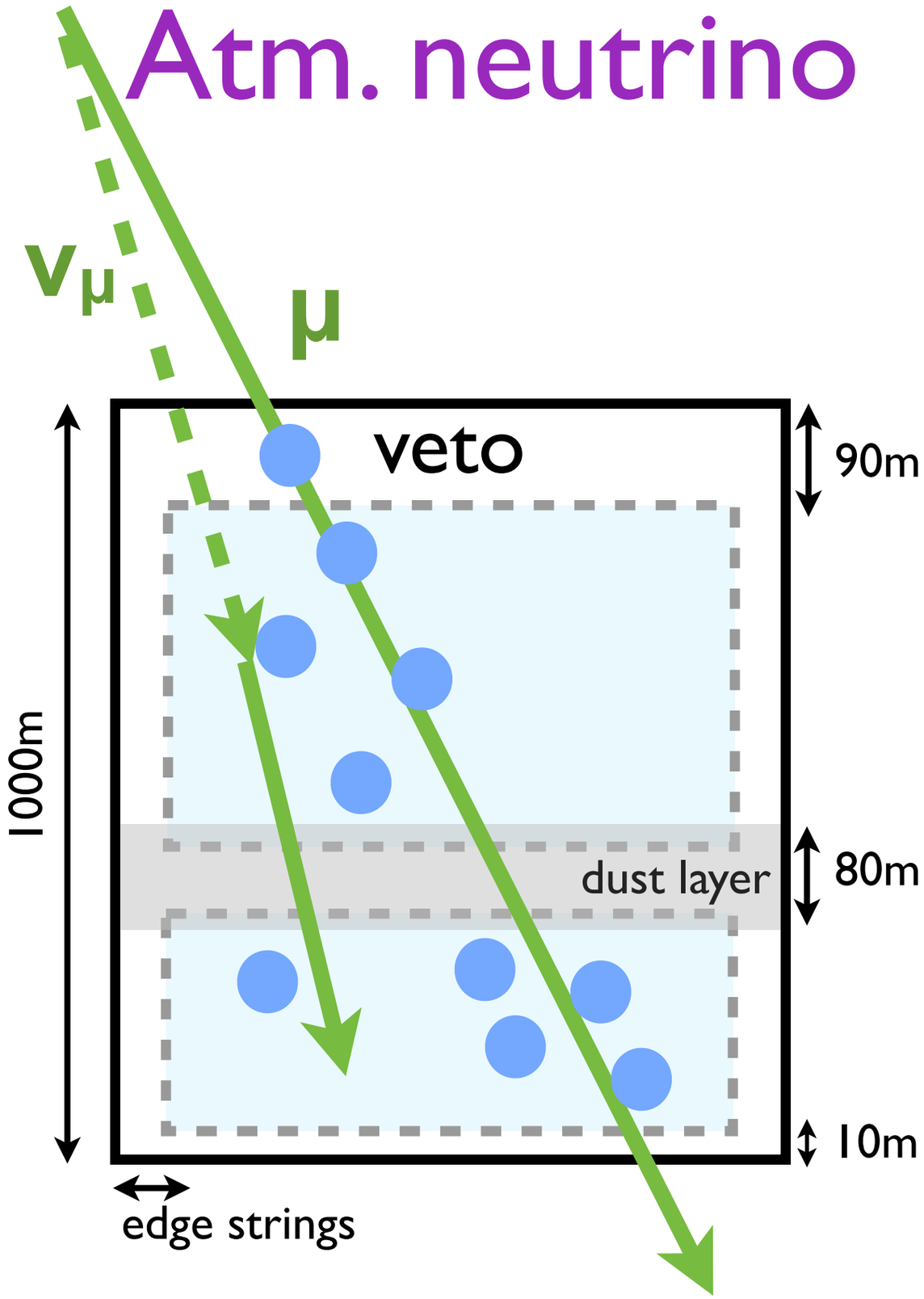
- Ernie ~1.15 PeV (~$1.9 \cdot 10^{-4}$ J)
- Bert ~ 1.05 PeV (~$1.7 \cdot 10^{-4}$ J)
- Energy is the visible energy of the cascade, could originate from NC event, ν_τ CC, or ν_e CC
- Angular resolution on cascade events at this energy ~10
- Energy resolution is about 15% on the deposited energy

Ernie & Bert are not GZK, but ...

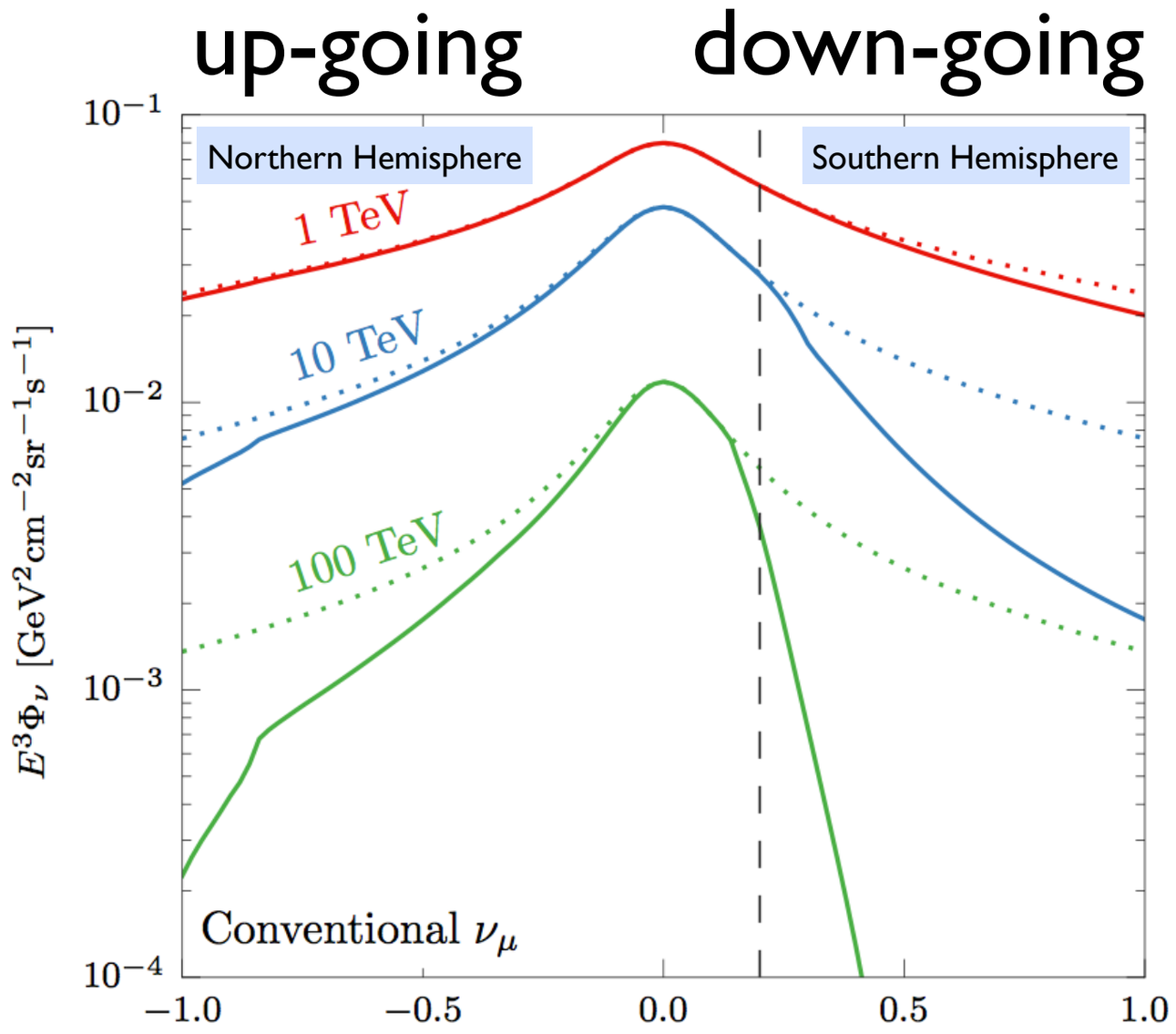
Veto and Self-veto

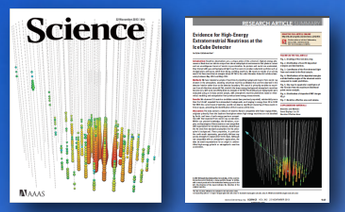
IceCube Collaboration **Phys.Rev. D91 (2015) no.2, 022001** (arxiv:1410.1749)

Atm. neutrino



Down-going high-energy neutrinos can be nearly background free identified as astro-physical neutrinos

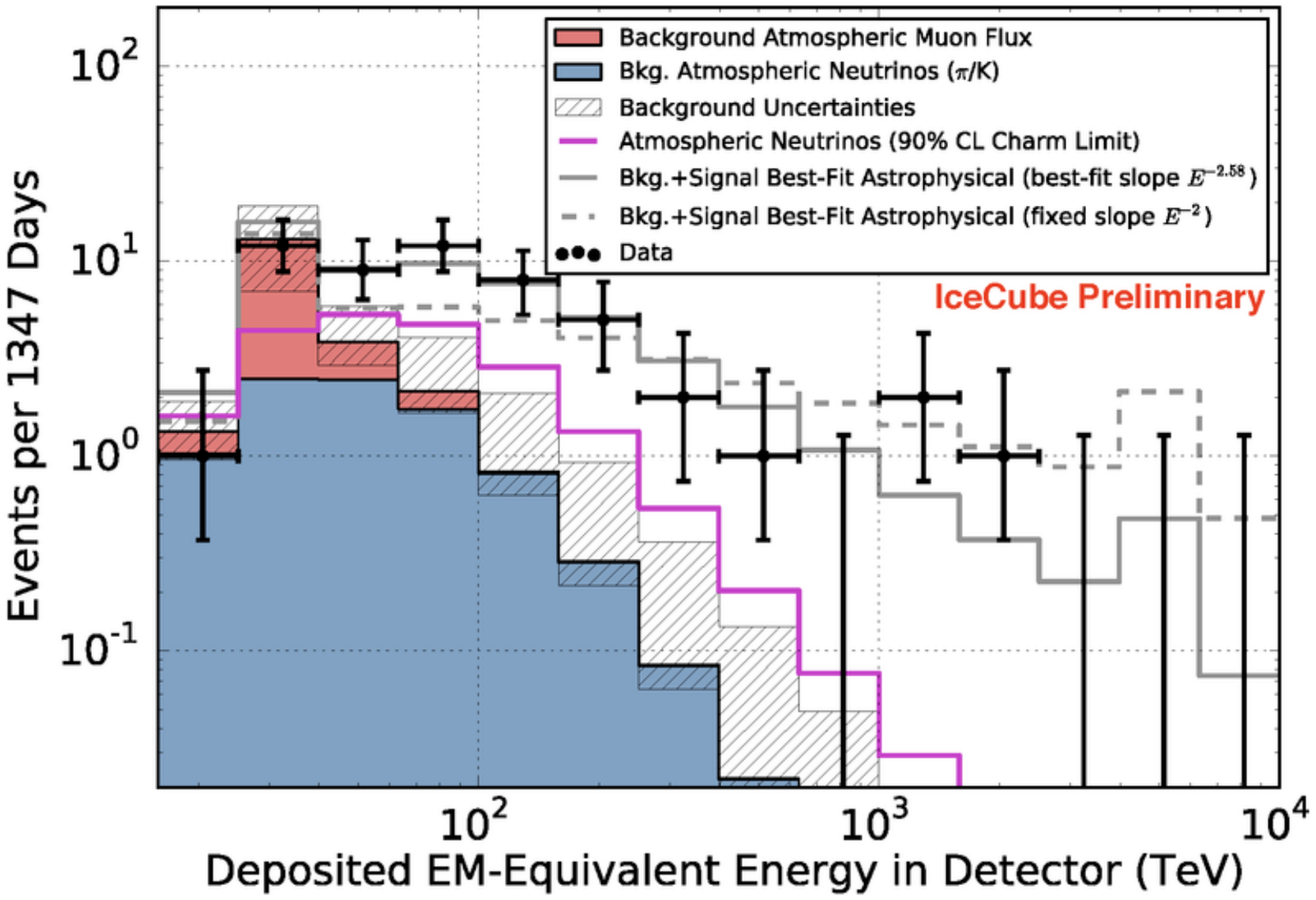




High-energy neutrino search 4yrs

54 events (15 track-like, 39 showers) observed

Expectation from conventional atm. muons and neutrinos ~21.6



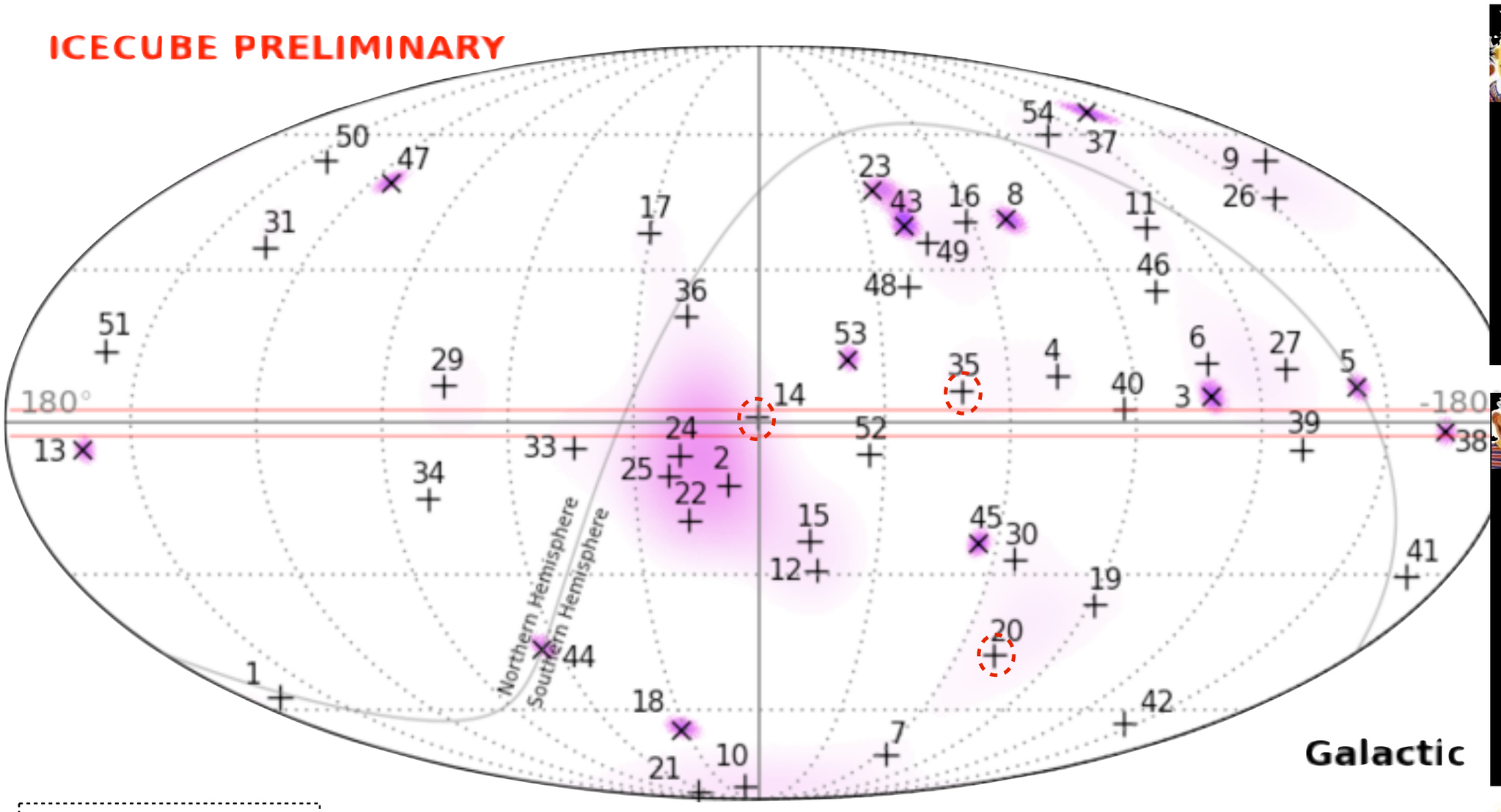
ICRC 2015 proceedings
IceCube Collaboration, *Science 342, 1242856 (2013)*,
IceCube Collaboration, *Phys. Rev. Lett 113, 101101 (2014)*

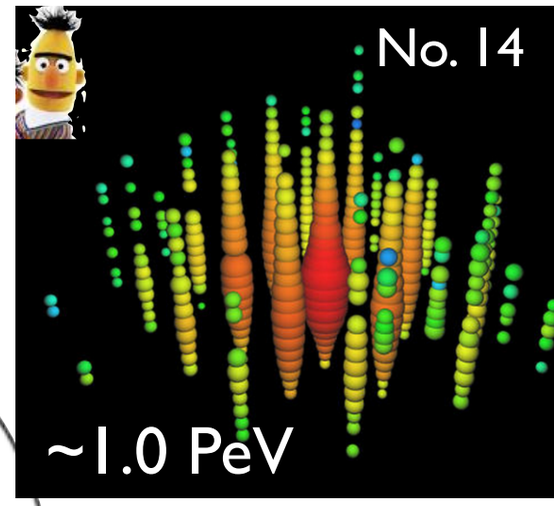
- Mesons including charm quarks in the atmosphere decay immediately to produce neutrinos, known as prompt neutrinos which are not observed yet.
- ERS, or Enberg et al. Phys. Rev. D 78, 043005 (2008) is used as a baseline prompt model
- Significance are based on the exact neutrino flux model, not including the uncertainty of the model.
- Atmospheric Bkg : CR Muon (12.6±5.1), Conv. Neutrino ($9.0^{+8.0}_{-2.2}$),
- Over 60 TeV < E < 2000 TeV, the spectrum best fit with $E^{-2.58}$
- E^{-2} spectrum predicts too may neutrinos above ~2 PeV. So, a cutoff or steeper spectrum needed.

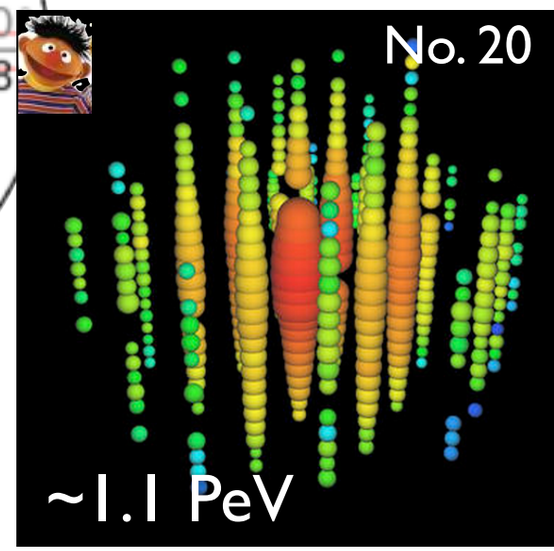
~7 sigma rejection of atmospheric-only hypothesis

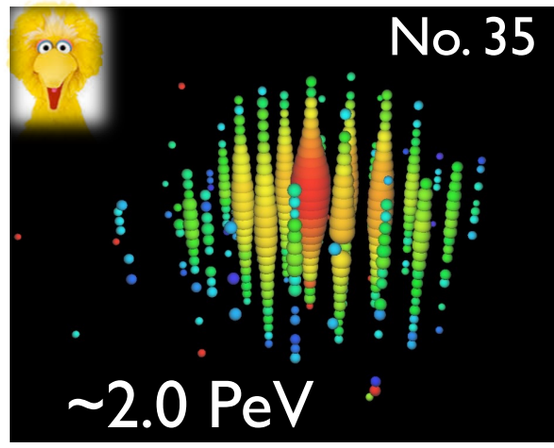
Skymap HESE-4yrs

IceCube Collaboration, *Science* 342, 1242856 (2013)









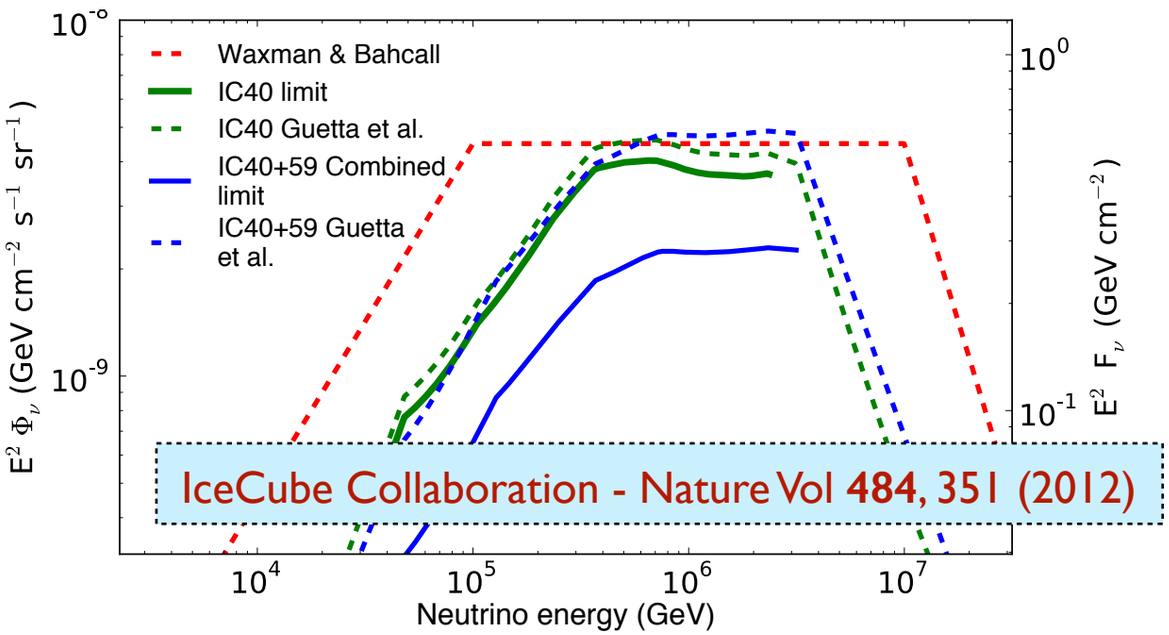
no significant correlations -- spacial or temporal

p-value for cascade events "clustering" 44%

Origin of the high-energy neutrinos ?

Extra Galactic

Gamma Ray Burst

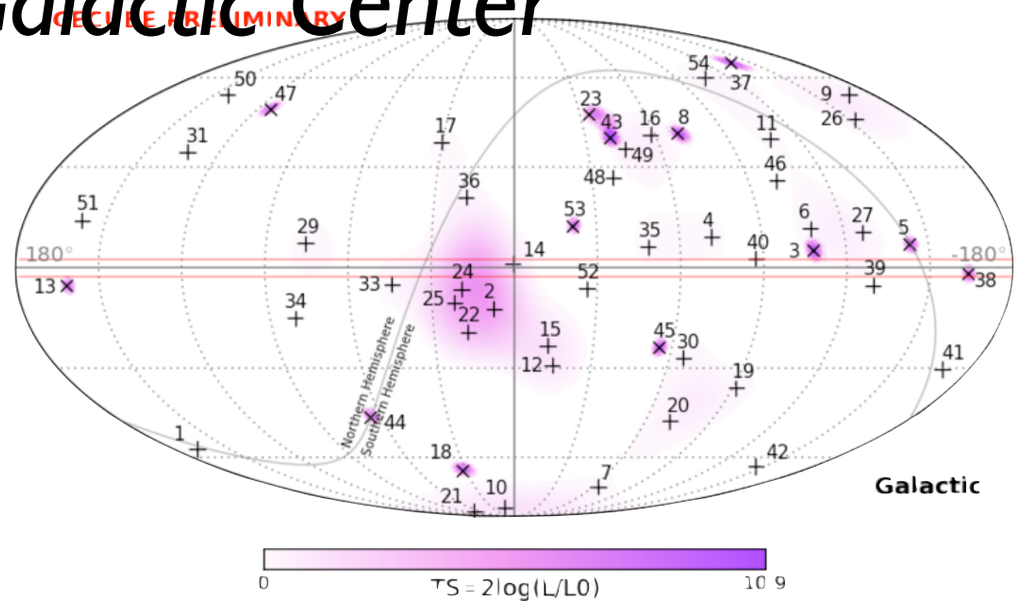


Active Galactic Nuclei / Starburst Galaxies

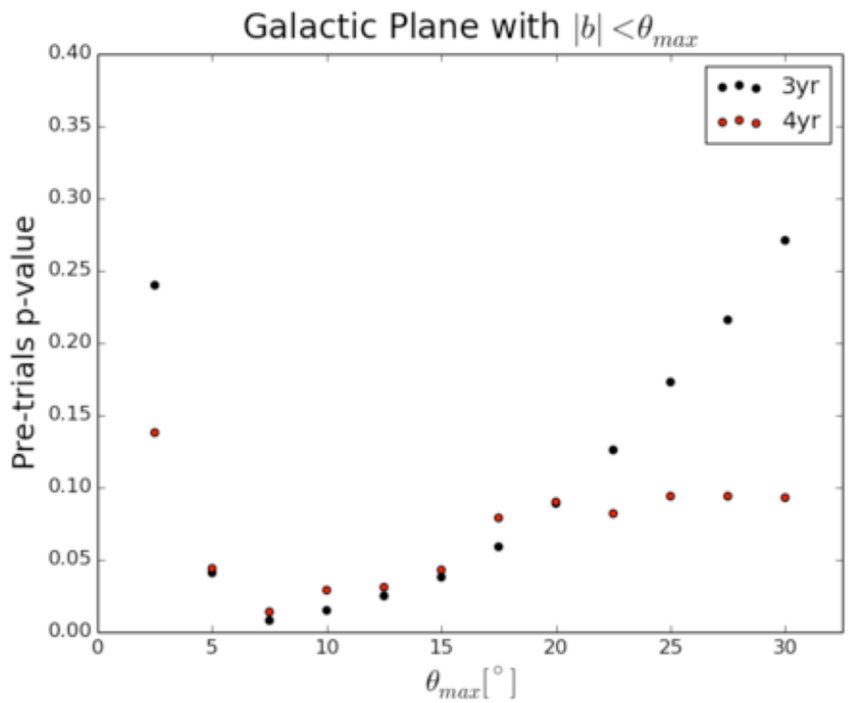
Starburst	M82	148.97	69.68	0.07	0.15
Radio Galaxies	NGC 1275	49.95	41.51	0.0	–
	Cyg A	299.87	40.73	0.9	0.03
	3C 123.0	69.27	29.67	0.0	–
	M87	187.71	12.39	0.0	–
	Cen A	201.37	-43.02	0.03	0.49

Galactic

Galactic Center

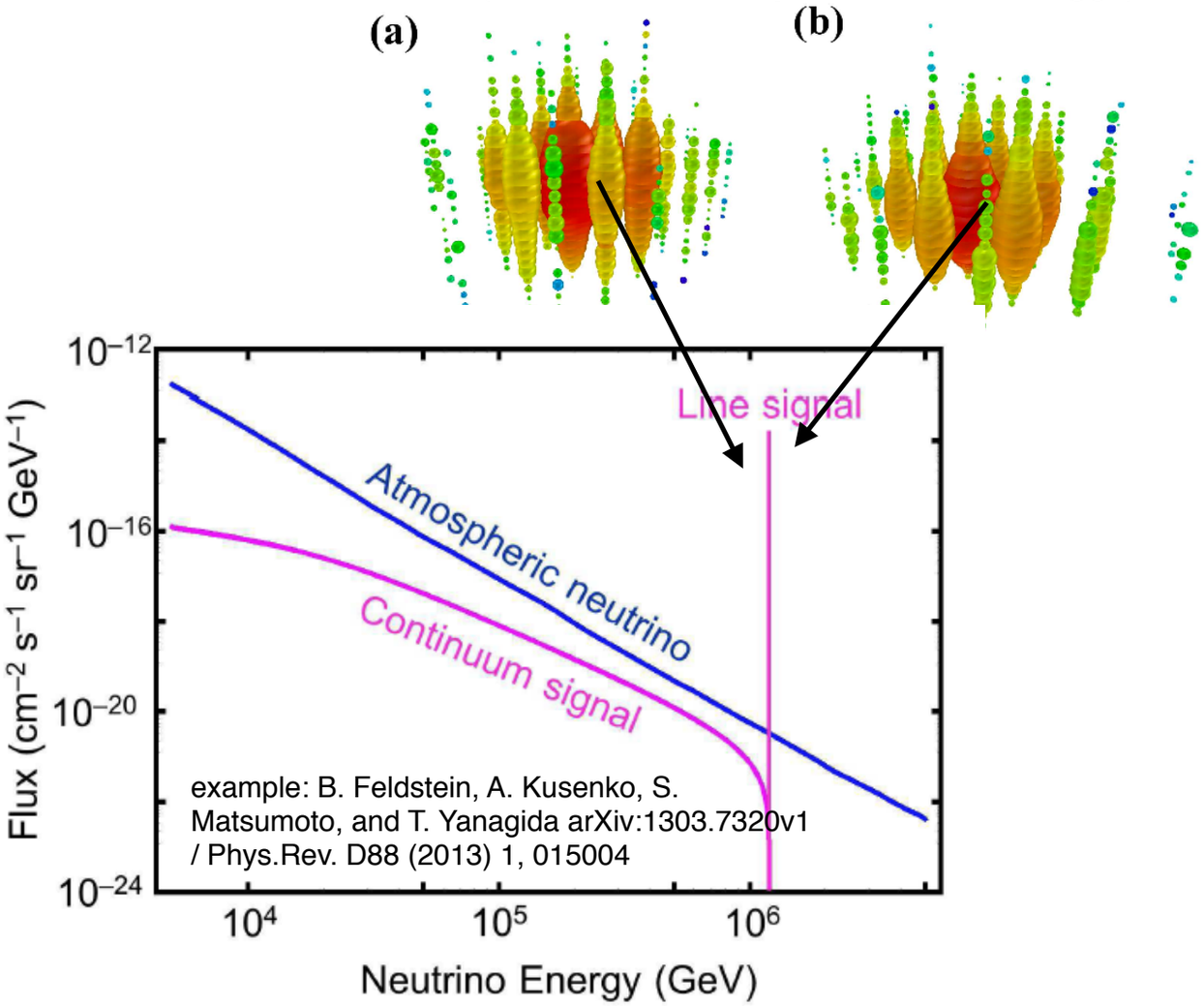


Galactic Plane

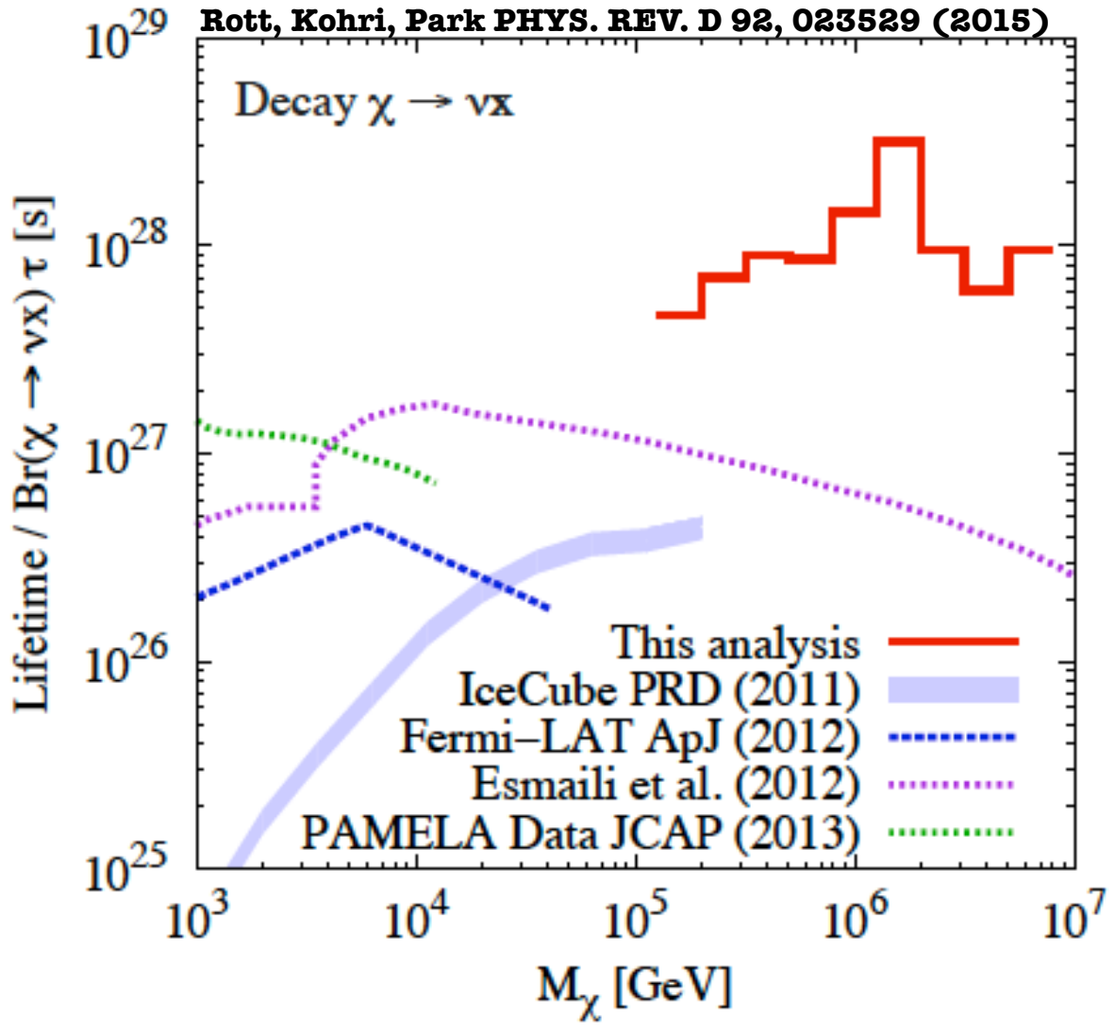


Heavy Dark Matter Decay

- Heavy Decaying Dark Matter (example $\chi \rightarrow \nu h$)
- Focus on most detectable feature (neutrino line)
- Backgrounds steeply falling with energy, highest energy events provide best sensitivity
- Continuum and spacial distribution could help identify a signal
- Bounds from Fermi-LAT and PAMELA derived from search for bb annihilation channel (dominant decay channel of Higgs).



Bound on lifetime ~10^{28}s derived with IceCube data



Heavy DM bounds with neutrinos, see also
Murase and Beacom JCAP 1210 (2012) 043
Esmaili, Ibarra, and Perez JCAP 1211 (2012) 034
El Aisati, Gustafsson, Hambye 1506.02657

Dedicated IceCube analysis should improve on these bounds Analyses on-going

Heavy Dark Matter Decay

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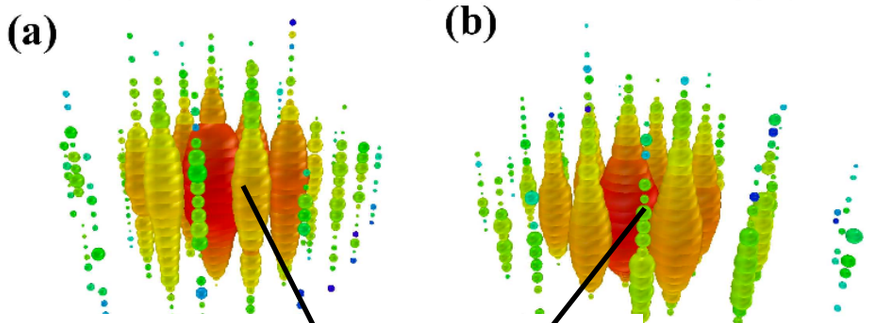

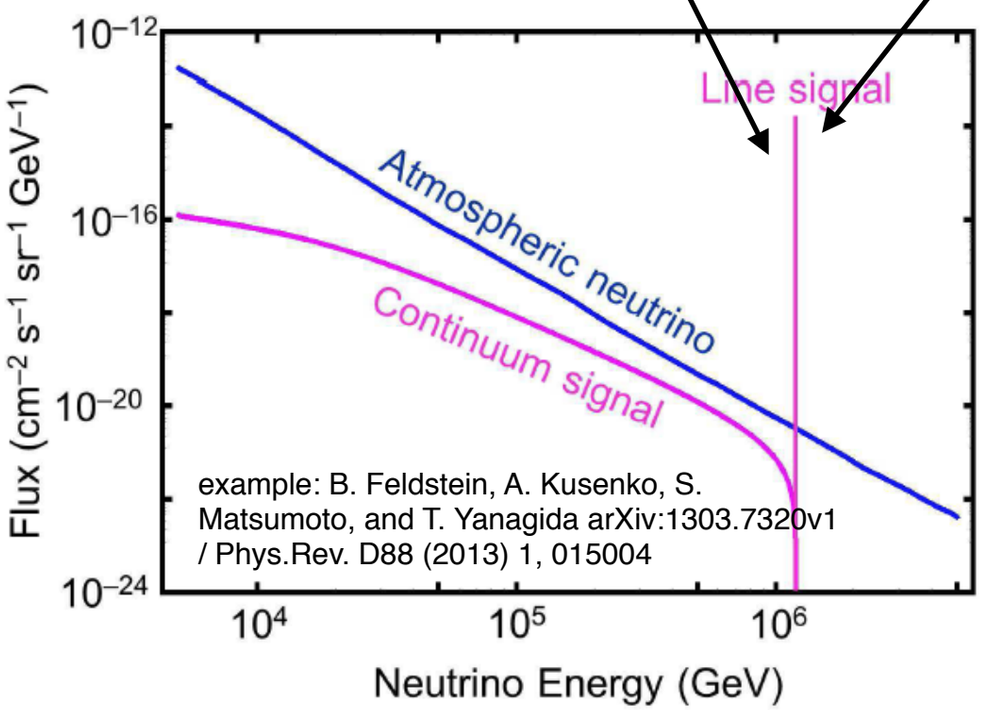

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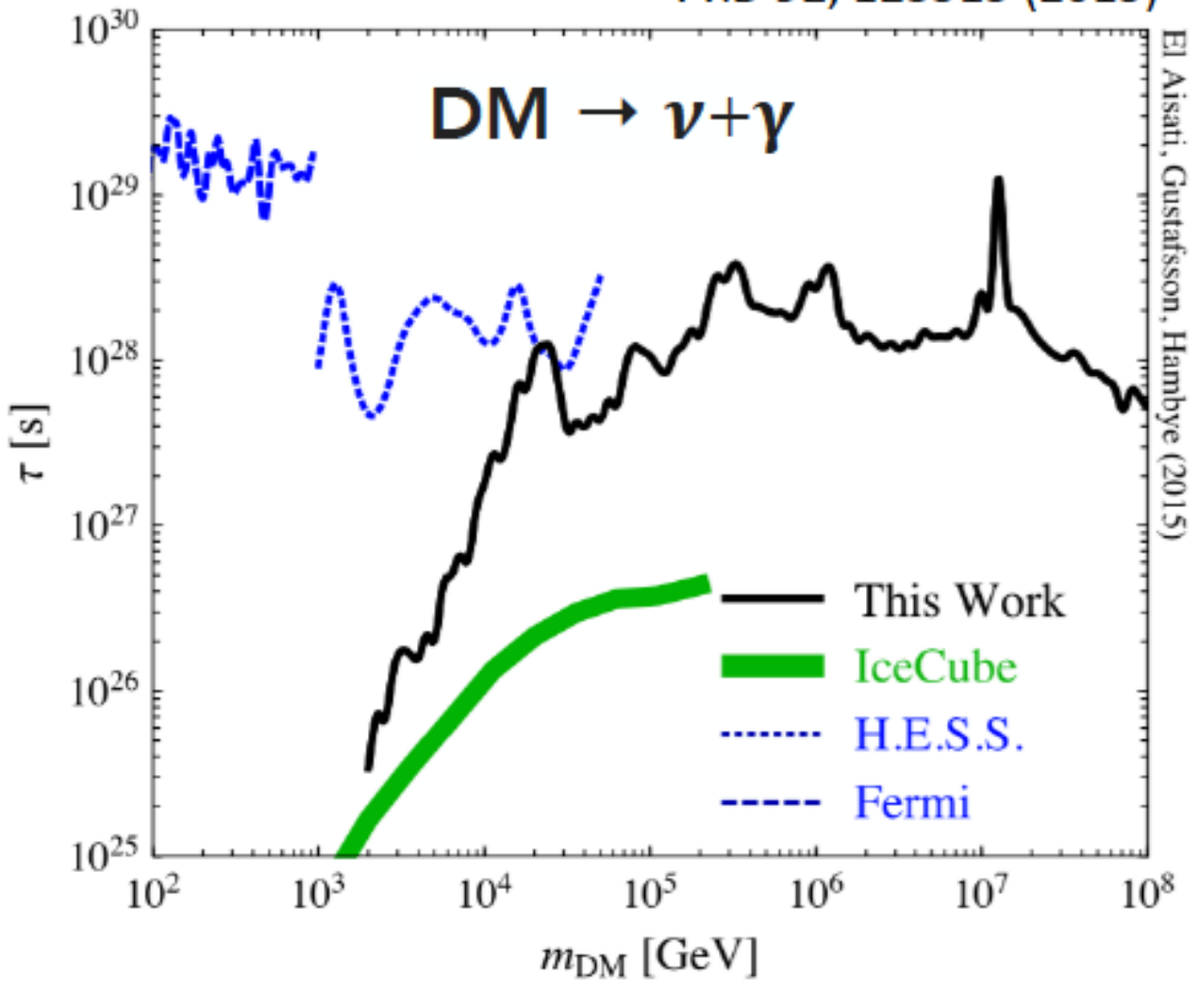
PRD 92, 123515 (2015)



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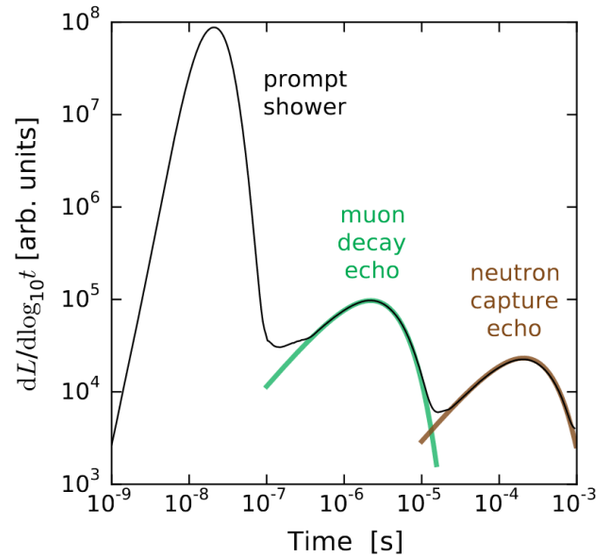
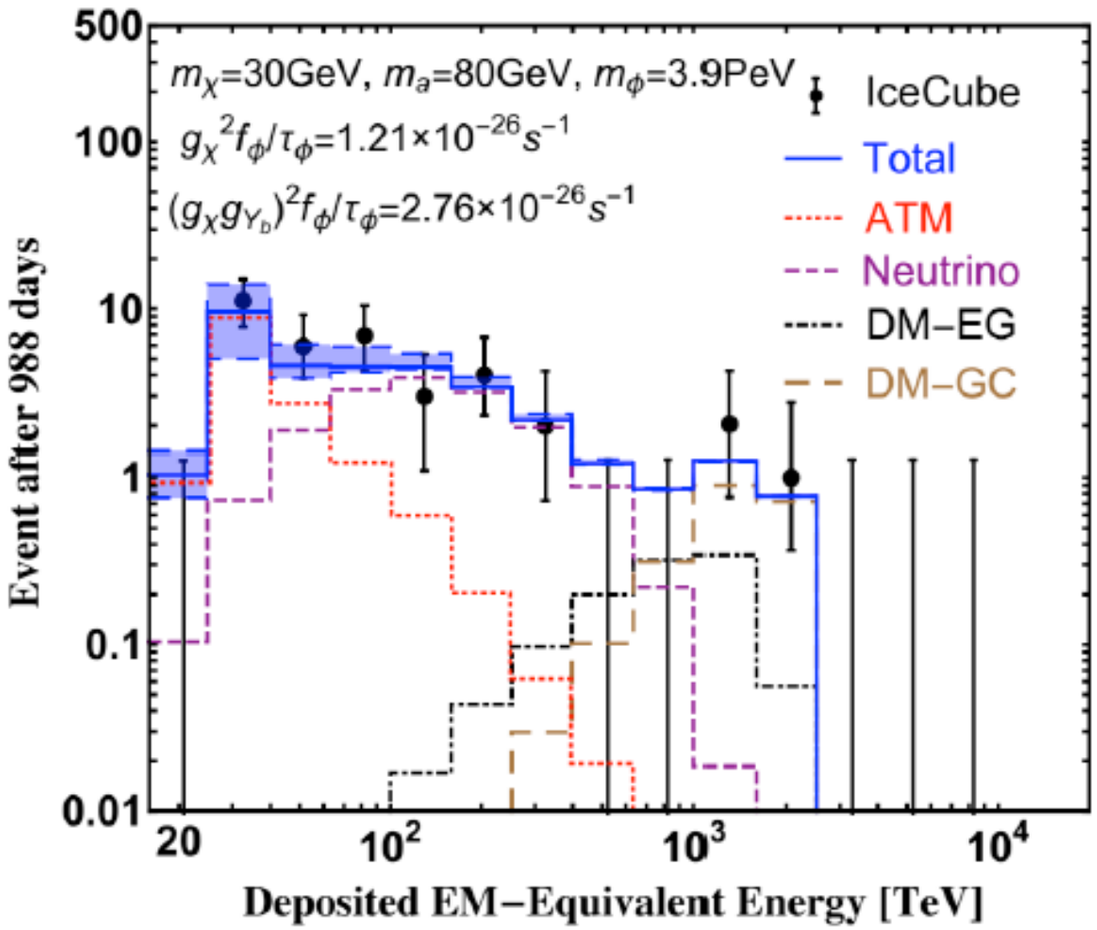
Dedicated IceCube analysis should improve on these bounds Analyses on-going

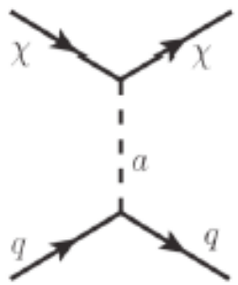
Boosted Dark Matter

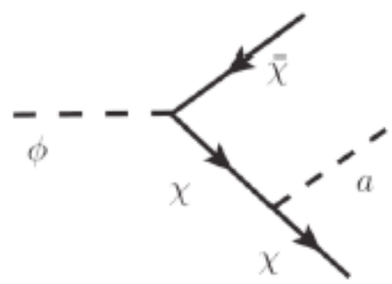
- "Boosted Dark Matter Search"
  - Following search proposed by Kopp, Liu, Wan (2015)
  - using "Echo Technique" Li, Bustamante, Beacom (2016)





Very heavy dark matter particle ϕ decays to lighter stable dark matter χ →boost!





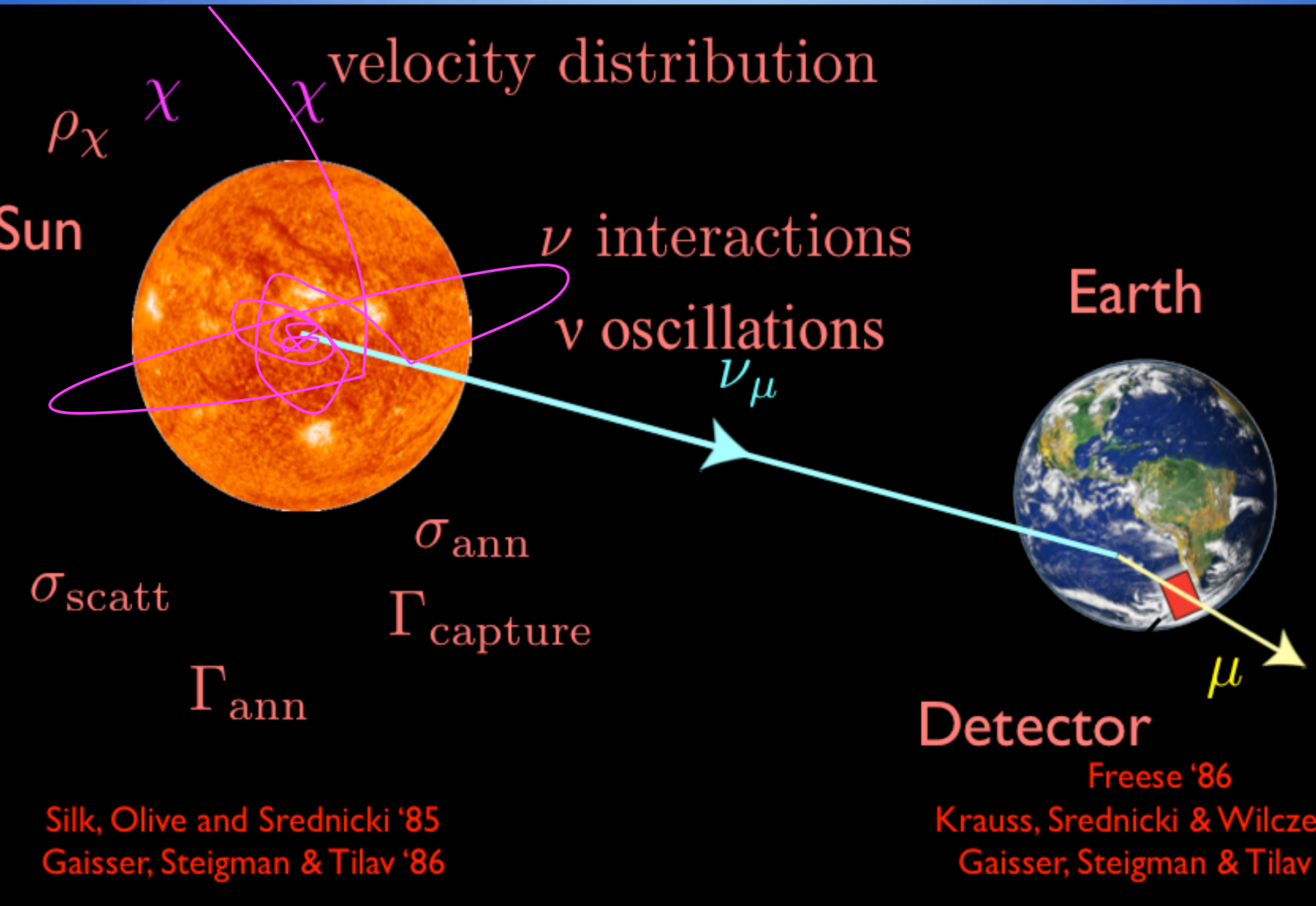
Recoil (only hadronic cascades)

$\phi \to \chi\,\bar{\chi}\,a$, $a \to b\,\bar{b}$ →ν's

May sound crazy, but is just an example for exotic interactions in IceCube detectable via recoil

Dark Matter Capture in the Sun

Solar WIMPs



Silk, Olive and Srednicki '85
Gaisser, Steigman & Tilav '86

Freese '86
Krauss, Srednicki & Wilczek '86
Gaisser, Steigman & Tilav '86

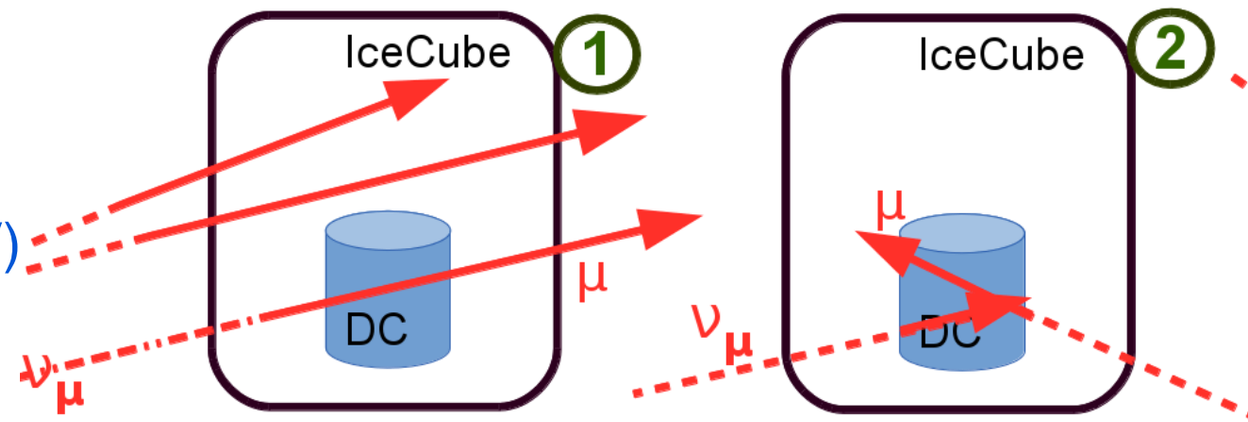
3yrs IceCube Solar WIMP Analysis

IceCubeColl., arXiv:1612.05949v1

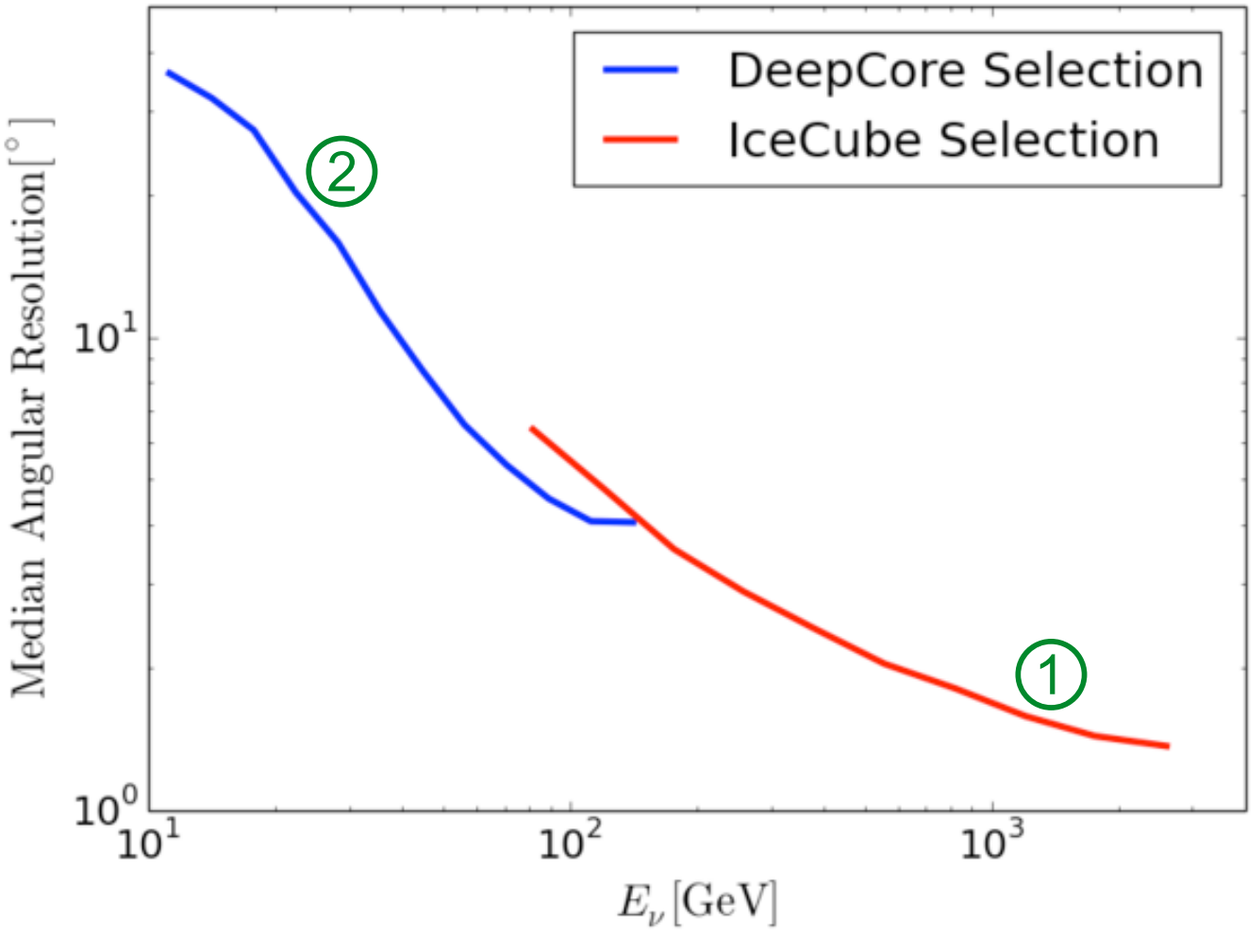
- Three years of data in 86-string configuration used (**May 2011 - May 2014**)
  - Only **up-going** events (Sun below the horizon) results in **532days of livetime**
- Two independent analysis performed
  - ① **IceCube**: Higher energy focus ($m_\chi > 100$GeV)
  - ② **DeepCore**: Low-energy focus ($m_\chi = 30$GeV - 100GeV)

- *Up-going*
- *IceCube Dominated*
- *No Containment*

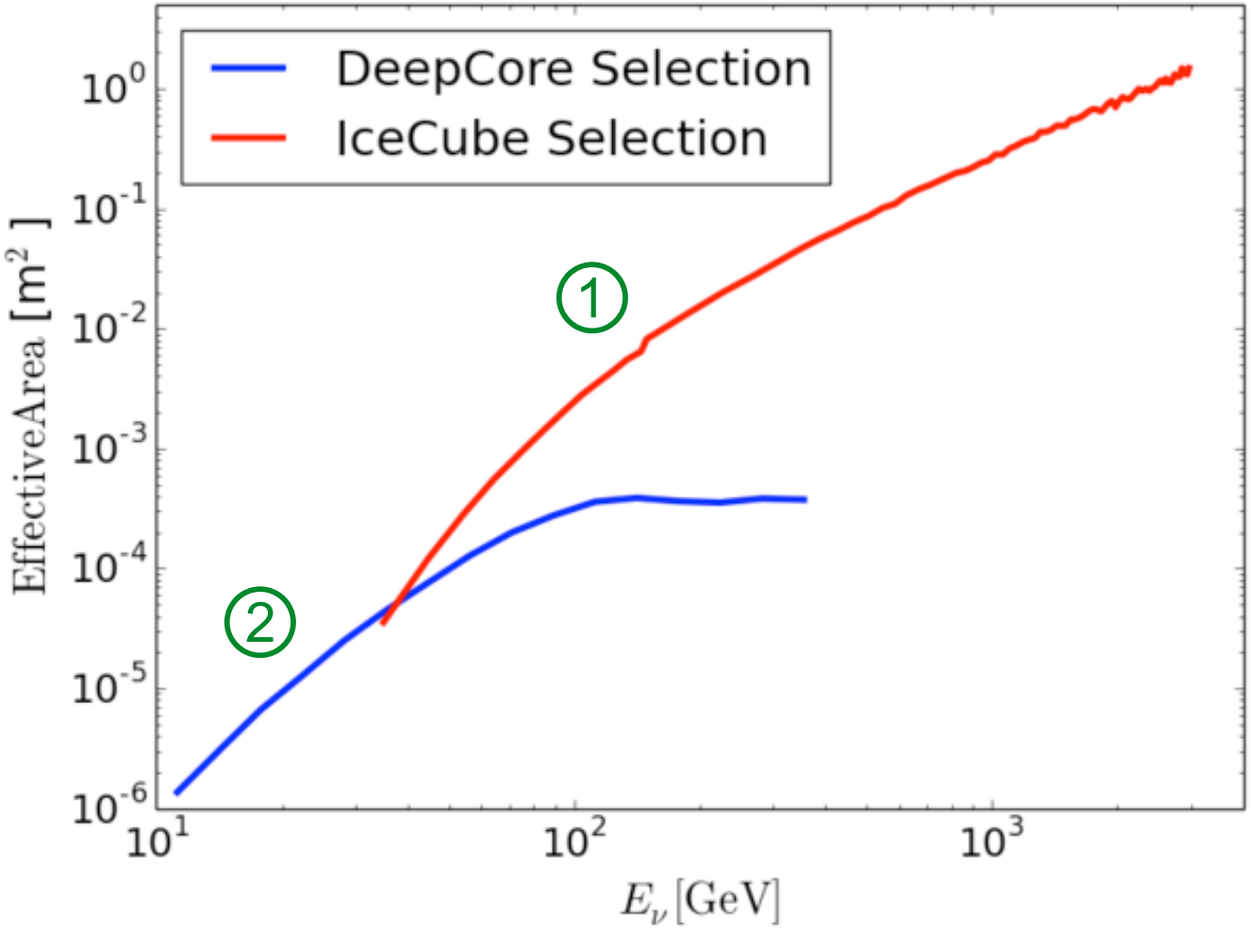
- *Up-going*
- *DeepCore Dominated*
- *Strong Containment*



Median angular resolutions

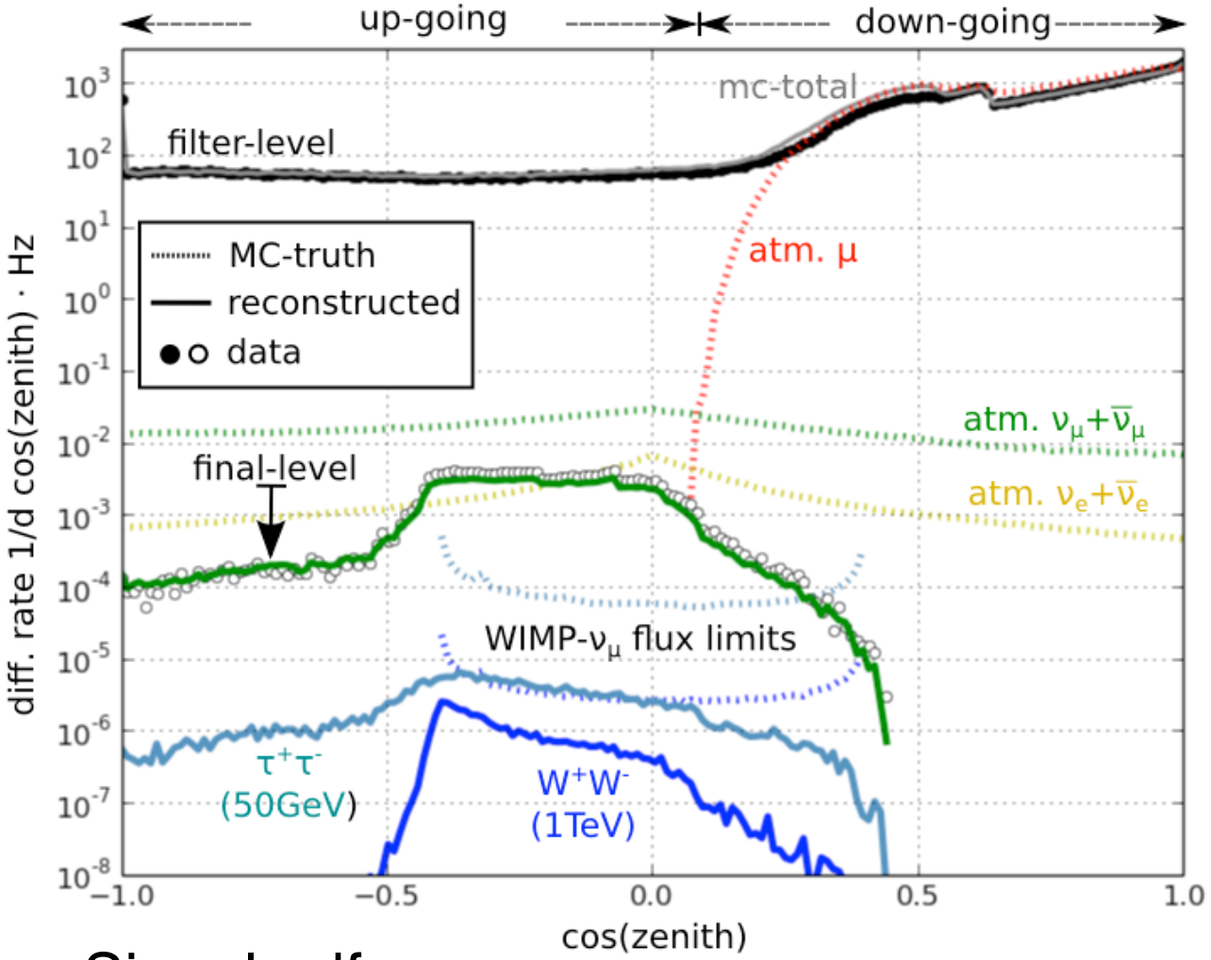


Effective Areas

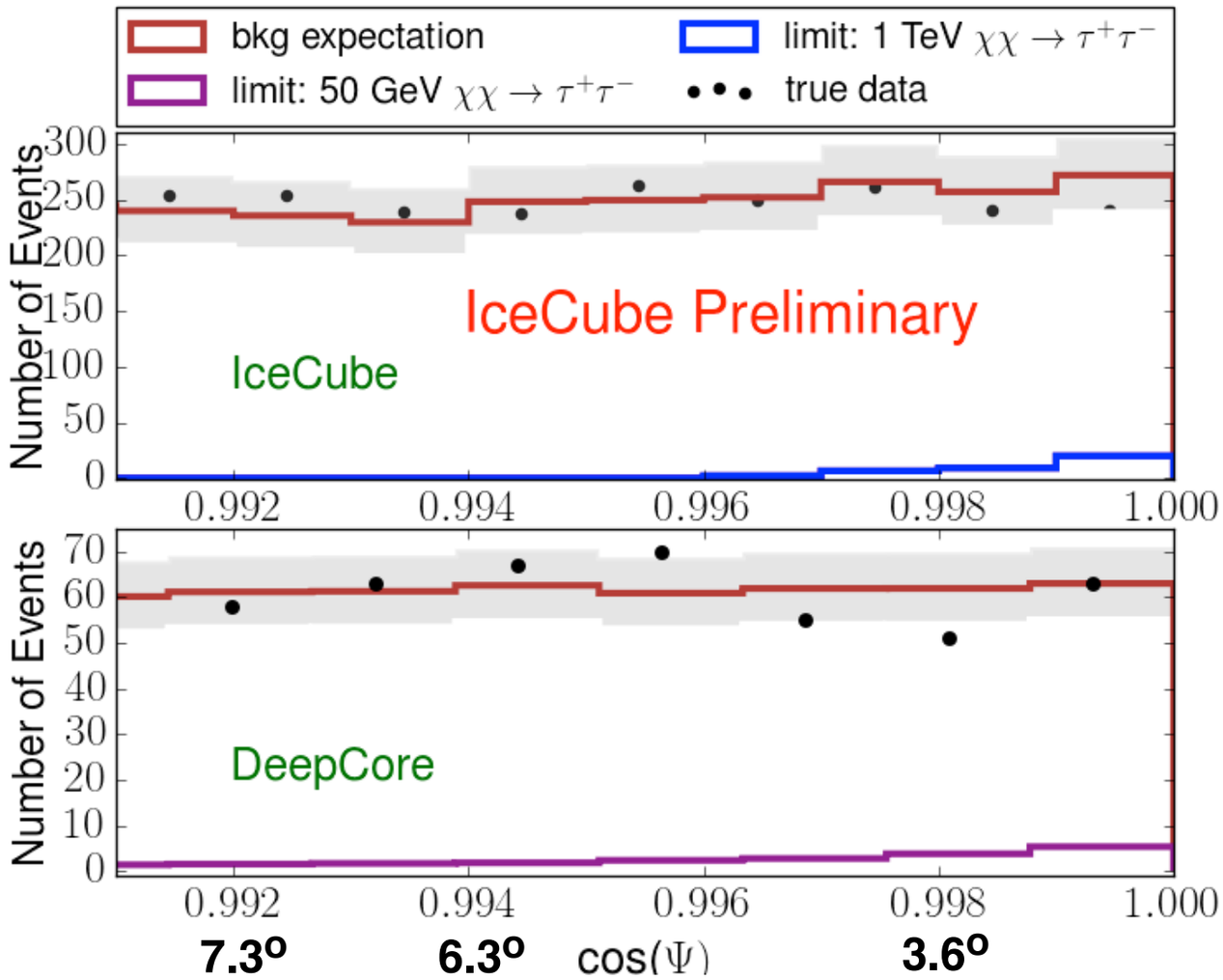


3yrs IceCube Solar WIMP Analysis

IceCubeColl., arXiv:1612.05949v1



Observed events



Signal pdf:

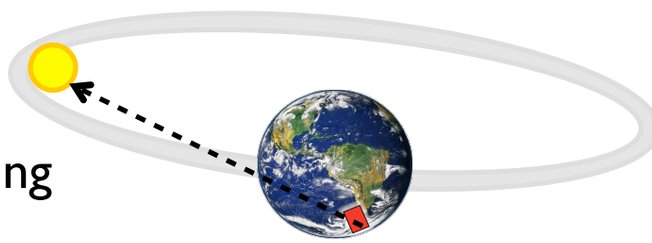
Spectral part

$$S_i(|\vec{x}_i - \vec{x}_{\rm sun}(t_i)|, E_i, m_\chi, c_{\rm ann}) = \mathscr{K}(|\vec{x}_i - \vec{x}_{\rm sun}(t_i)|, \kappa_i) \times \mathscr{E}_{m_\chi, c_{\rm ann}}(E_i)$$

Monovariate Fisher Bingham distribution from directional statistics

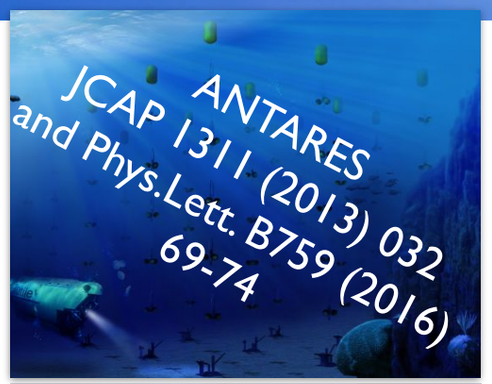
Background pdf: $\mathscr{B}_i(tx_i, E_i) = B(\delta_i) \times P(E_i|\phi_{\rm atm})$

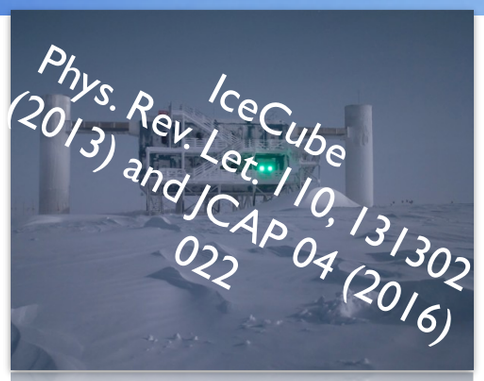
Likelihood: $\mathscr{L}(n_{\rm s}) = \prod_N \left(\frac{n_{\rm s}}{N} S_i + (1 - \frac{n_{\rm s}}{N})\mathscr{B}_i\right)$

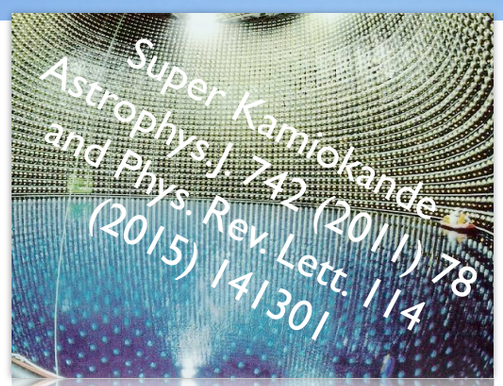


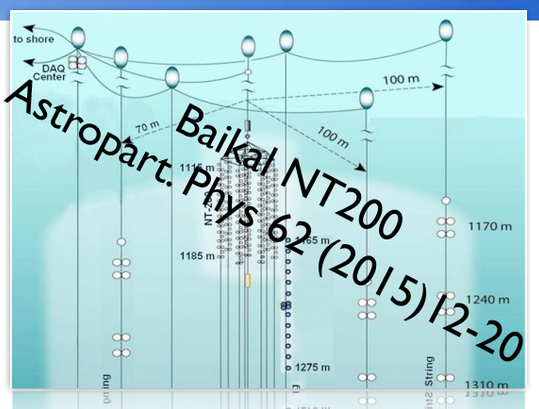
- Use track events for better pointing
- Search for an excess of events from the direction of the Sun
- Observed events consistent with background only expectations

Solar WIMPs Summary

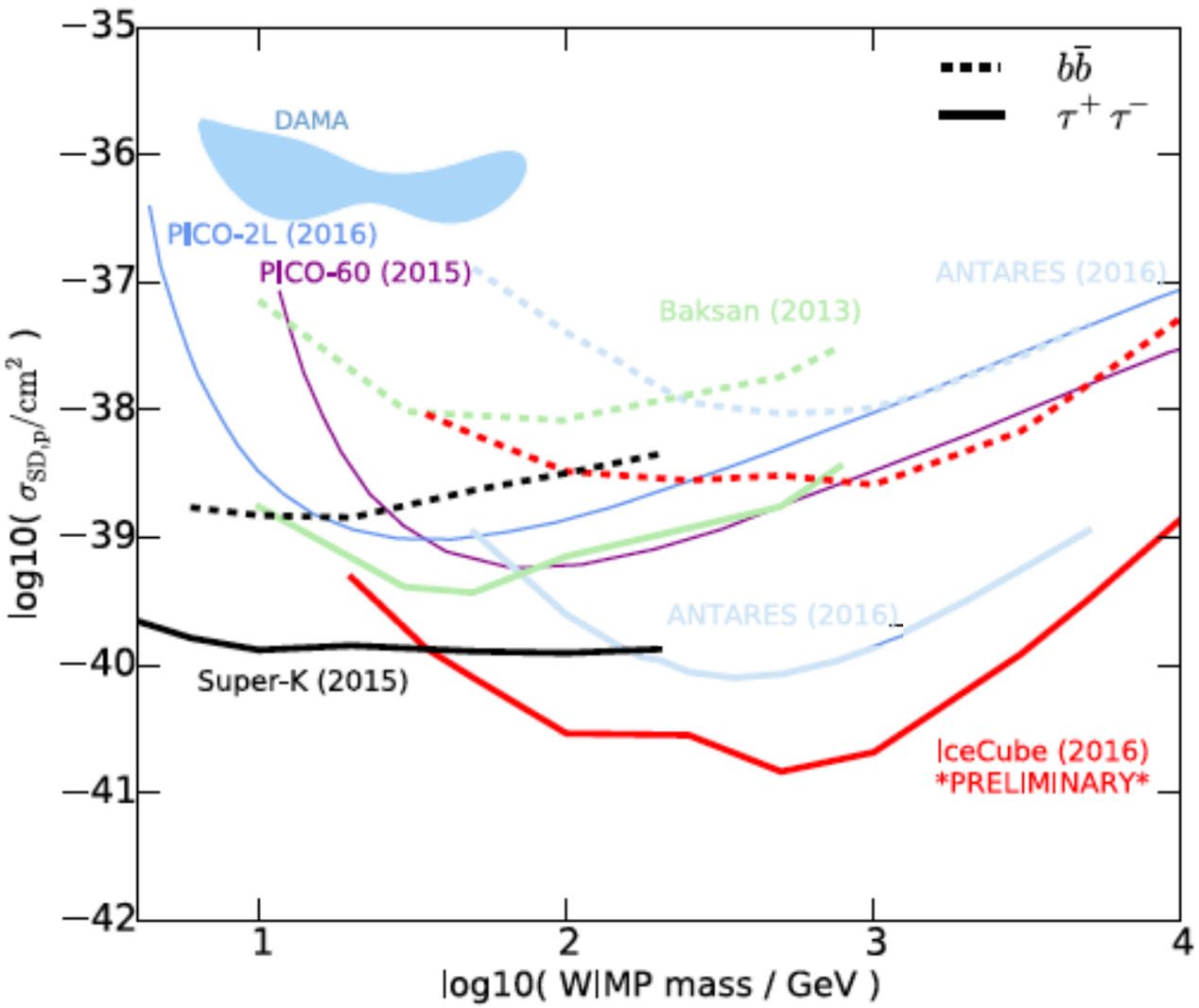




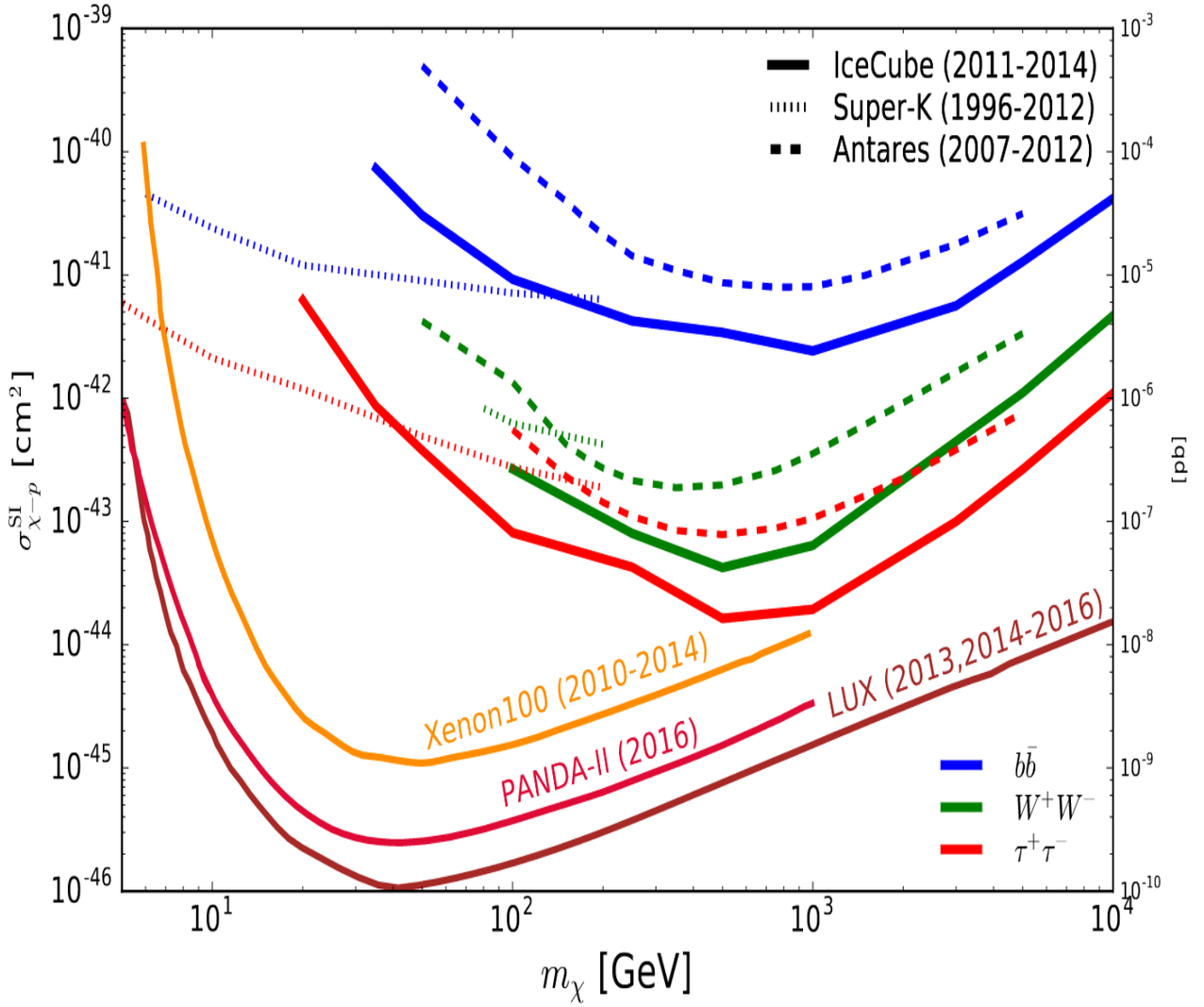




Spin-dependent scattering



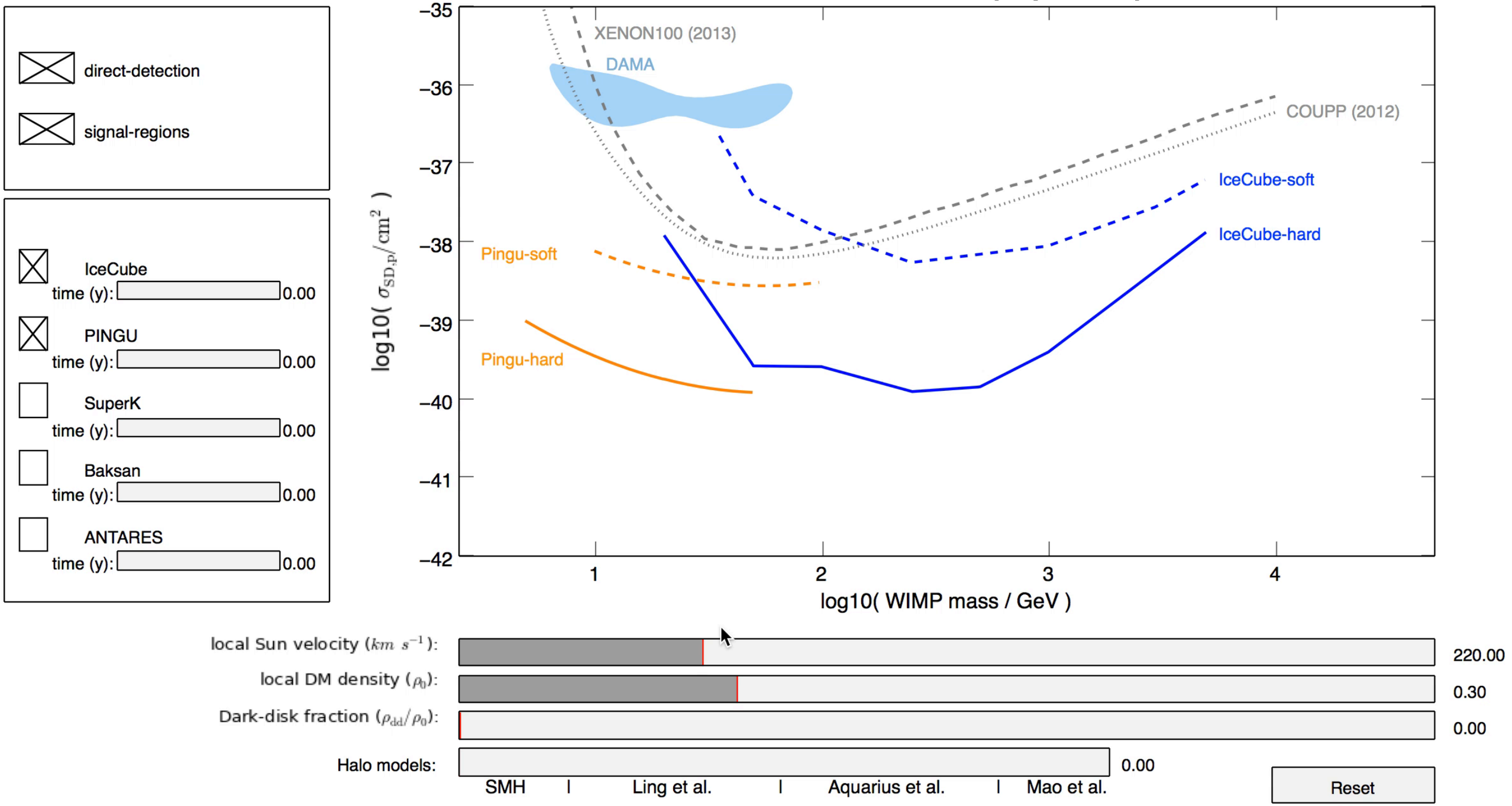
Spin-independent scattering



Impact of astrophysical uncertainties

M. Danninger & C. Rott "Solar WIMPs Unraveled" – Physics of the Dark Universe (Nov 2014)

Interactive tool to study impact of astrophysical parameters

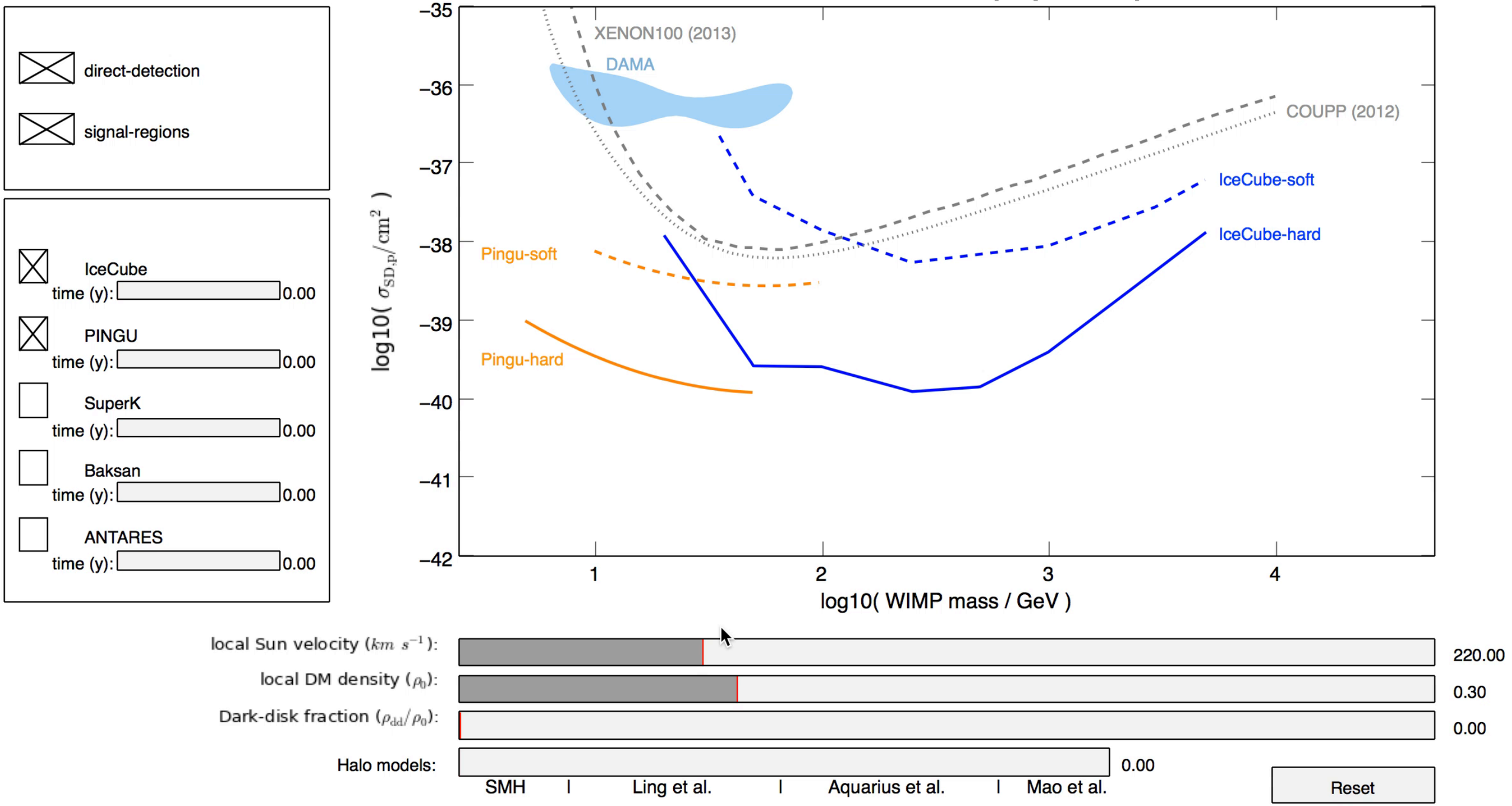


https://mdanning.web.cern.ch/mdanning/public/Interactive_figures/

Impact of astrophysical uncertainties

M. Danninger & C. Rott "Solar WIMPs Unraveled" – Physics of the Dark Universe (Nov 2014)

Interactive tool to study impact of astrophysical parameters



https://mdanning.web.cern.ch/mdanning/public/Interactive_figures/

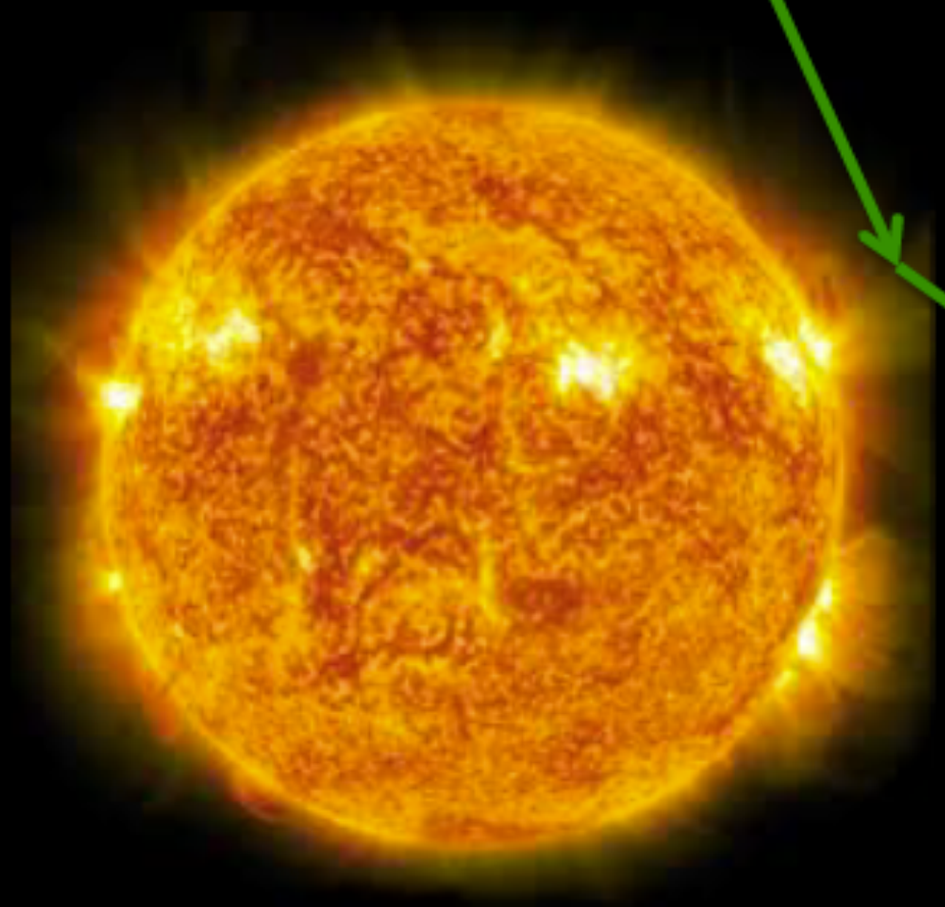
Solar Neutrino Floor

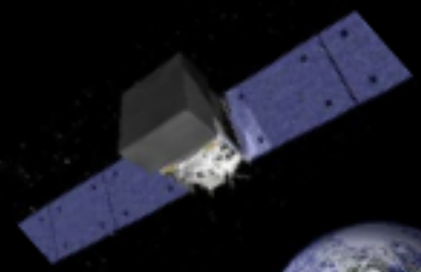
see K. Ng, J. Beacom, A. Peter, C. Rott PRD 2016

In preparation Ng, Beacom, Peter, Rott

Sun – Cosmic-Ray Beam Dump

- Leptonic

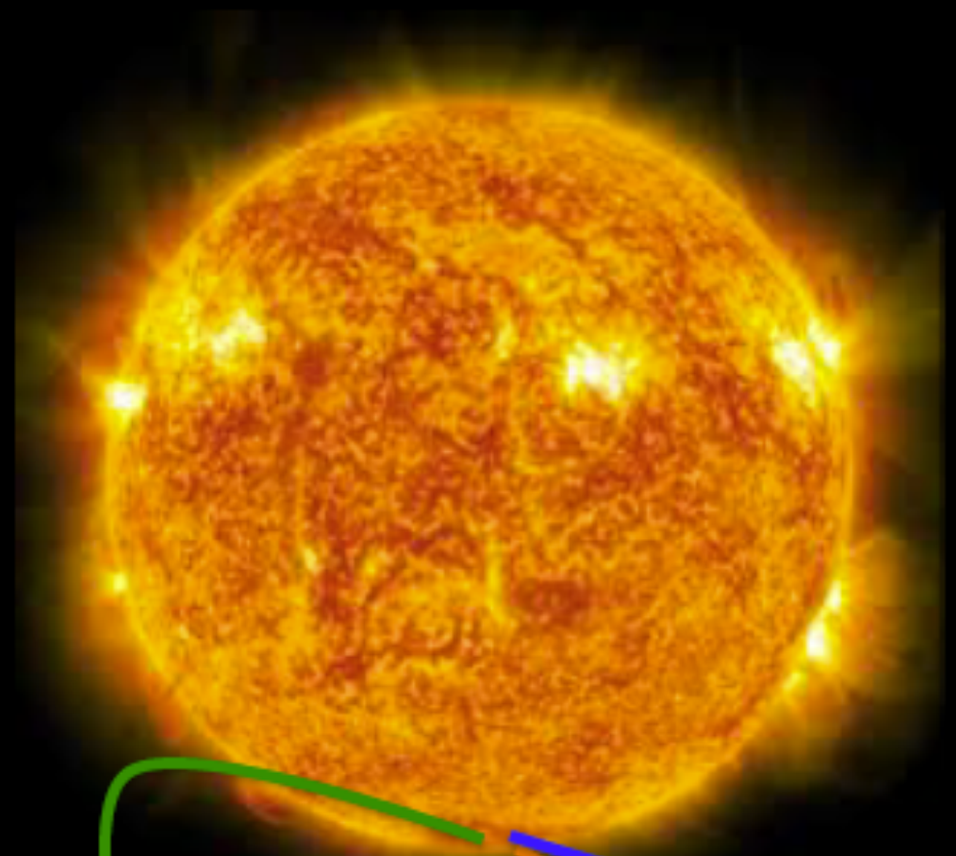




Moskalenko, Porter, Digel (2006)
Orlando, Strong (2007)

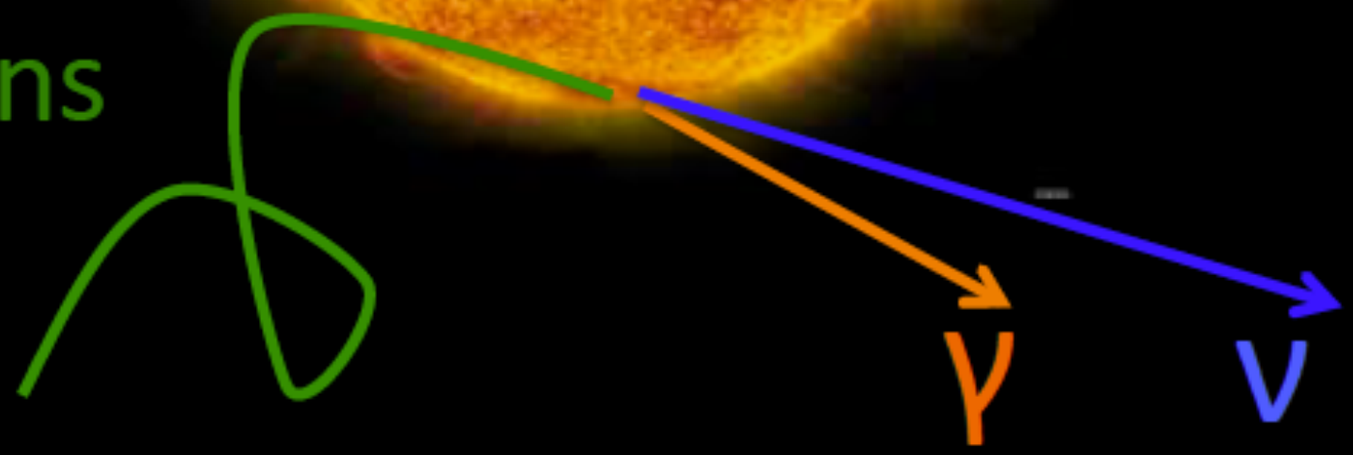
Sun – Cosmic-Ray Beam Dump

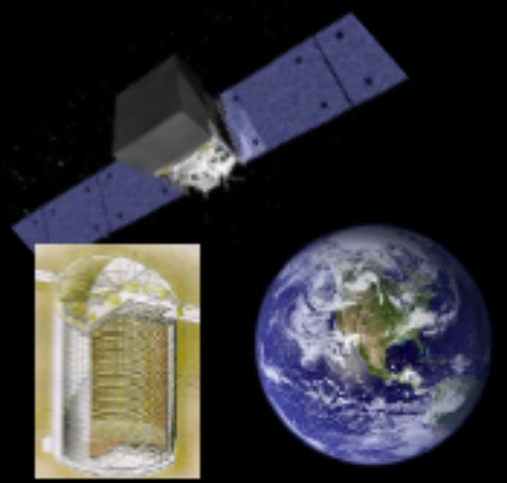
- Hadronic



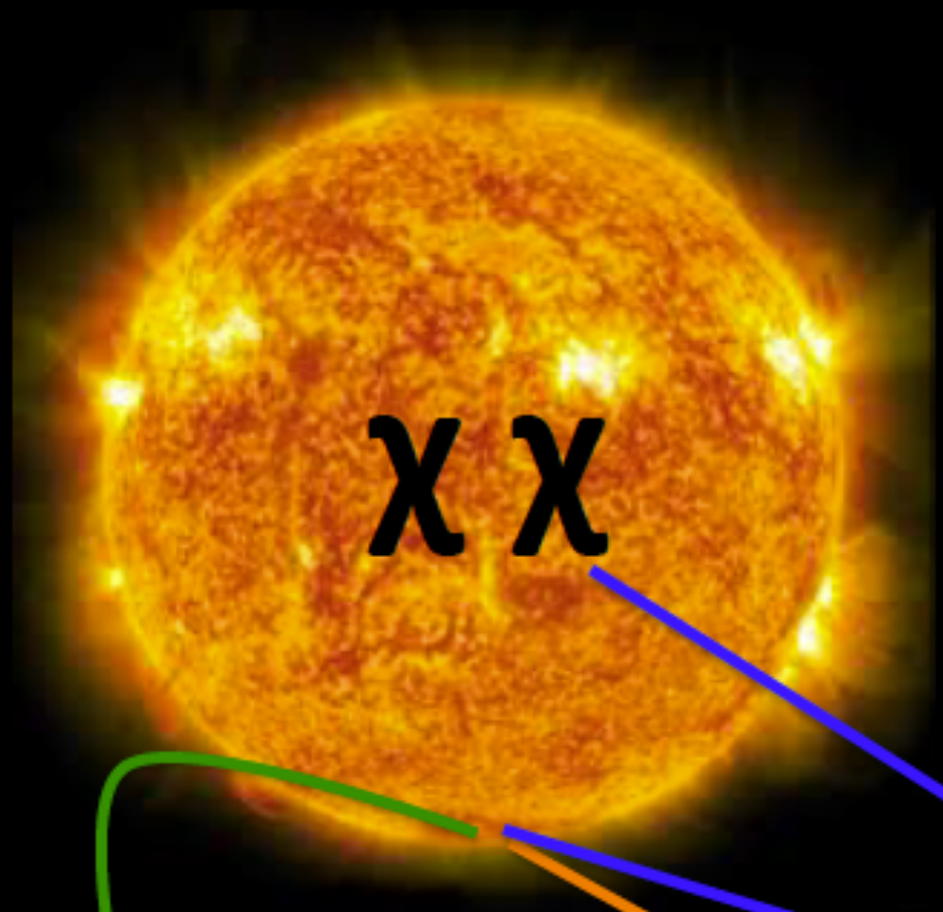
Seckel, Stanev, Gaisser (1991)
Moskalenko, Karakula (1993)
Ingelman, Thunman (1996)

CR protons

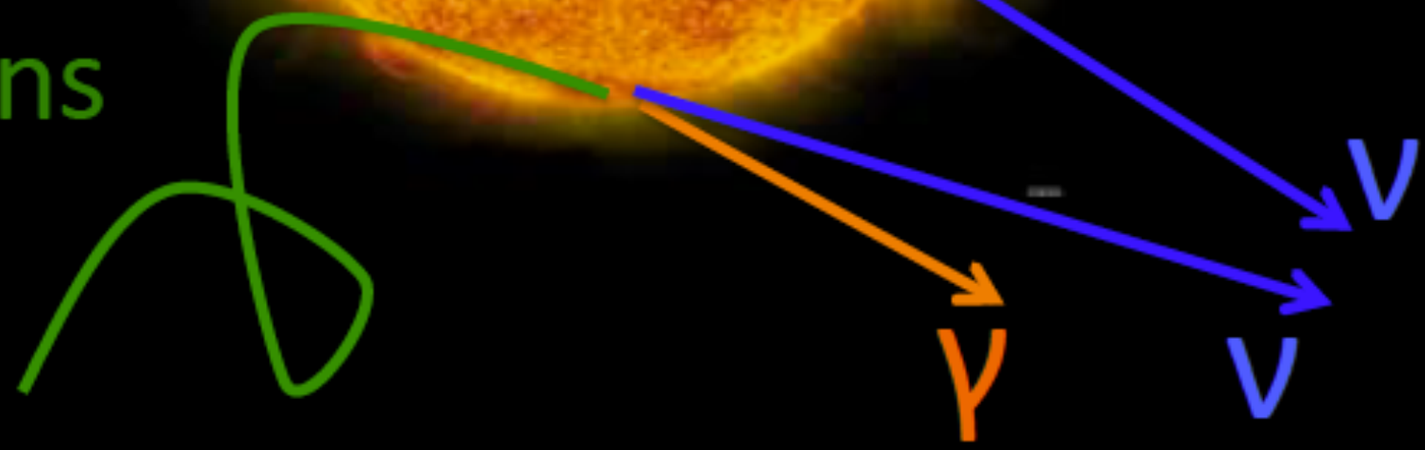


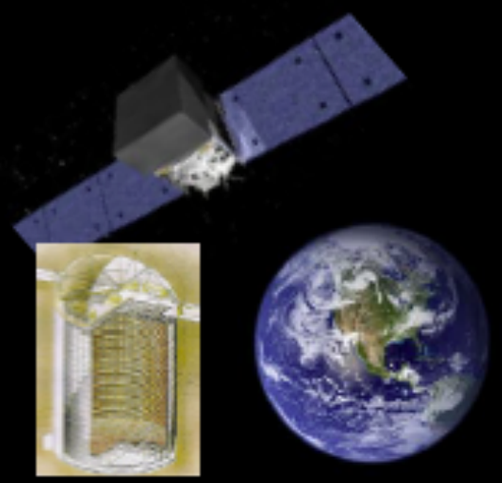


Cosmic Rays vs Dark Matter



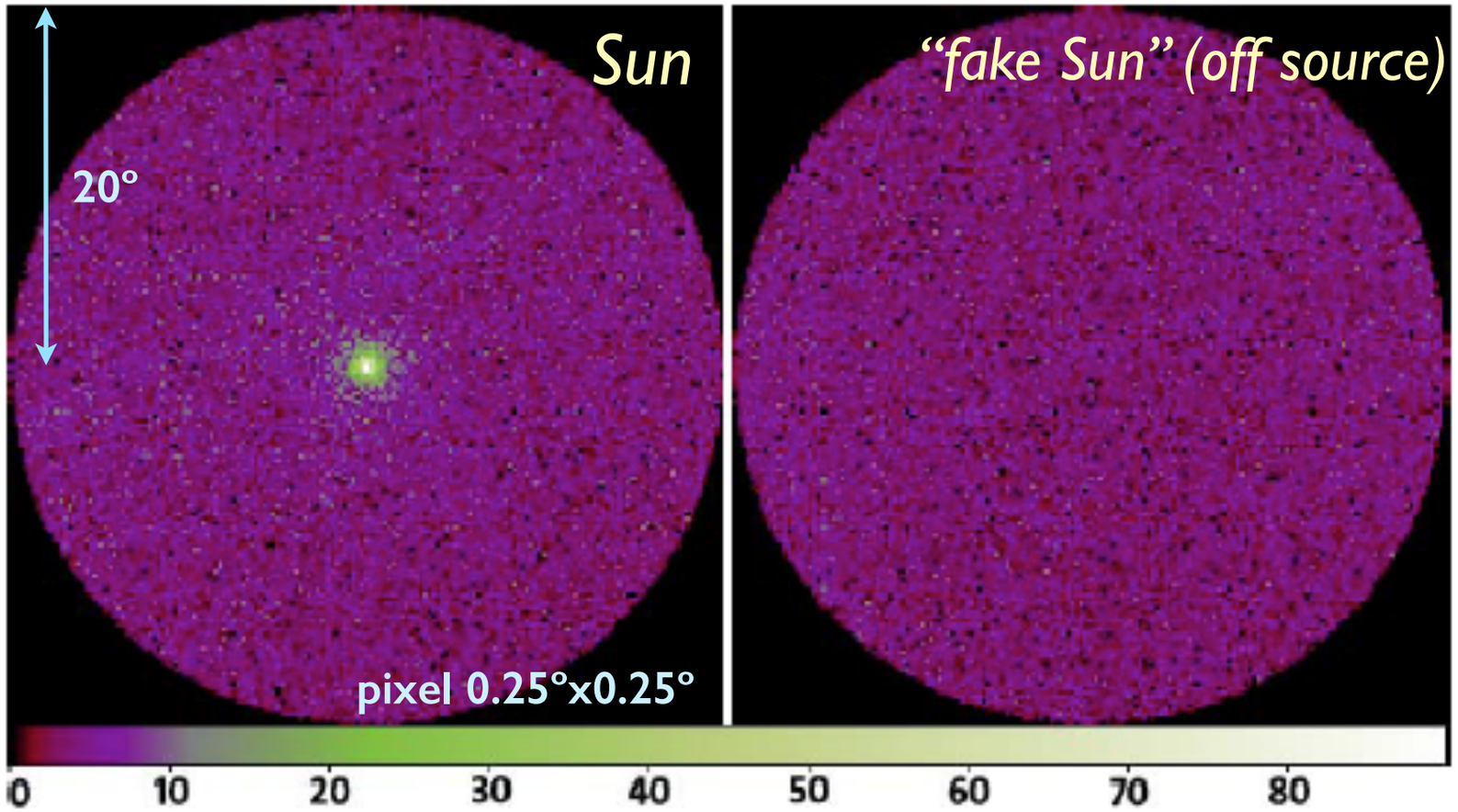
CR protons



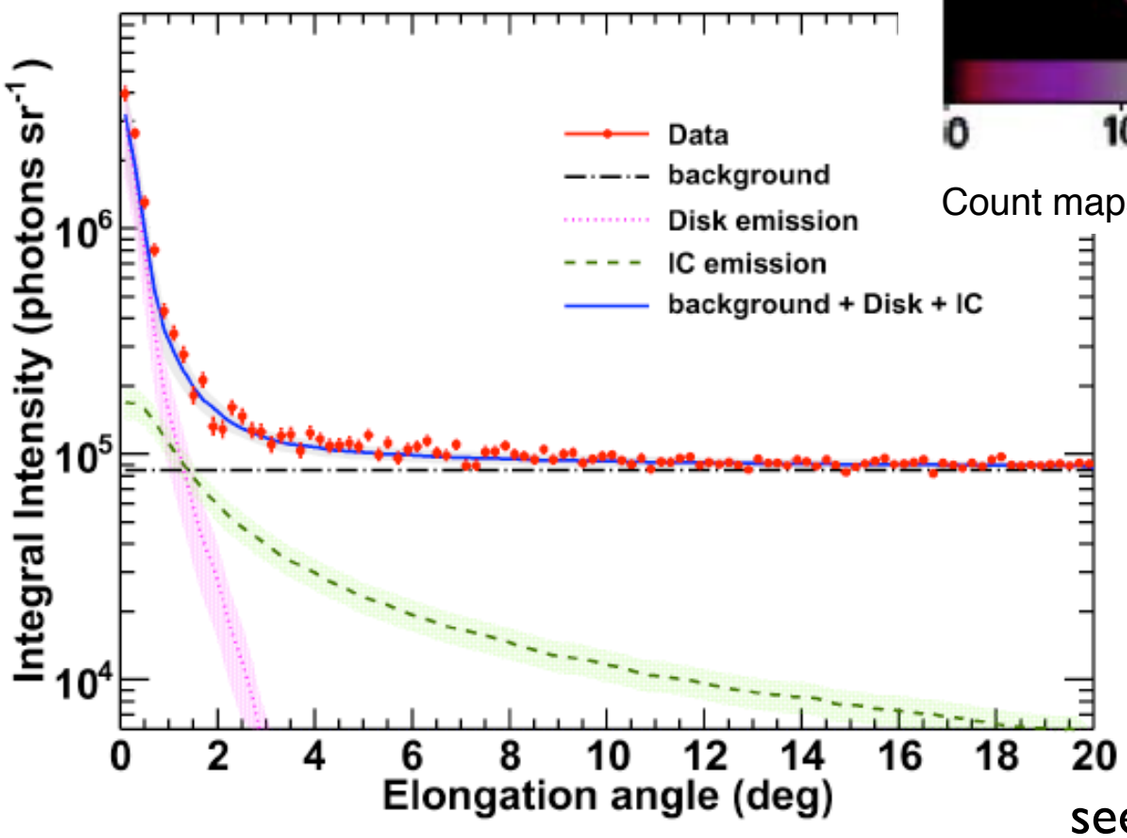


Gamma-ray's from the Sun

- 1.5 yrs of data during solar minimum
  - Aug 2008 - Feb 2010
- Standard Fermi analysis selection criteria



Count maps for events >100 MeV taken between August 2008 and February 2010

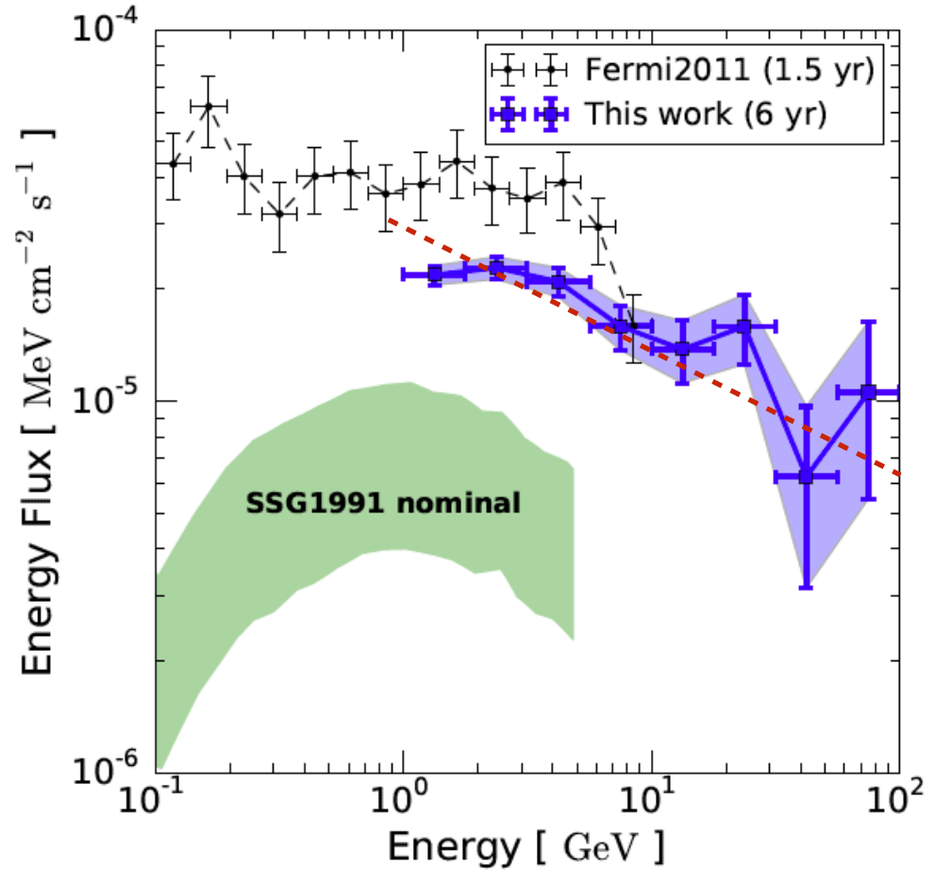


- Extended and disk emission is observed

see Fermi-LAT Collaboration: http://arxiv.org/pdf/1104.2093.pdf

Gamma-ray's from the Sun

- 6 yrs of data
  - Aug 2008 - Aug 2014
- Fermi science tools version v9r33p0



- Observed gamma-ray flux cannot be described by current models
- Significant time variation in solar-disk gamma-rays observed (<10GeV)
- Gamma-ray flux from the Sun extends beyond 100GeV

see K. Ng, J. Beacom, A. Peter, C. Rott PRD 2016

Gamma-ray's from the Sun

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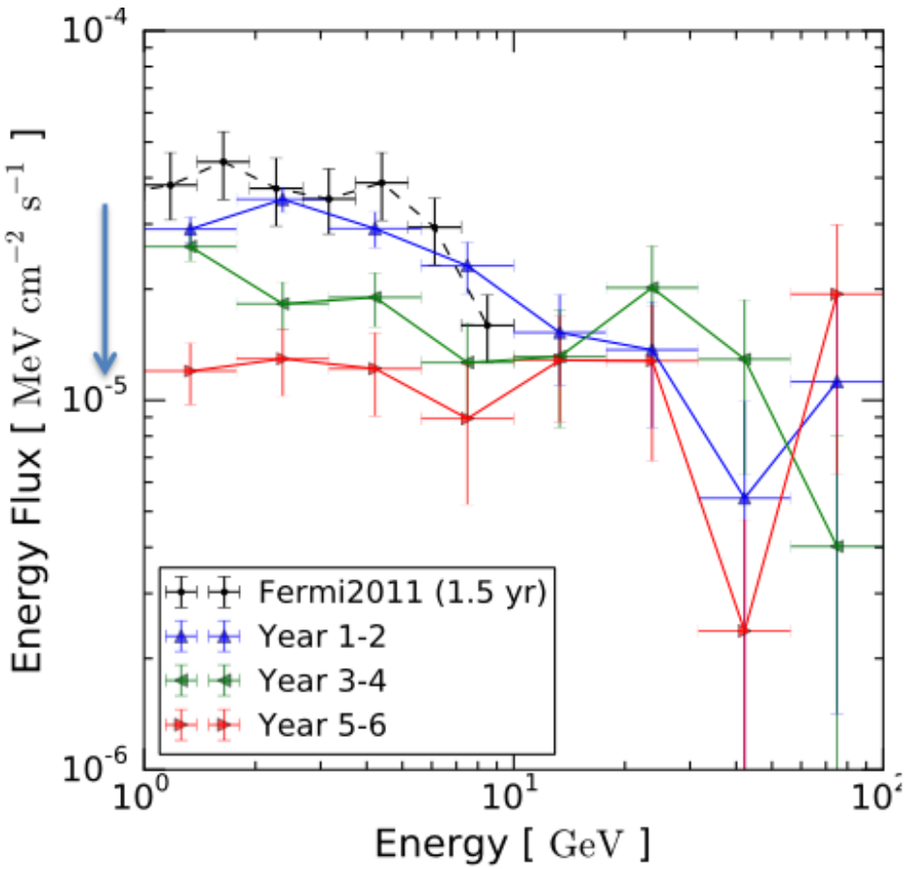

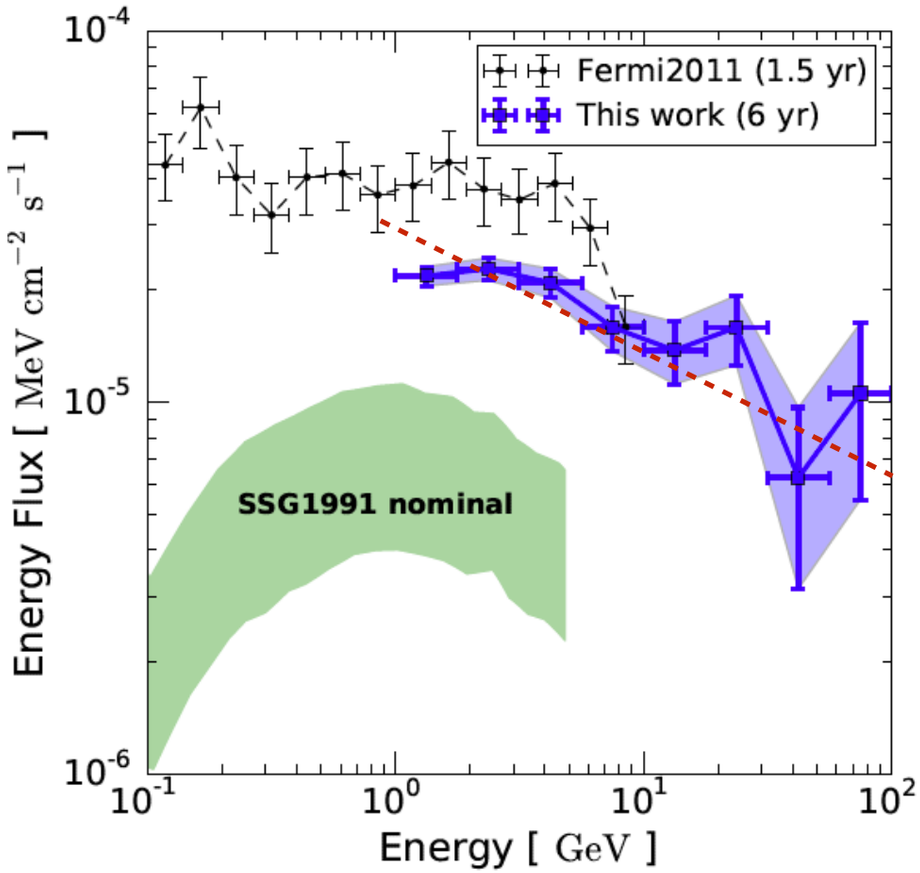

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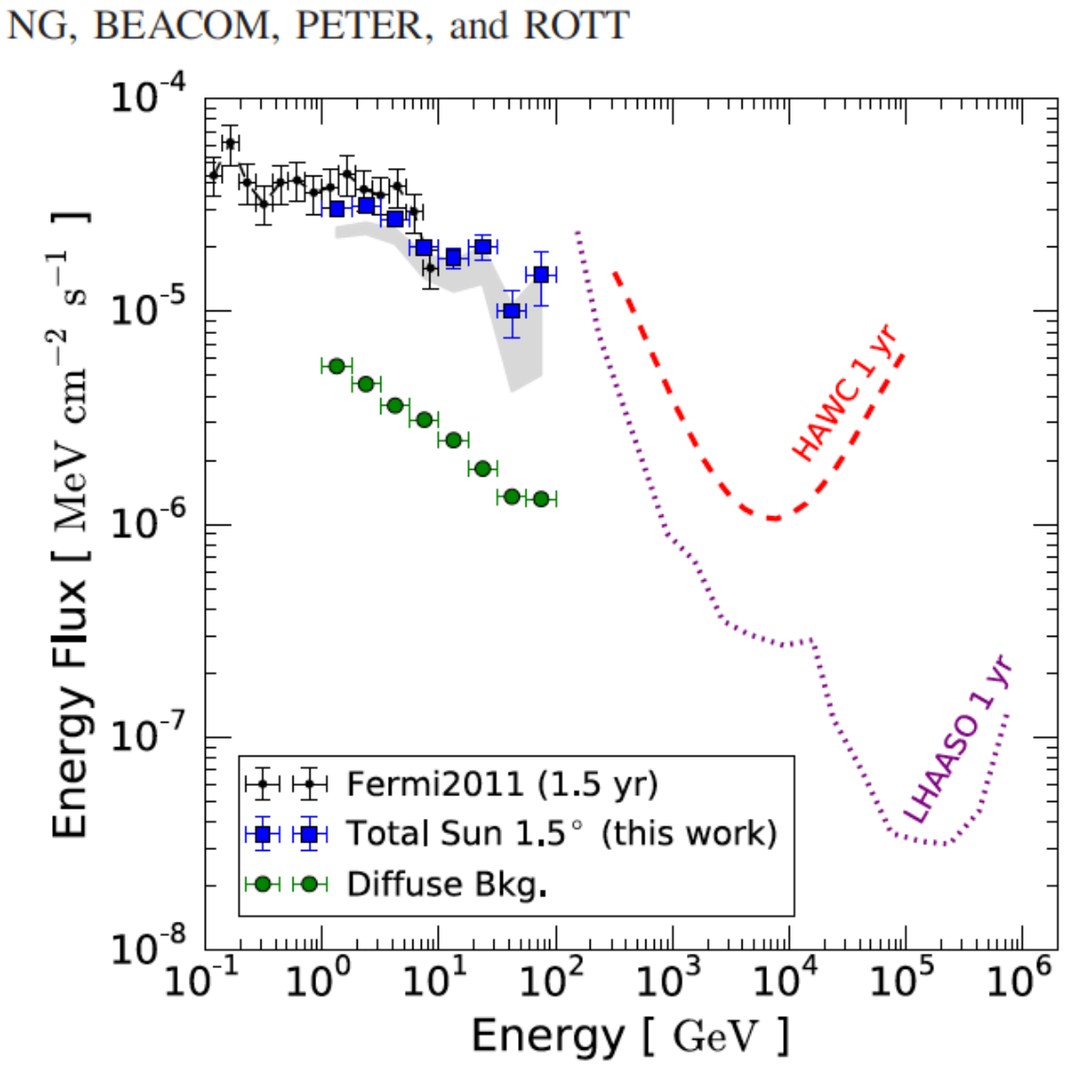
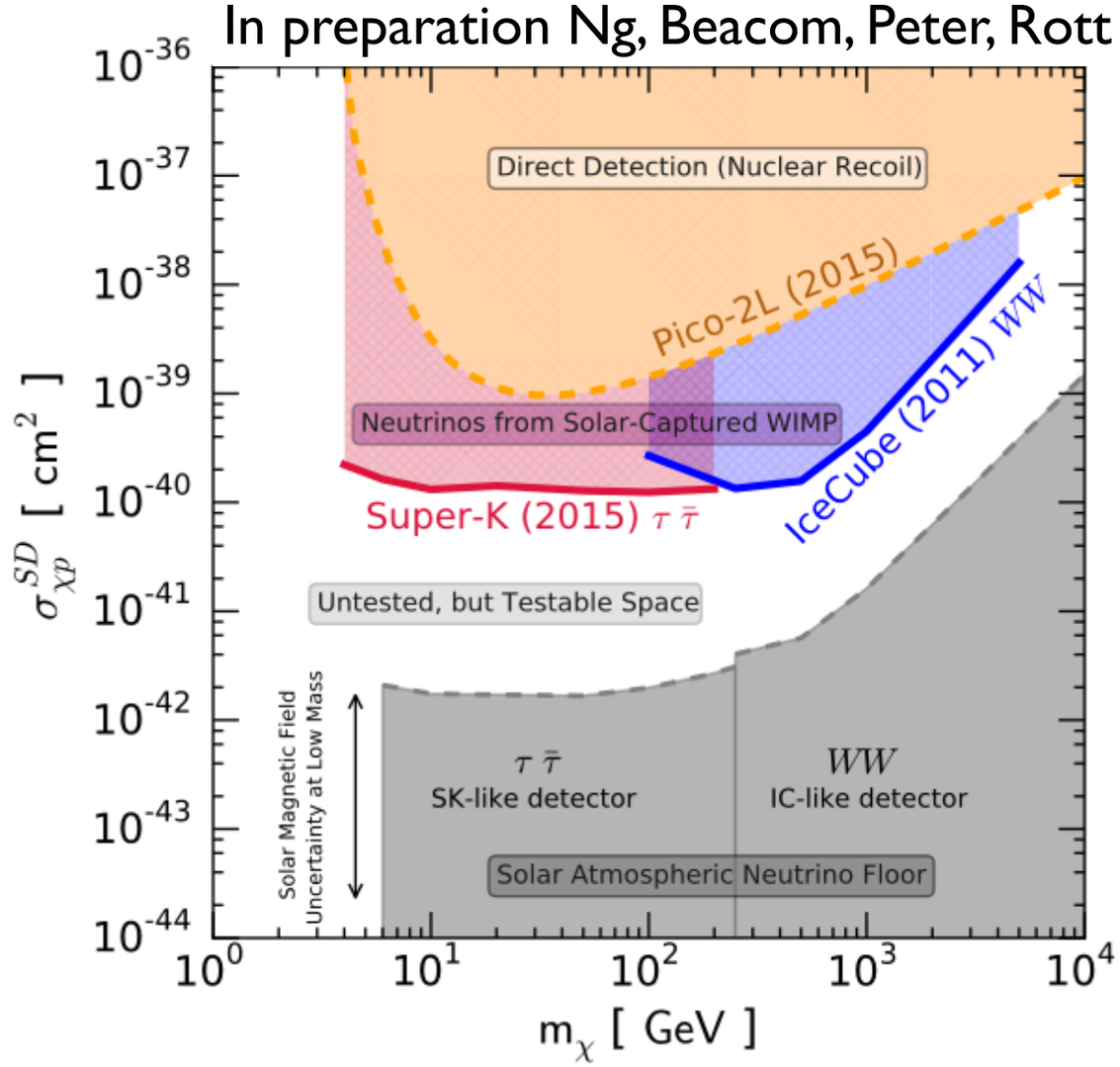
Gamma-ray's from the Sun





- Sun is a promising source for ground-based high altitude water Cherenkov detectors
- Background to dark matter search from the Sun, that soon will be relevant (and first high-energy neutrino point source ??)

see K. Ng, J. Beacom, A. Peter, C. Rott PRD 2016

Low Energy Neutrinos from the Sun

C. Rott, J. Siegal-Gaskins, J.F.Beacom Physical Review D 88, 055005 (2013) (arXiv1208.0827)
C.Rott, S.In, J.Kumar, D.Yaylali JCAP11(2015)039

Low-Energy Neutrinos from the Sun

Possible annihilation channels:

$qq, gg, cc, ss, bb, tt, W^+W^-, ZZ, \tau^+\tau^-, \mu^+\mu^-, \nu\nu, e^+e^-, \gamma\gamma$

few neutrinos

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few neutrinos

some "high energy" neutrinos in decays

⇒ basis of present day searches

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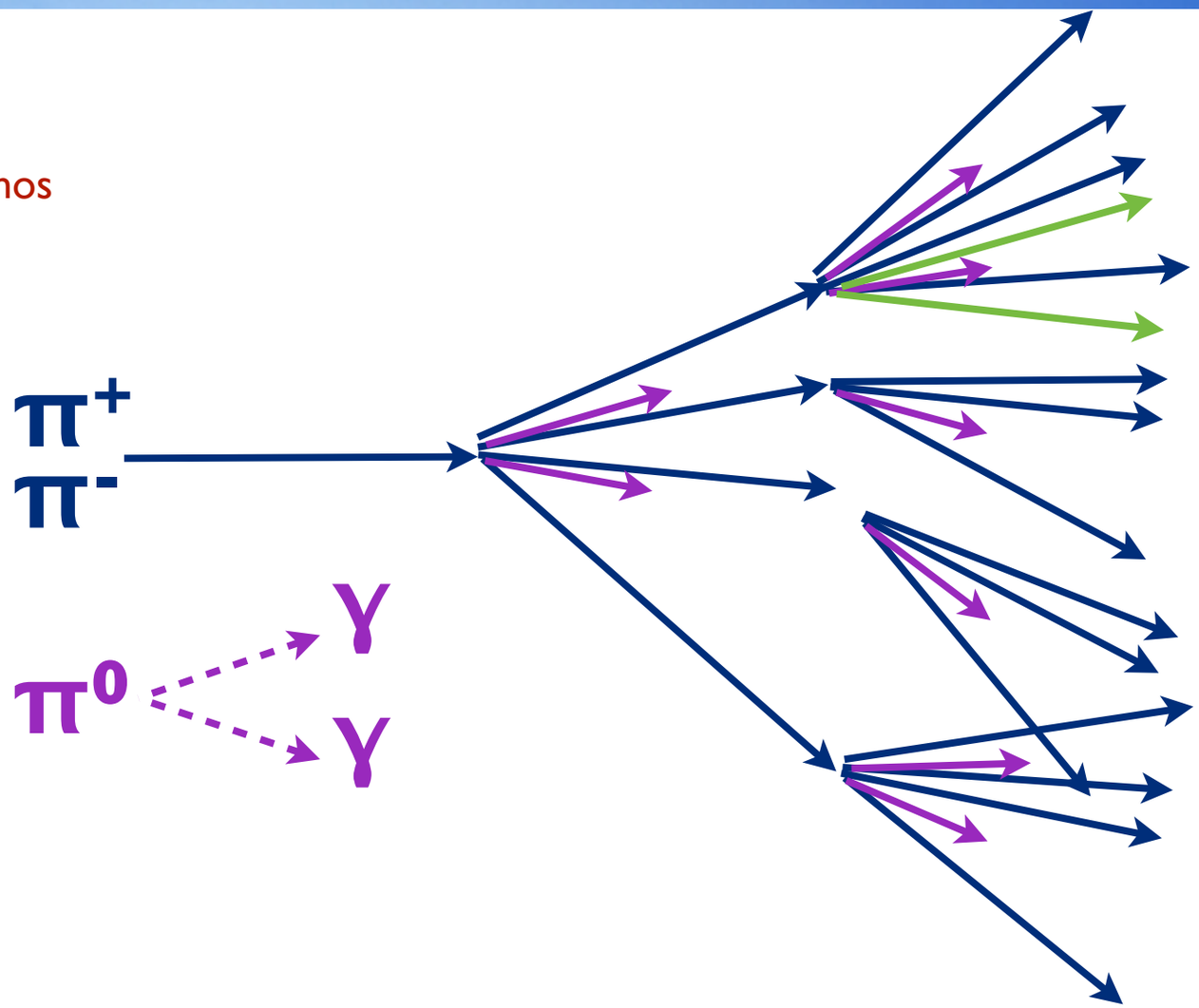

Low-Energy Neutrinos from the Sun

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qq,gg,cc,ss,bb,tt,W^+W^-, ZZ, $\tau^+\tau^-$,$\mu^+\mu^-$, νν, e^+e^-,γγ

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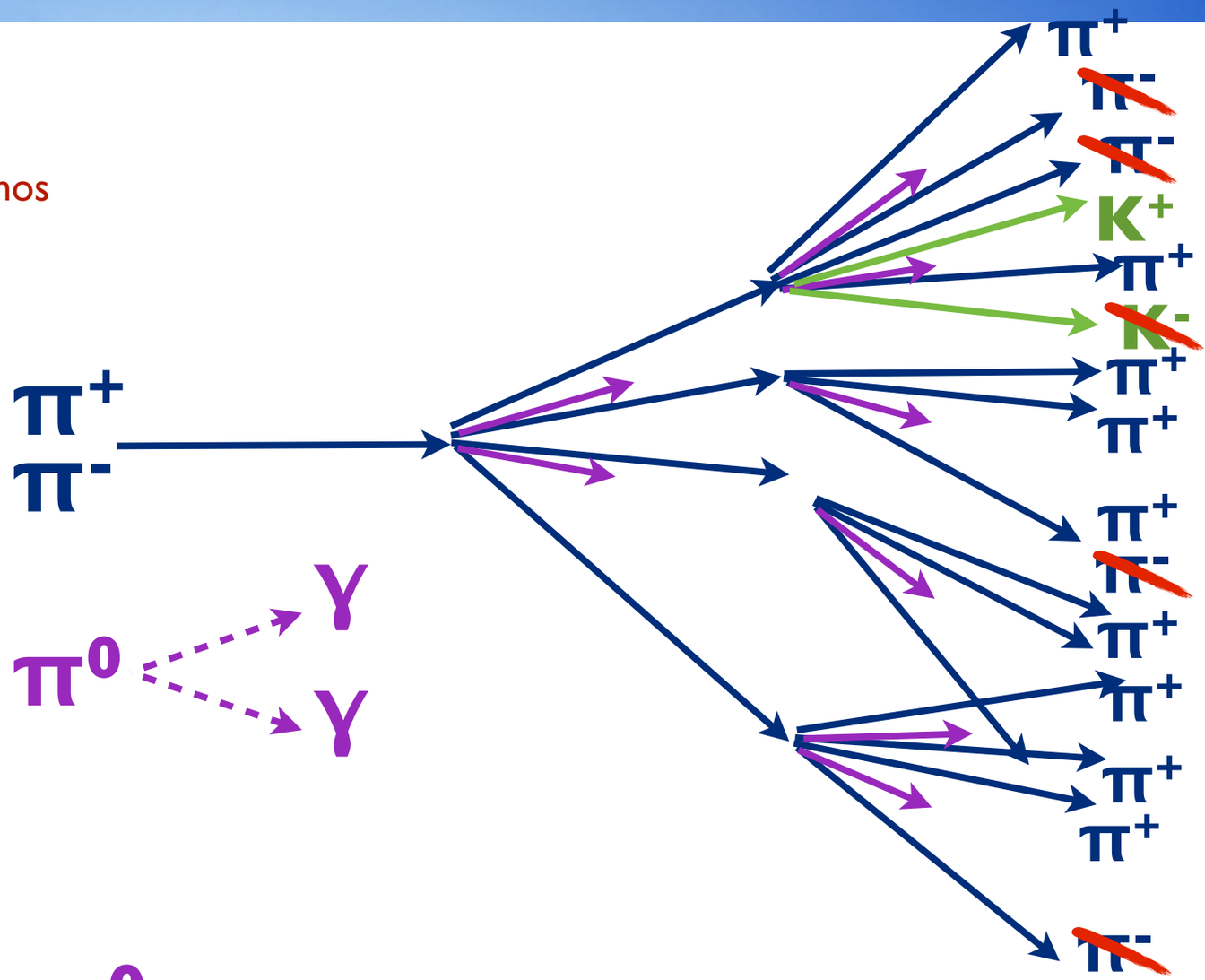

π^0

- Lifetime too short to interact

π^-

- Interaction length short compared to losses
- Produces secondary particles in collision with protons
- Dominant energy loss term is π^0 production

Low-Energy Neutrinos from the Sun

Possible annihilation channels:

qq,gg,cc,ss,bb,tt,W^+W^-, ZZ, $\tau^+\tau^-$,$\mu^+\mu^-$, νν, e^+e^-,γγ

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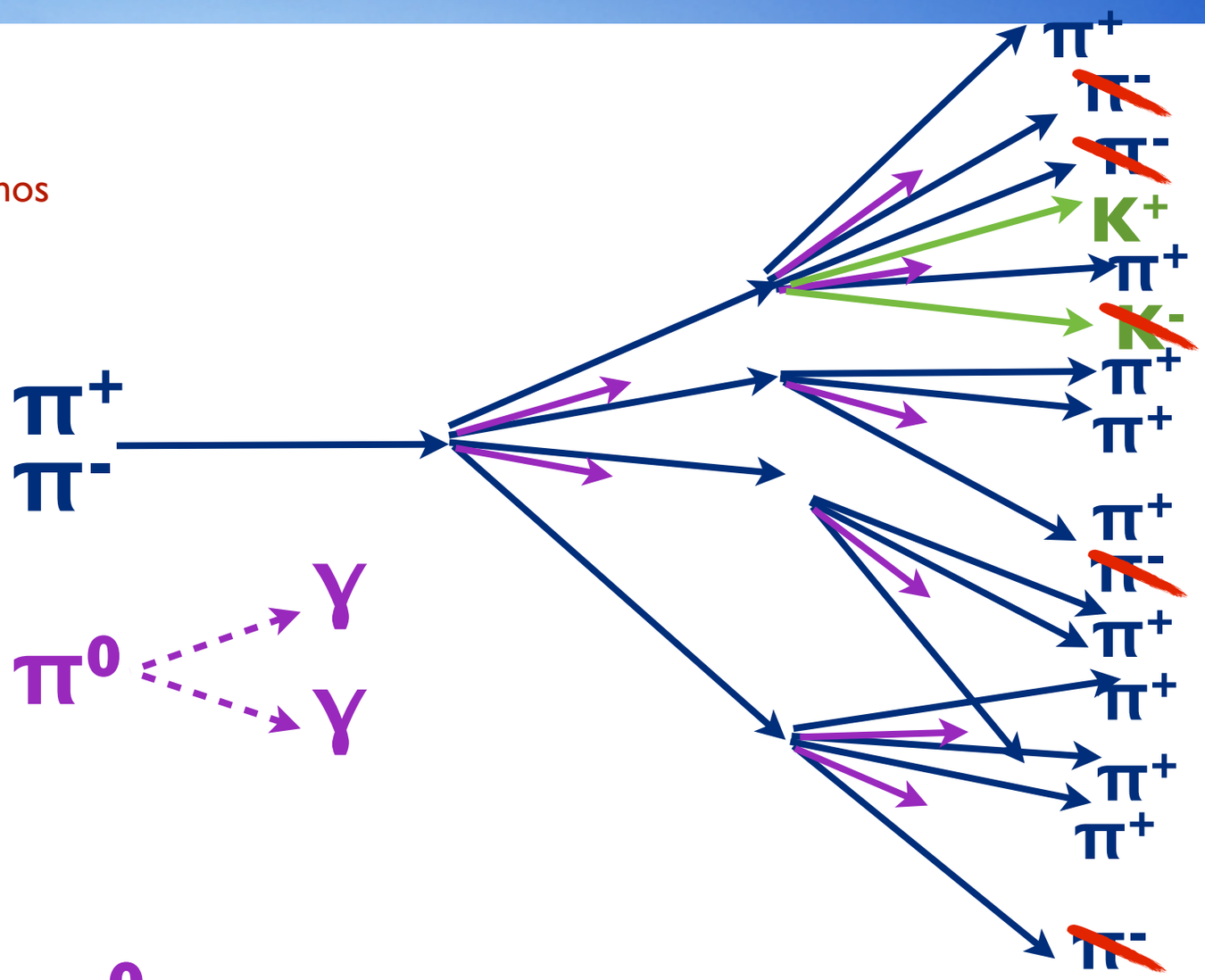


Charged pions and kaons decay at rest producing mono-energetic neutrinos

$\pi^+ \rightarrow \mu^+\nu_\mu$ E_ν= 29.8MeV

$K^+ \rightarrow \nu_\mu \mu^+$ E_ν=235.5MeV

$\mu^+ \rightarrow e^+\nu_e\bar{\nu}_\mu$

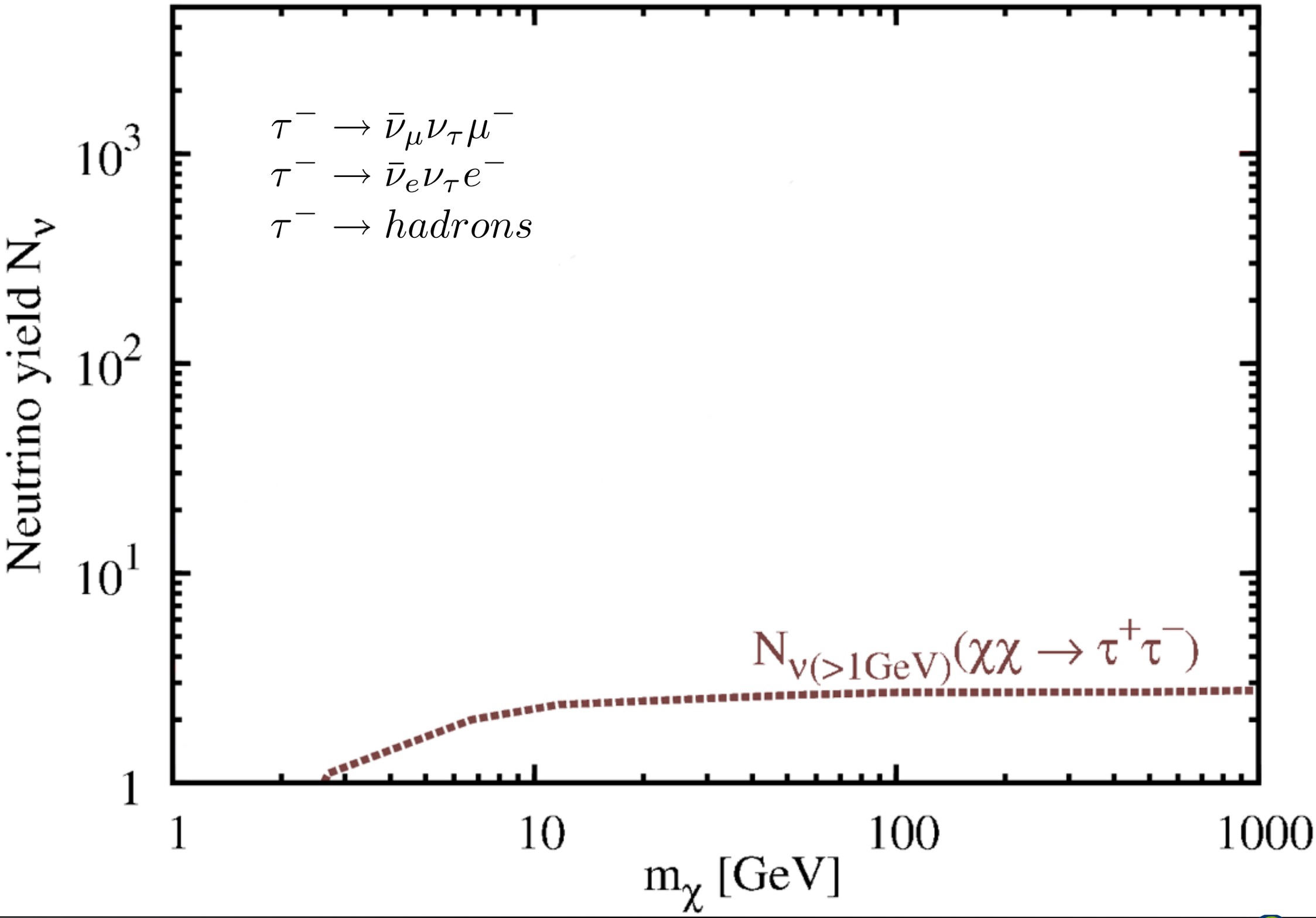
π^0

- Lifetime too short to interact

π^-

- Interaction length short compared to losses
- Produces secondary particles in collision with protons
- Dominant energy loss term is π^0 production

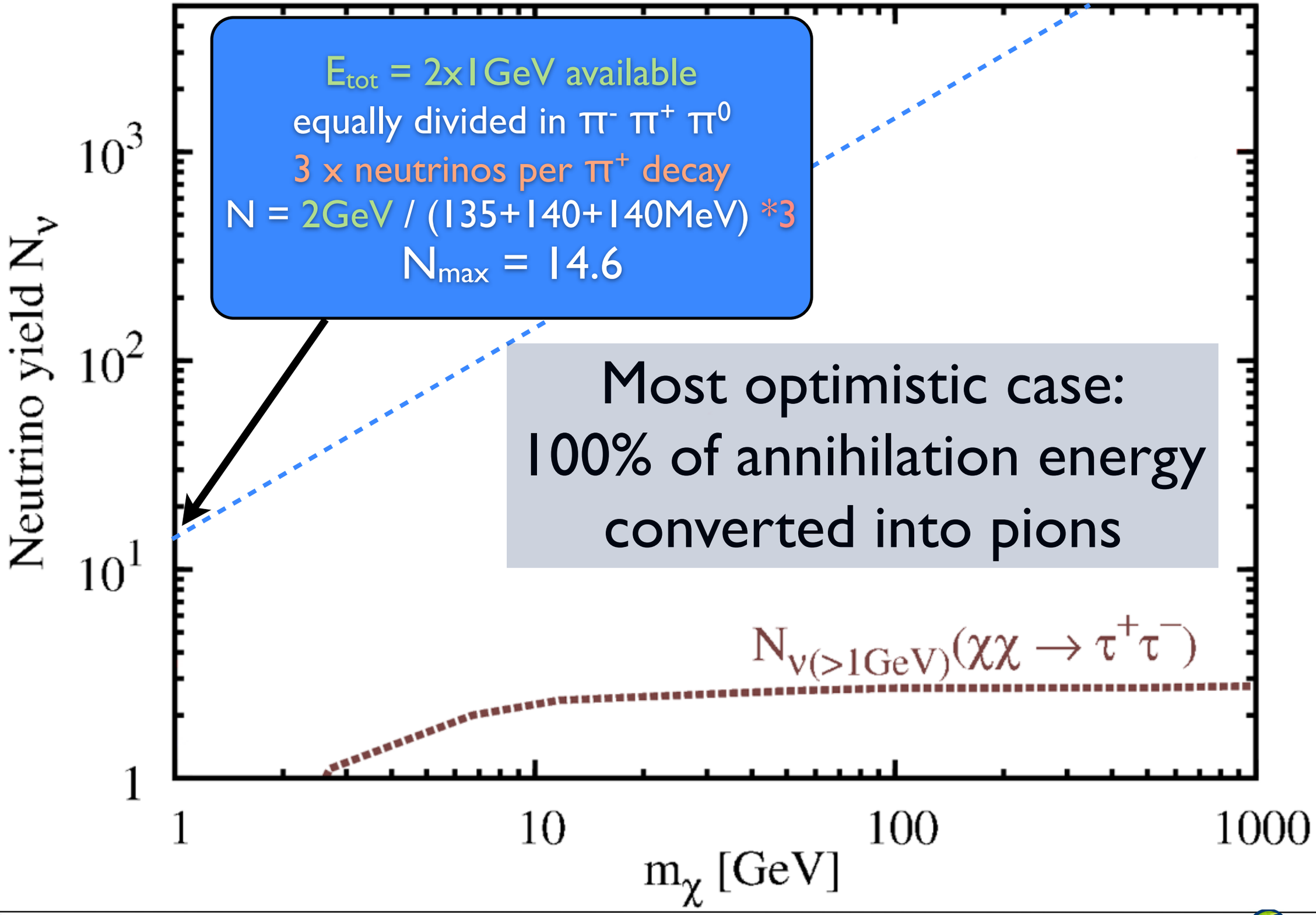
What is the Neutrino yield ?



What's the Neutrino yield ?

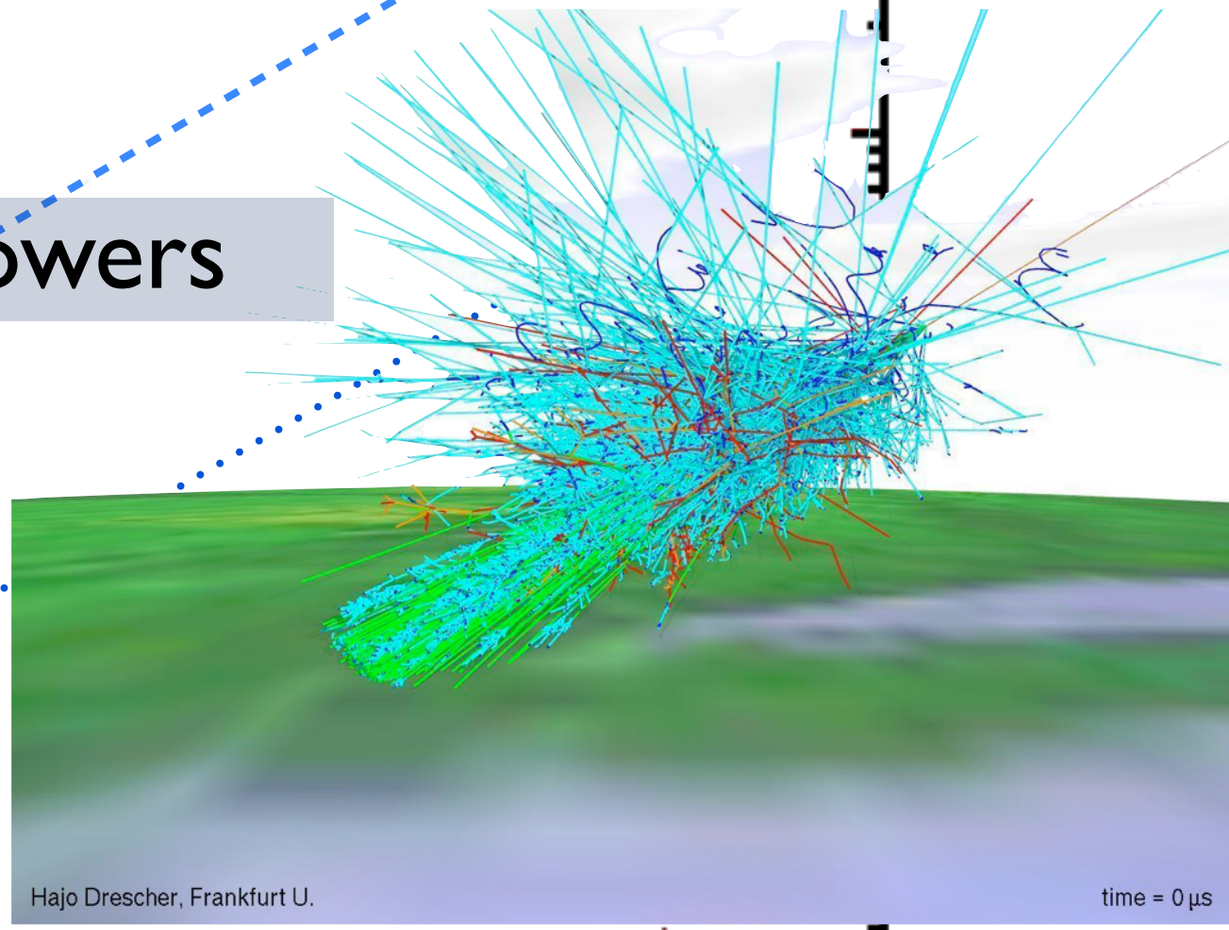
E_{tot} = 2x1GeV available
equally divided in π^- π^+ π^0
3 x neutrinos per π^+ decay
N = 2GeV / (135+140+140MeV) *3
N_{max} = 14.6

Most optimistic case:
100% of annihilation energy
converted into pions



What's the Neutrino yield ?

scaling from airshowers



Neutrino yield N_ν

10^3

10^2

10^1

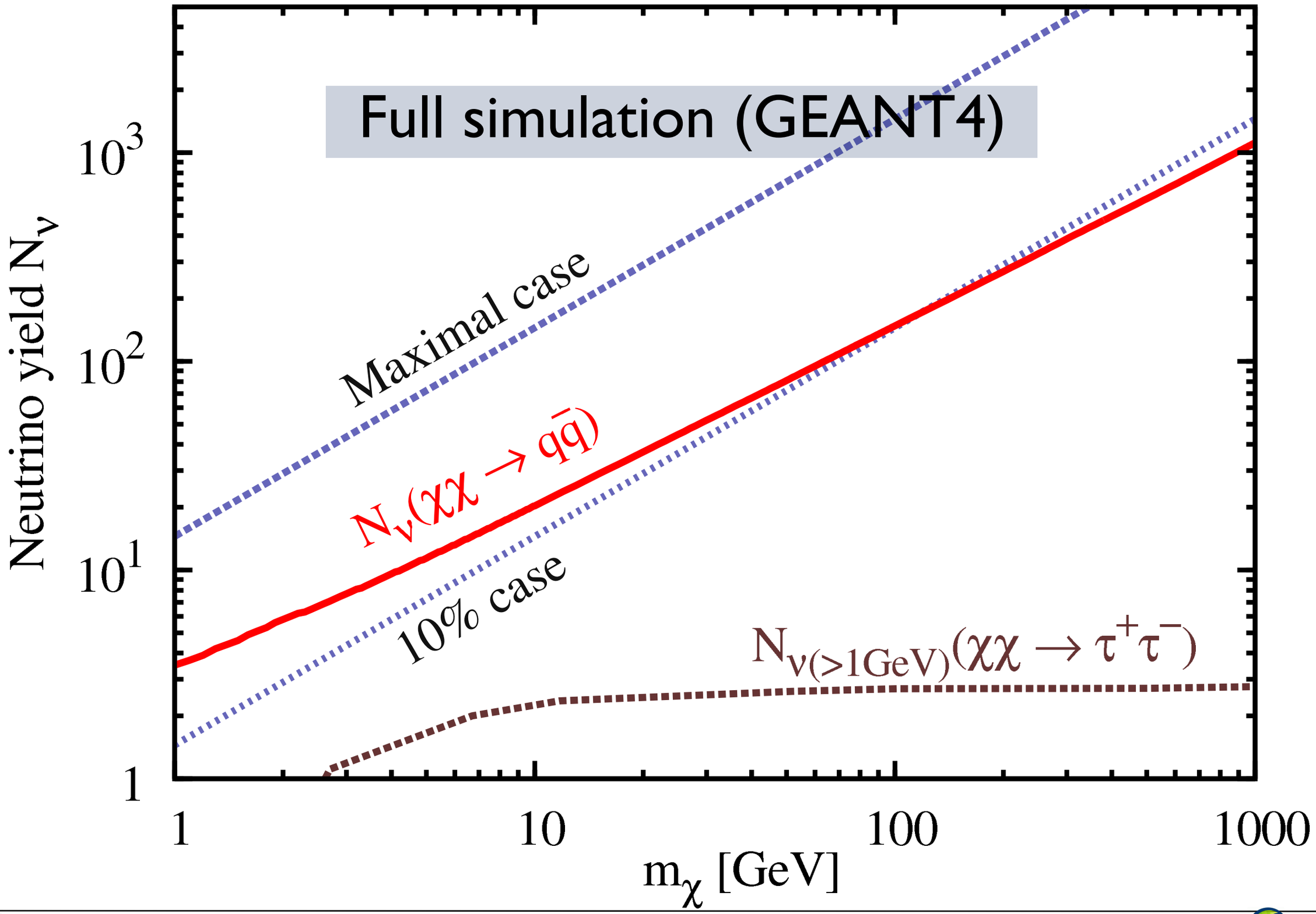
1

$N_{\nu(>1GeV)}(\chi\chi \to \tau^+\tau^-)$

1 10 100 1000

m_χ [GeV]

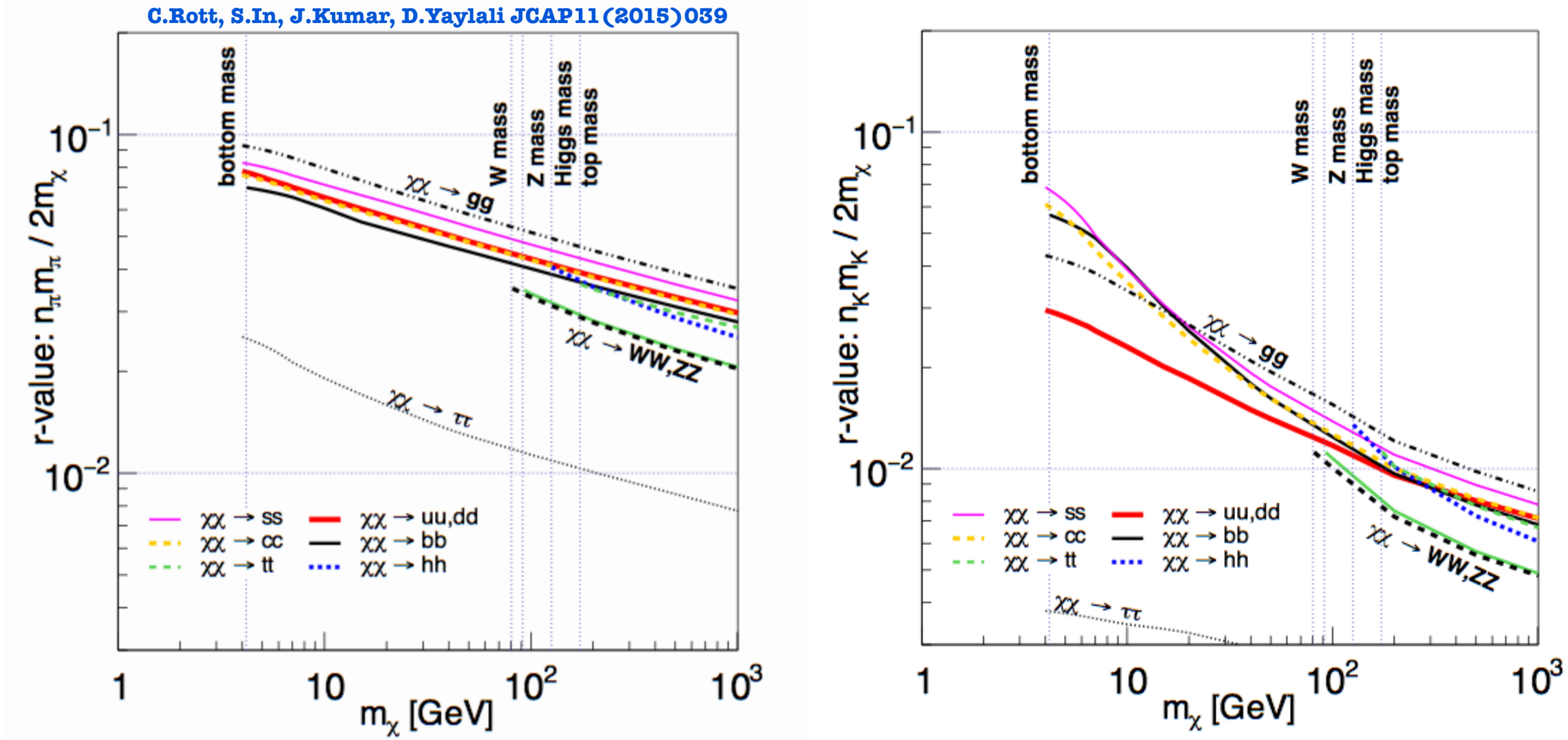
Neutrino yield



Pion and Kaon yields

π^+ r-value - fraction of center-of-mass energy which goes into π^+

K^+ r-value - fraction of center-of-mass energy which goes into K^+

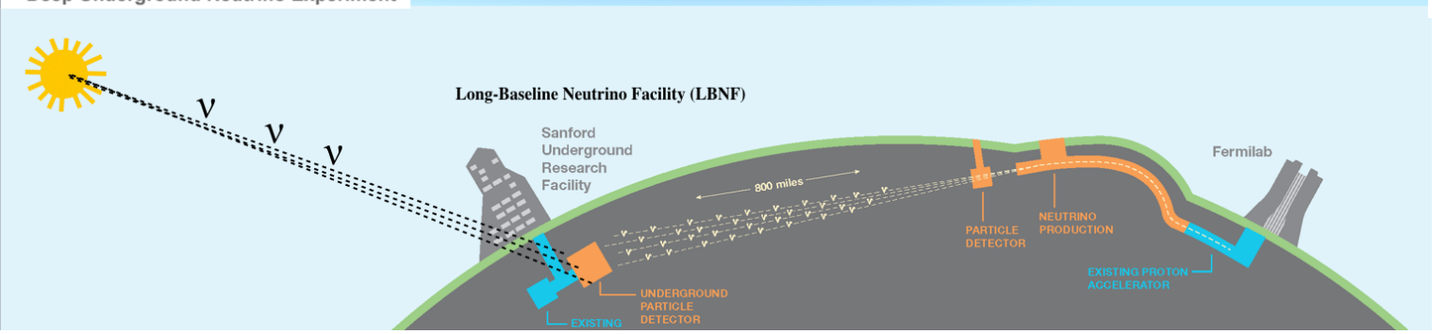


For low dark matter masses difference between flux from stopped pion and kaon decay at rest can be used to disentangle annihilation final states

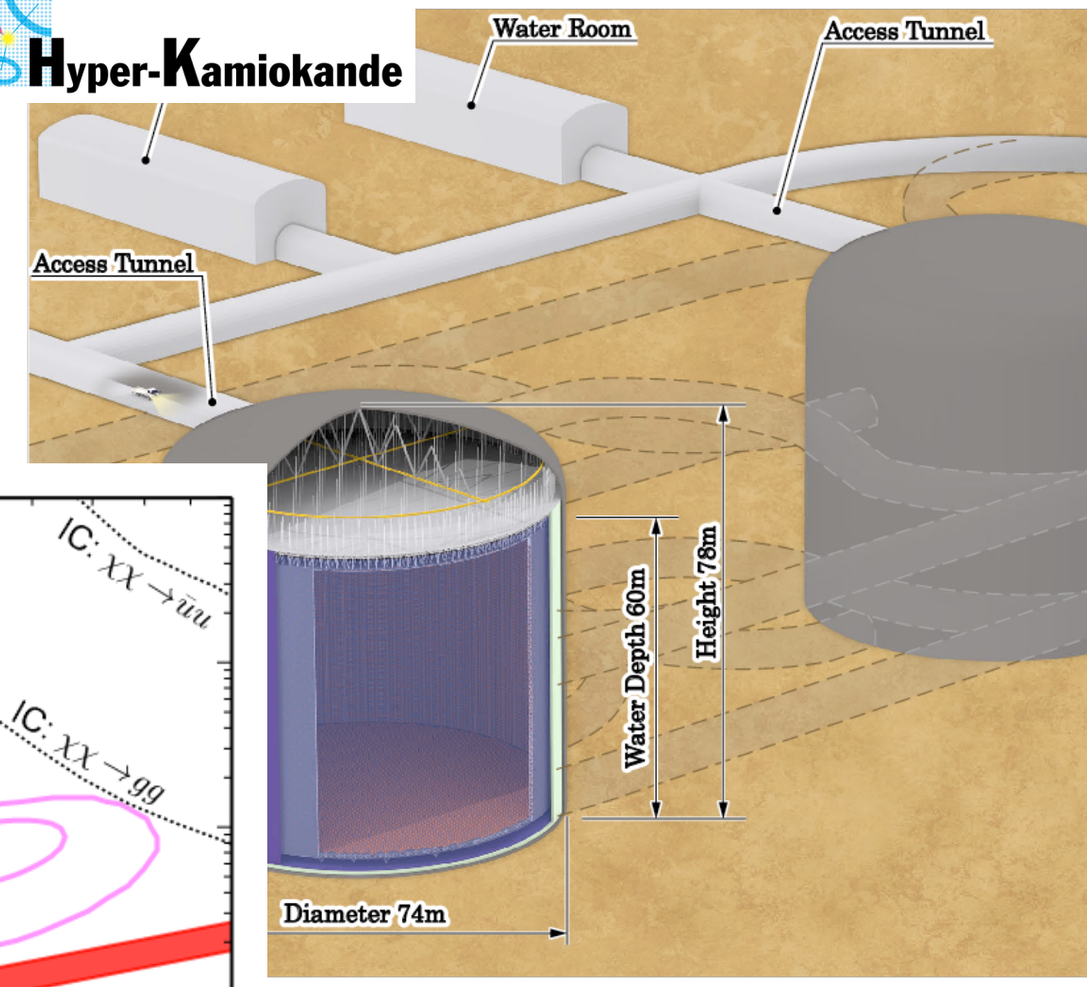
Sensitivity

DUNE
Deep Underground Neutrino Experiment

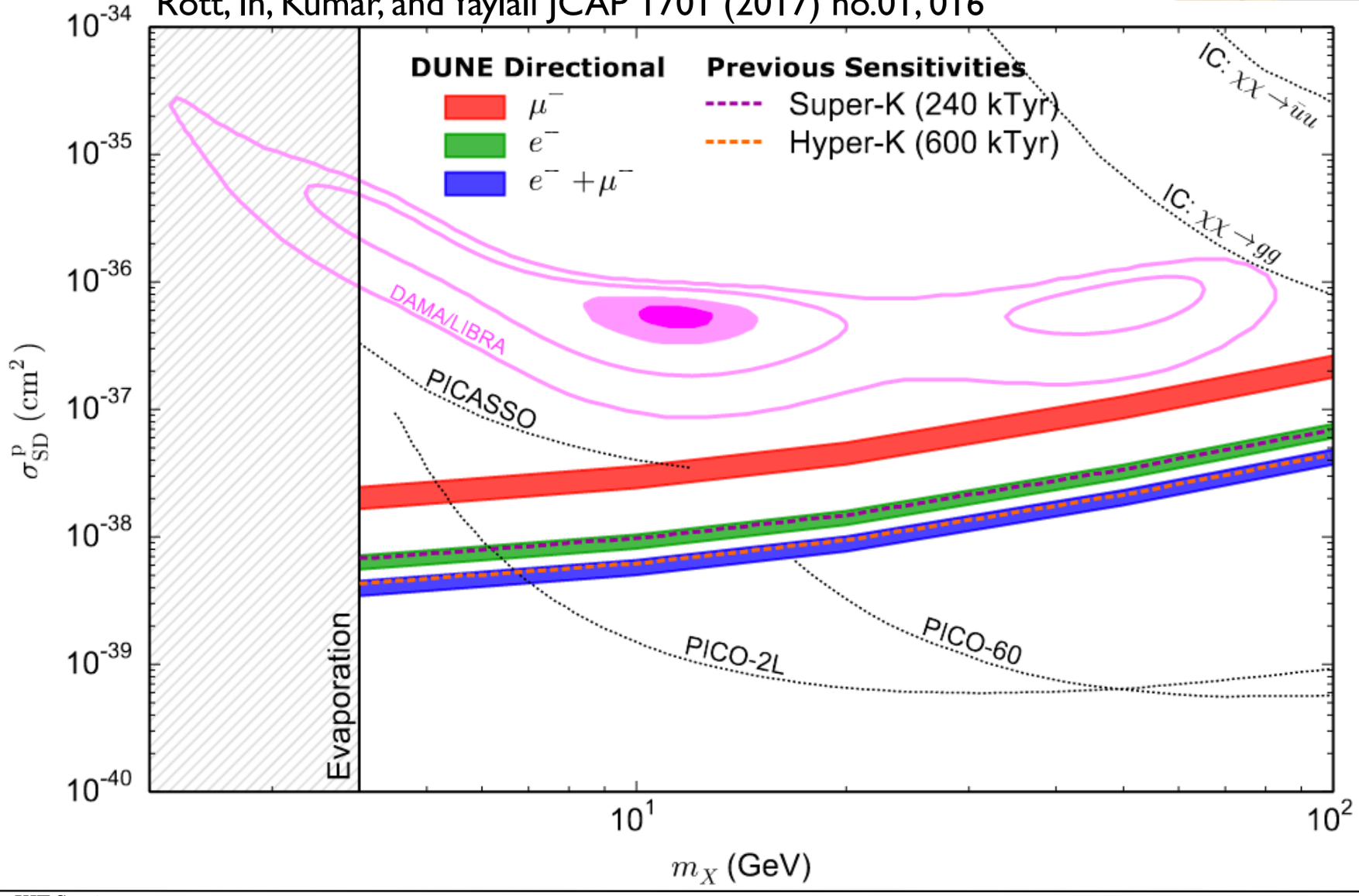
http://www.dunescience.org/

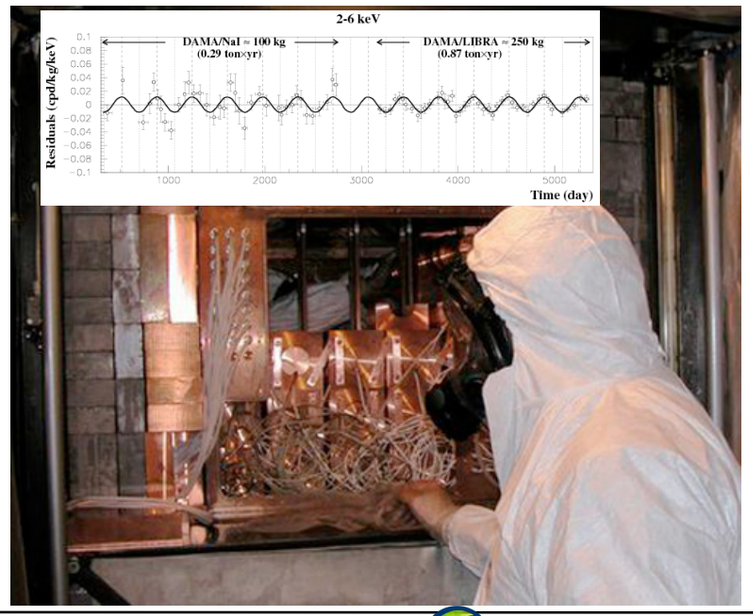


Hyper-Kamiokande



Rott, In, Kumar, and Yaylali JCAP 1701 (2017) no.01, 016

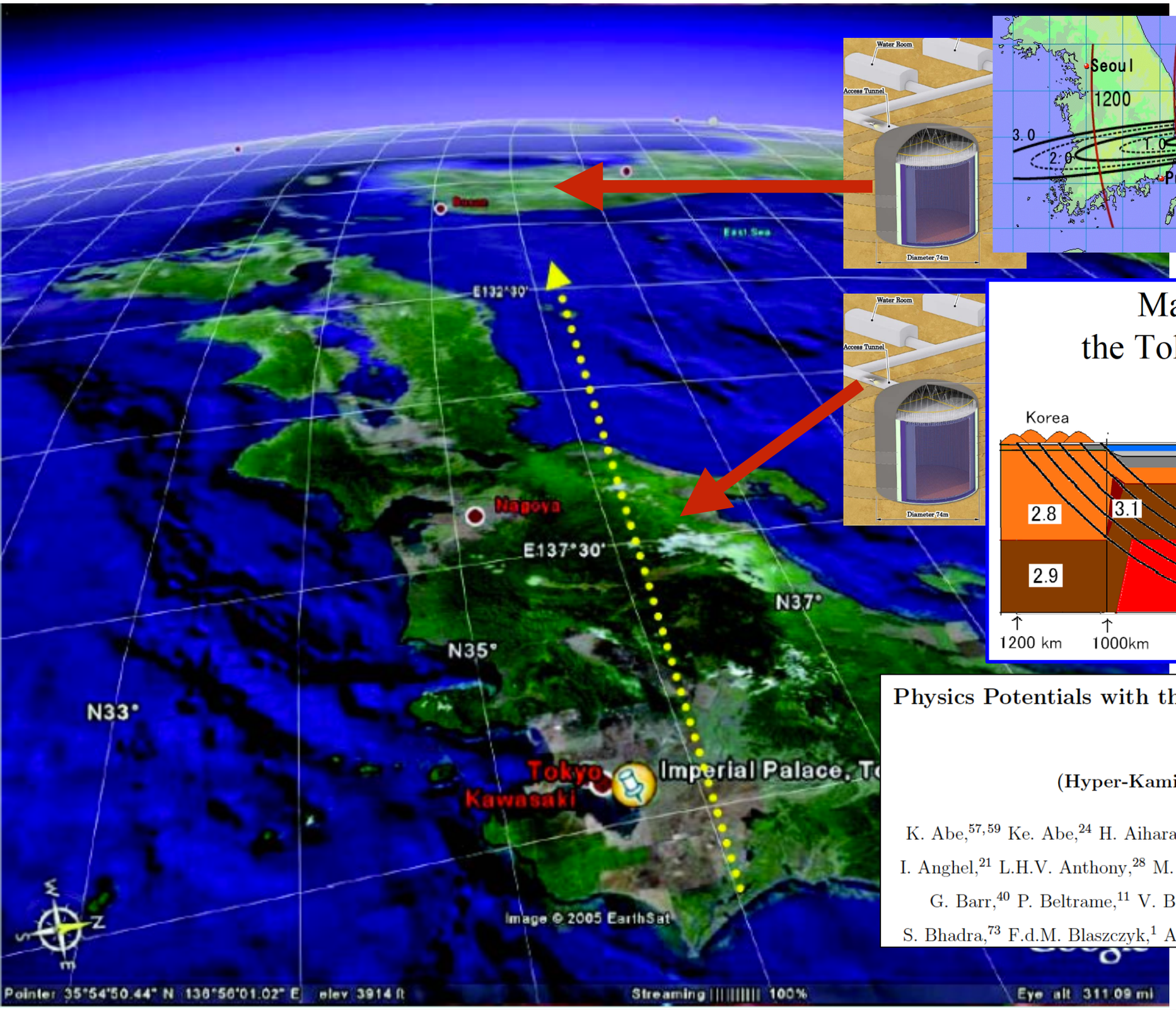


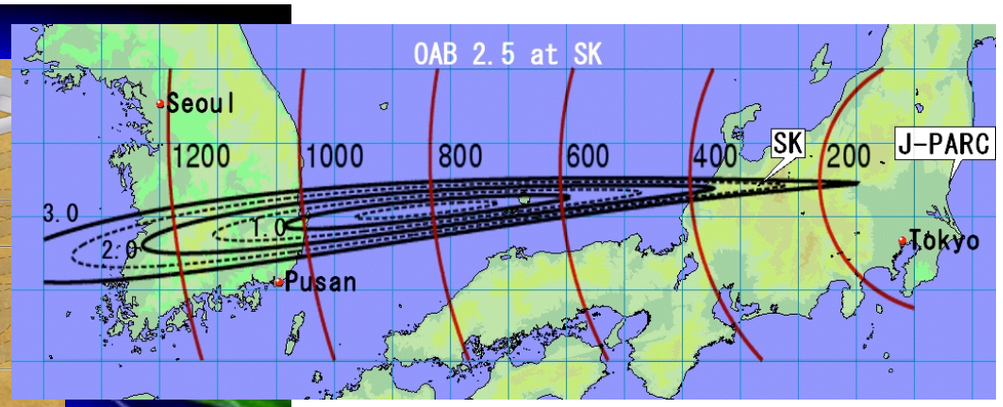


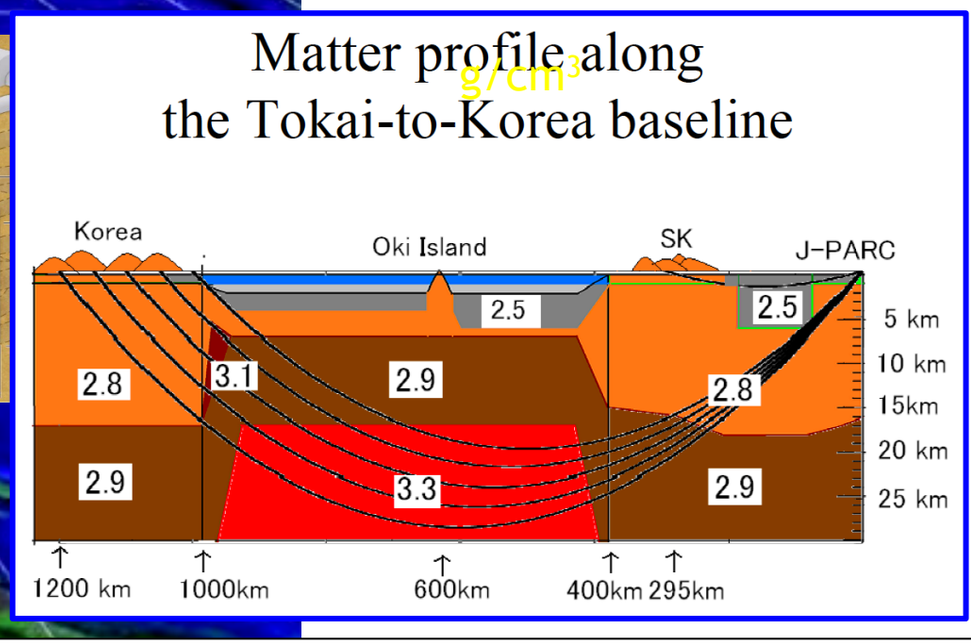
2nd Hyper-K Detector in Korea ?

Hyper-Kamiokande Proto-Collaboration
arXiv:1611.06118

Tokai-to-Hyper-K & Korea (T2HKK)







Physics Potentials with the Second Hyper-Kamiokande Detector in Korea

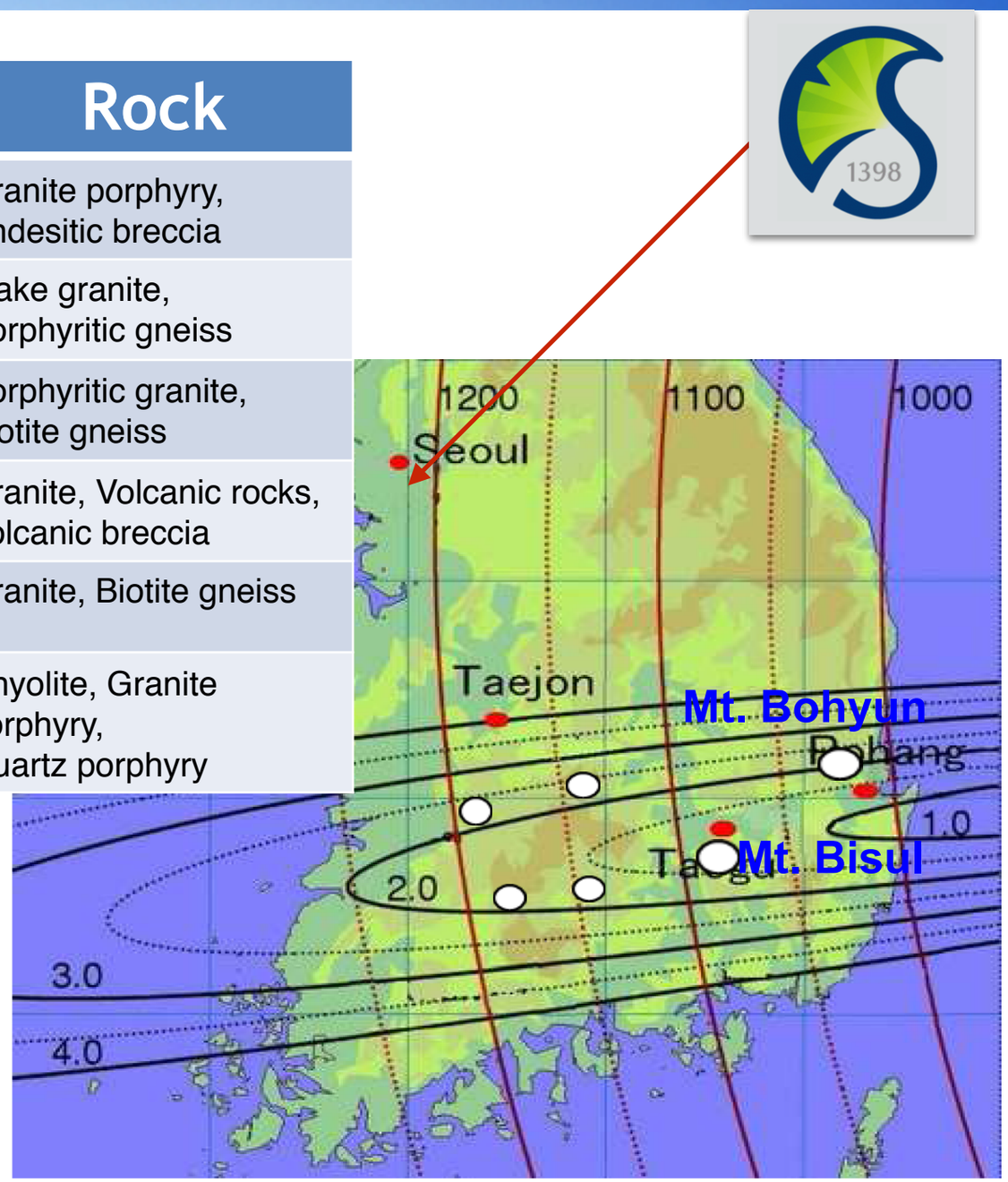
(Hyper-Kamiokande Proto-Collaboration)

K. Abe,[57,59] Ke. Abe,[24] H. Aihara,[59,60] A. Aimi,[18] R. Akutsu,[58] C. Andreopoulos,[28,43] I. Anghel,[21] L.H.V. Anthony,[28] M. Antonova,[20] Y. Ashida,[25] M. Barbi,[44] G.J. Barker,[66] G. Barr,[40] P. Beltrame,[11] V. Berardi,[16] M. Bergevin,[3] S. Berkman,[2] T. Berry,[45] S. Bhadra,[73] F.d.M. Blaszczyk,[1] A. Blondel,[12] S. Bolognesi,[6] S.B. Boyd,[66] A. Bravar,[12]

arXiv:1611.06118

Candidate Sites in Korea

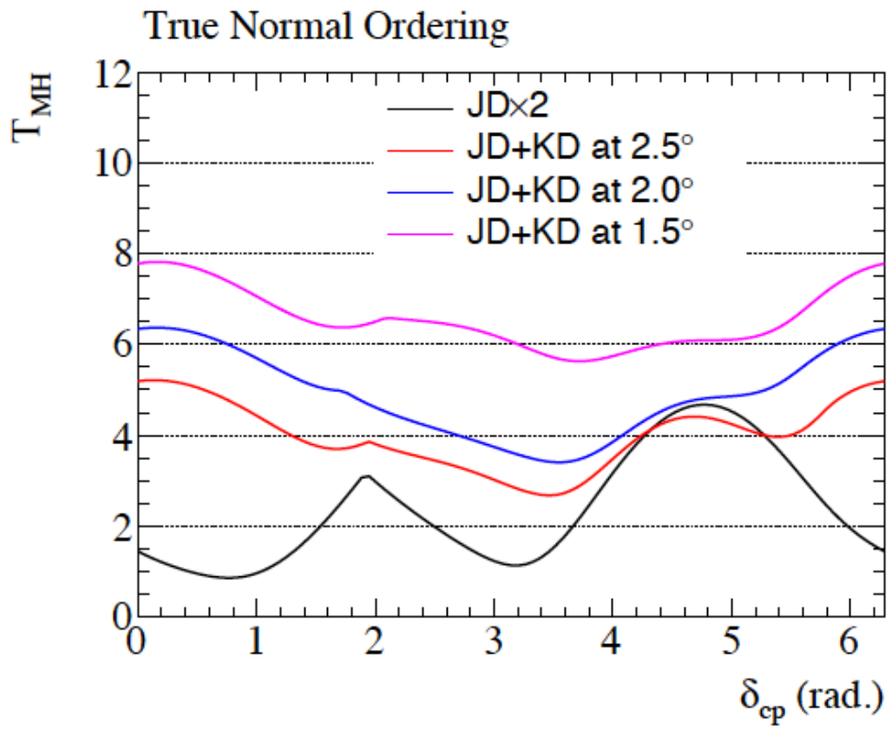
Site	OAB	Baseline	Height	Rock
Mt. Bisul	~1.3°	1088 km	1084 m	Granite porphyry, Andesitic breccia
Mt. Hwangmae	~1.8°	1140 km	1113 m	Flake granite, Porphyritic gneiss
Mt. Sambong	~1.9°	1180 km	1186 m	Porphyritic granite, Biotite gneiss
Mt. Bohyun	~2.2°	1040 km	1126 m	Granite, Volcanic rocks, Volcanic breccia
Mt. Minjuii	~2.2°	1140 km	1242 m	Granite, Biotite gneiss
Mt. Unjang	~2.2°	1190 km	1125 m	Rhyolite, Granite porphyry, Quartz porphyry

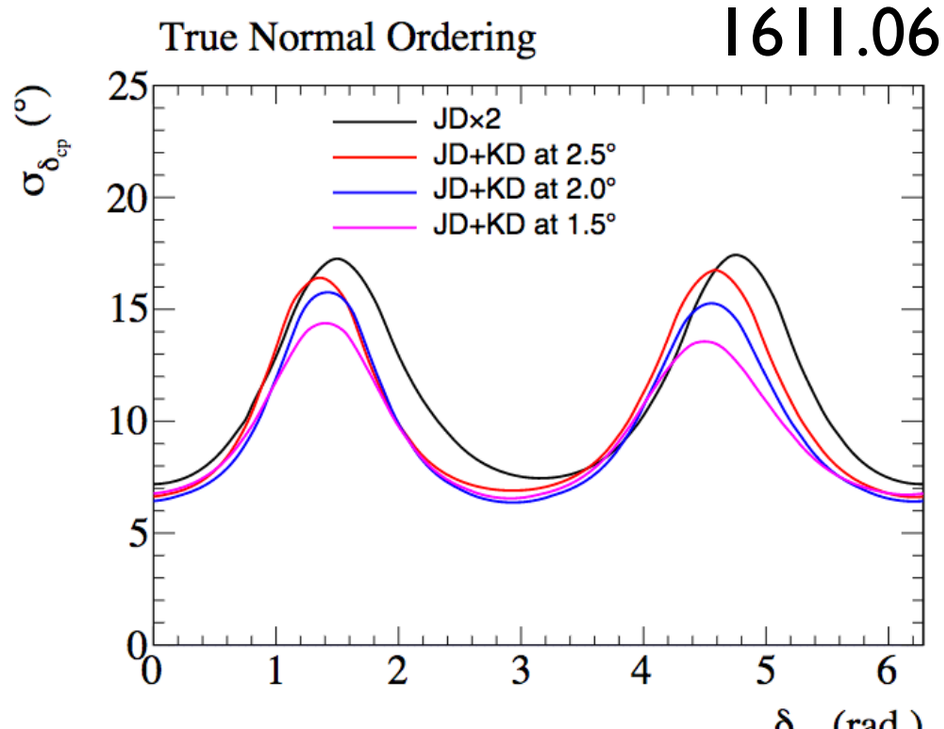


- Baselines length 1,000 ~ 1,200 km
- Off axis angle 1.3° ~ 3°
- Considering tunnel entrance positions overburdens are expected to be greater than 820 m (2,200 m.w.e.)

arXiv:1611.06118

Tokai-to-Hyper-K & Korea (T2HKK)





1611.06118

- Improved CP Precision, Mass hierarchy, …
- Better control of systematics
- Potential site benefits (larger over burden)
- Non-standard neutrino interactions

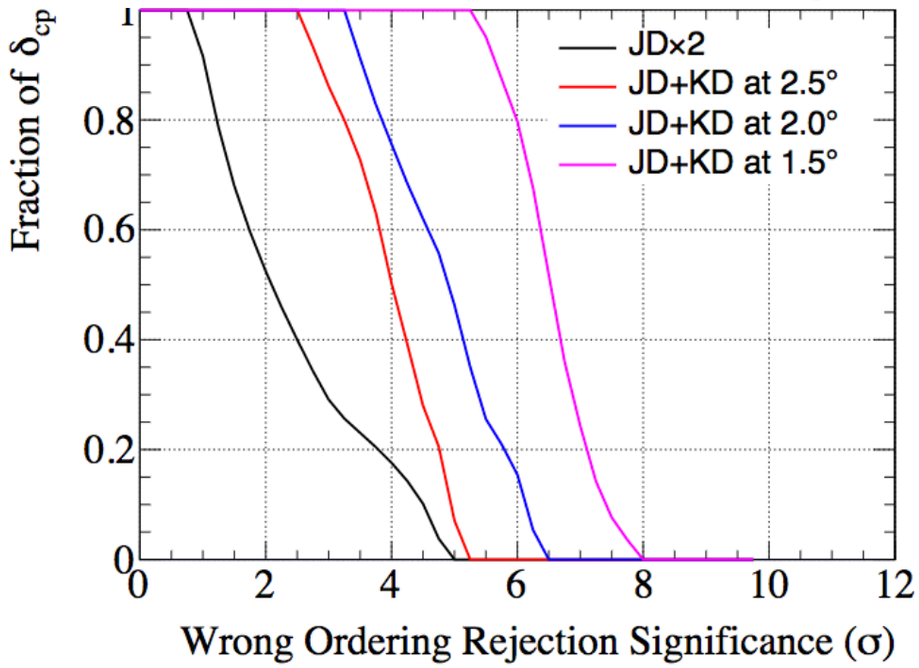
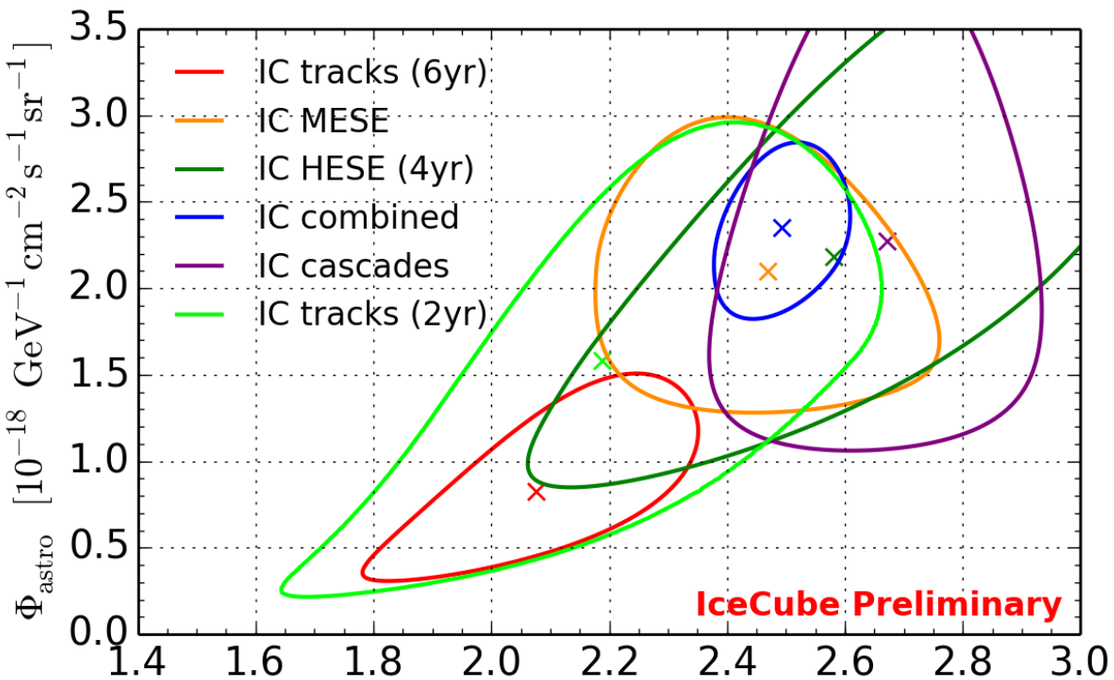
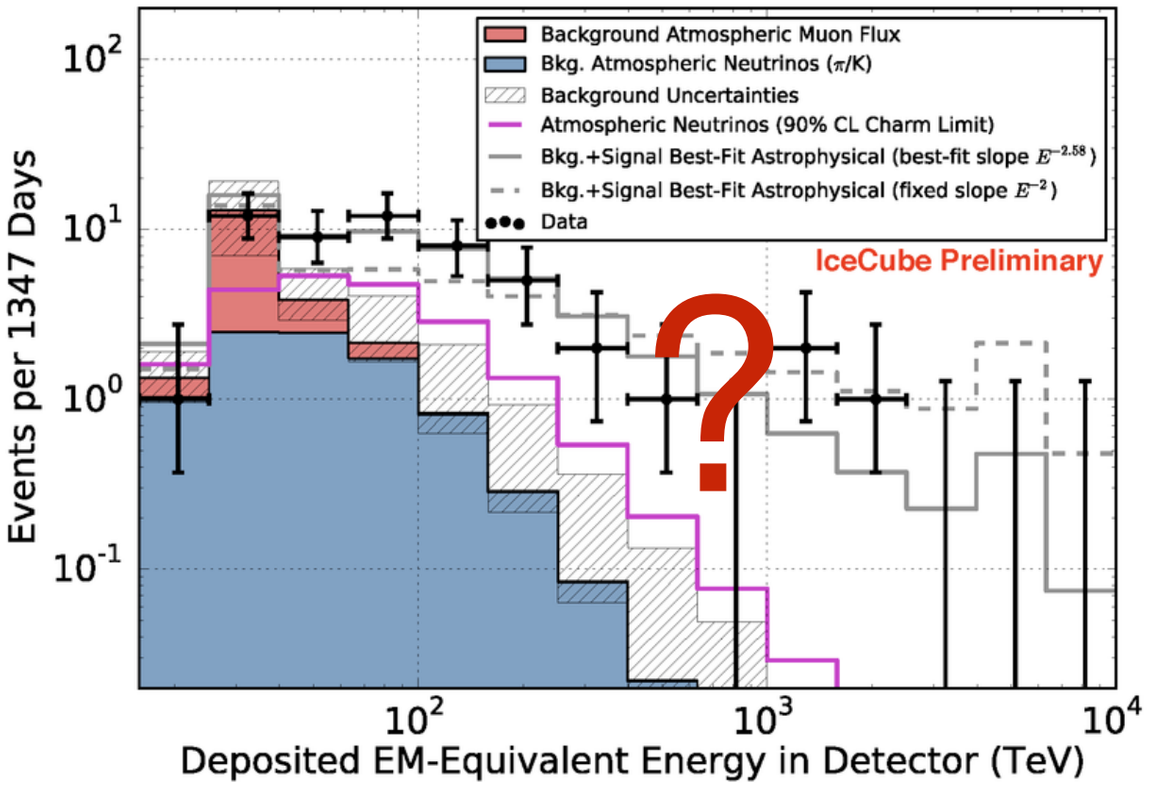


FIG. 19: The fraction of δ_{cp} values (averaging over the true mass ordering) for which the wrong hierarchy can be rejected with a given significance or greater.

Future Plans for IceCube …

Beyond Standard Model Physics at the PeV scale





Intense interest in high-energy neutrino region

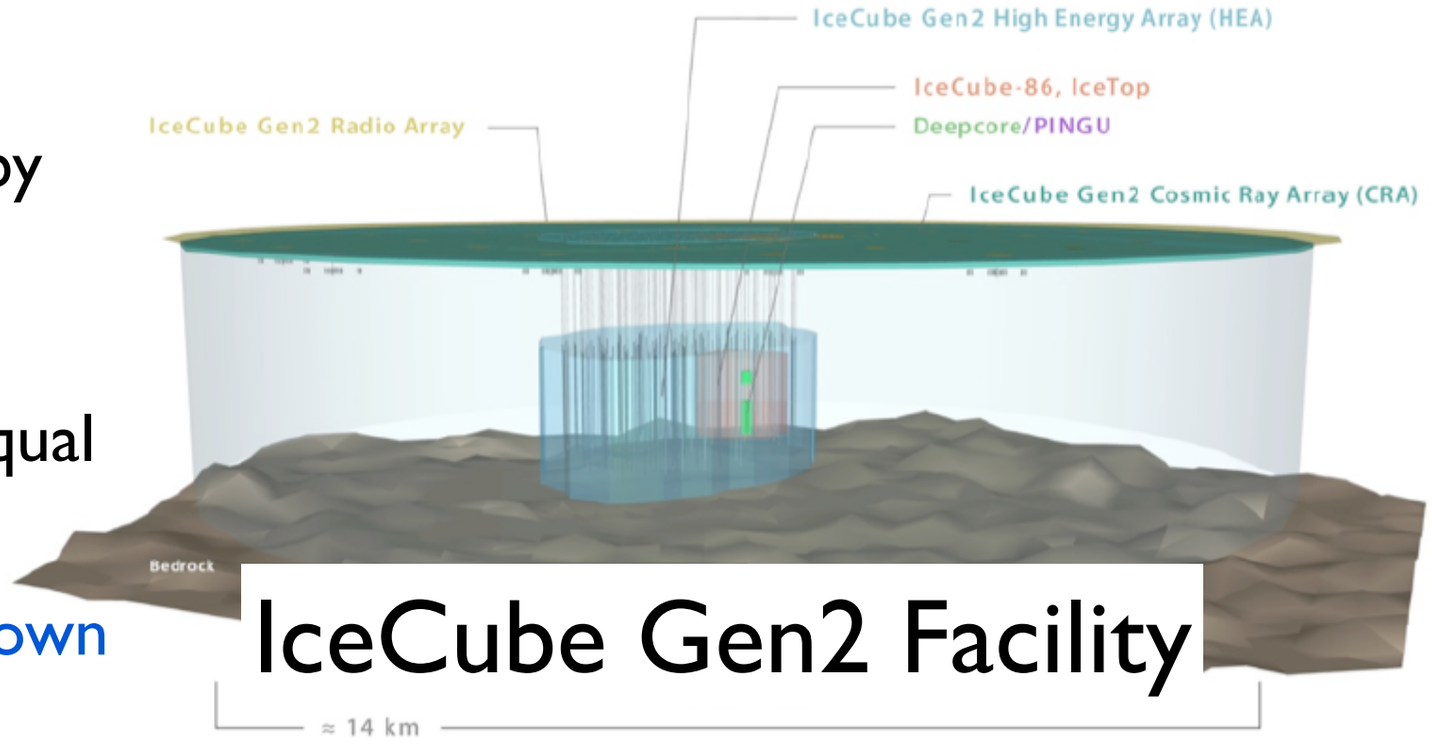
- Observations defy any simple explanation from a single generic source class
  - Multiple sources classes ?
  - Hints of new physics ?

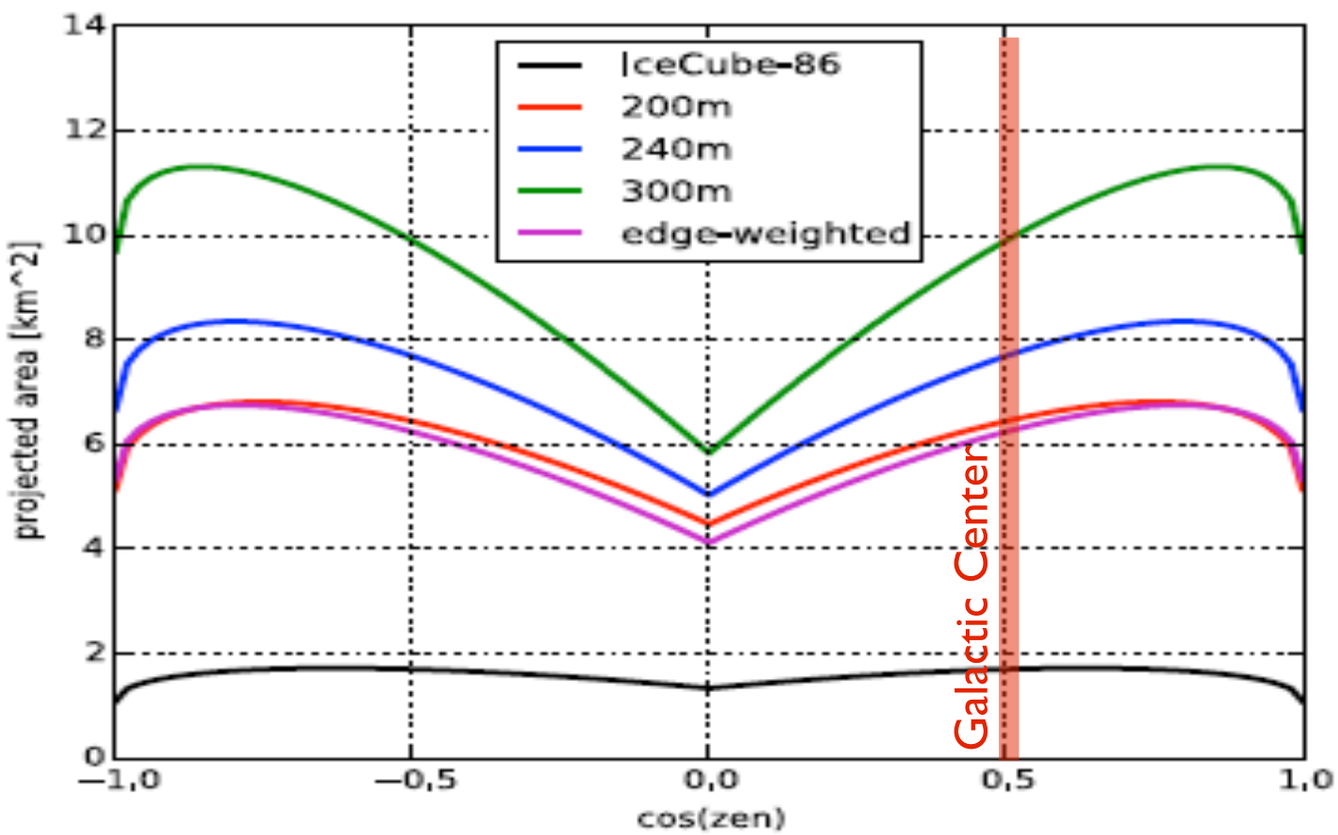
- PeV Scale Right Handed Neutrino Dark Matter
- Super Heavy Dark Matter
- Neutrino Portal Dark Matter
- Right-handed neutrino mixing via Higgs portal
- Heavy right-handed neutrino dark matter
- Leptophilic Dark Matter
- PeV Scale Supersymmetric Neutrino Sector Dark Matter
- Dark matter with two- and many-body decays
- Shadow dark matter
- Boosted Dark Matter
- …

Next generation - IceCube Gen2 Facility

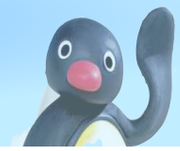
IceCube Gen2 arXiv:1412.5106

- IceCube has provided an amazing sample of events, but is still limited by the small number of events
- Observed astrophysical flux is consistent with a isotropic flux of equal amounts of all neutrino flavors
  - So far non of the analyses has shown any evidence for point sources
- Where are the point sources?
- What is the flavor composition?
- What is the spectrum? Cutoff?
- Transients ?
- Multi-messenger physics?
- GZK neutrinos?
- ....





PINGU - Precision IceCube Next Generation Upgrade



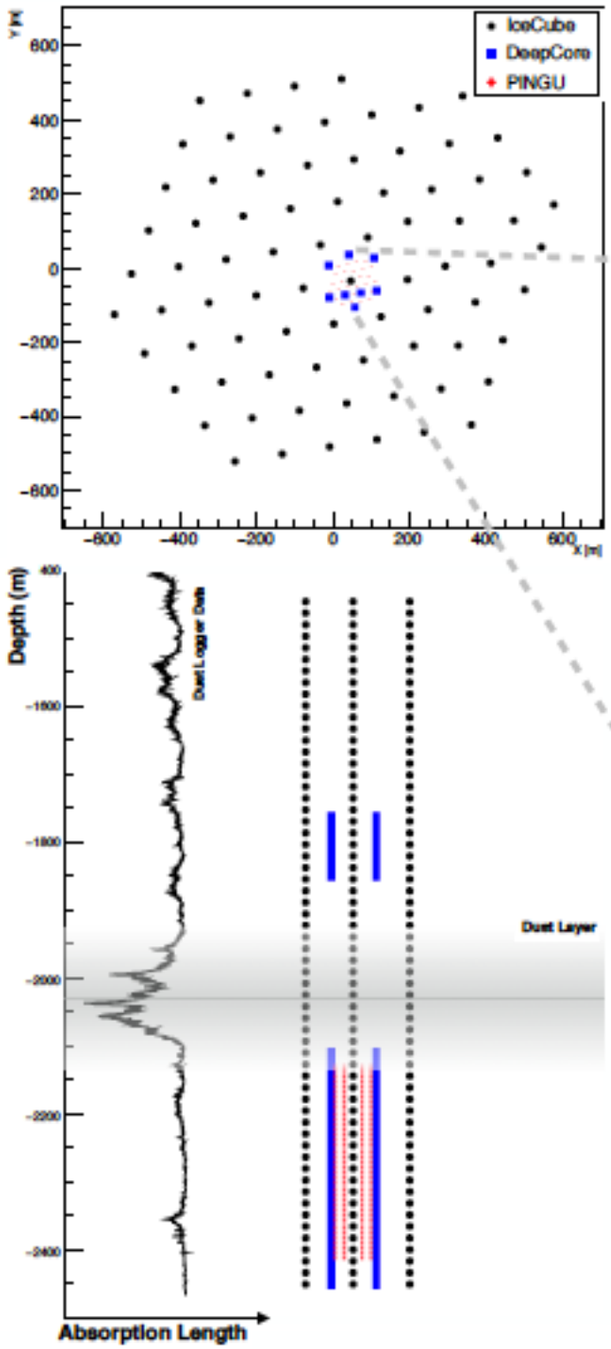
IceCube PINGU Collaboration arXiv:1401.2046

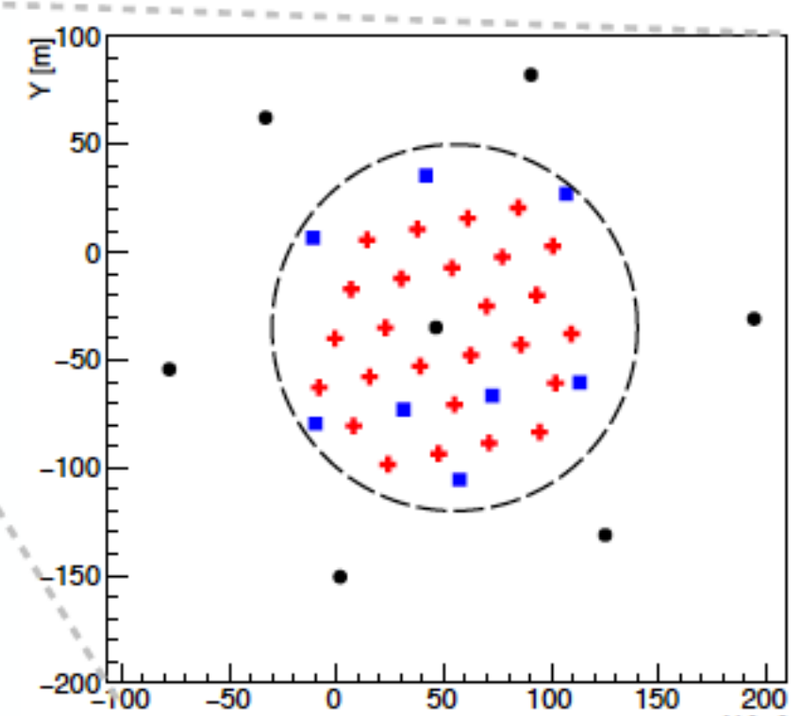
- PINGU upgrade plan
  - Instrument a volume of about 5MT with 20-26 strings
  - Rely on well established drilling technology and photo sensors
  - Create platform for calibration program and test technologies for future detectors
- Physics Goals:
  - Precision measurements of neutrino oscillations (mass hierarchy, ...)
  - Test low mass dark matter models

PINGU LOI to be updated shortly

Short version https://arxiv.org/pdf/1607.02671.pdf

New PINGU Geometry

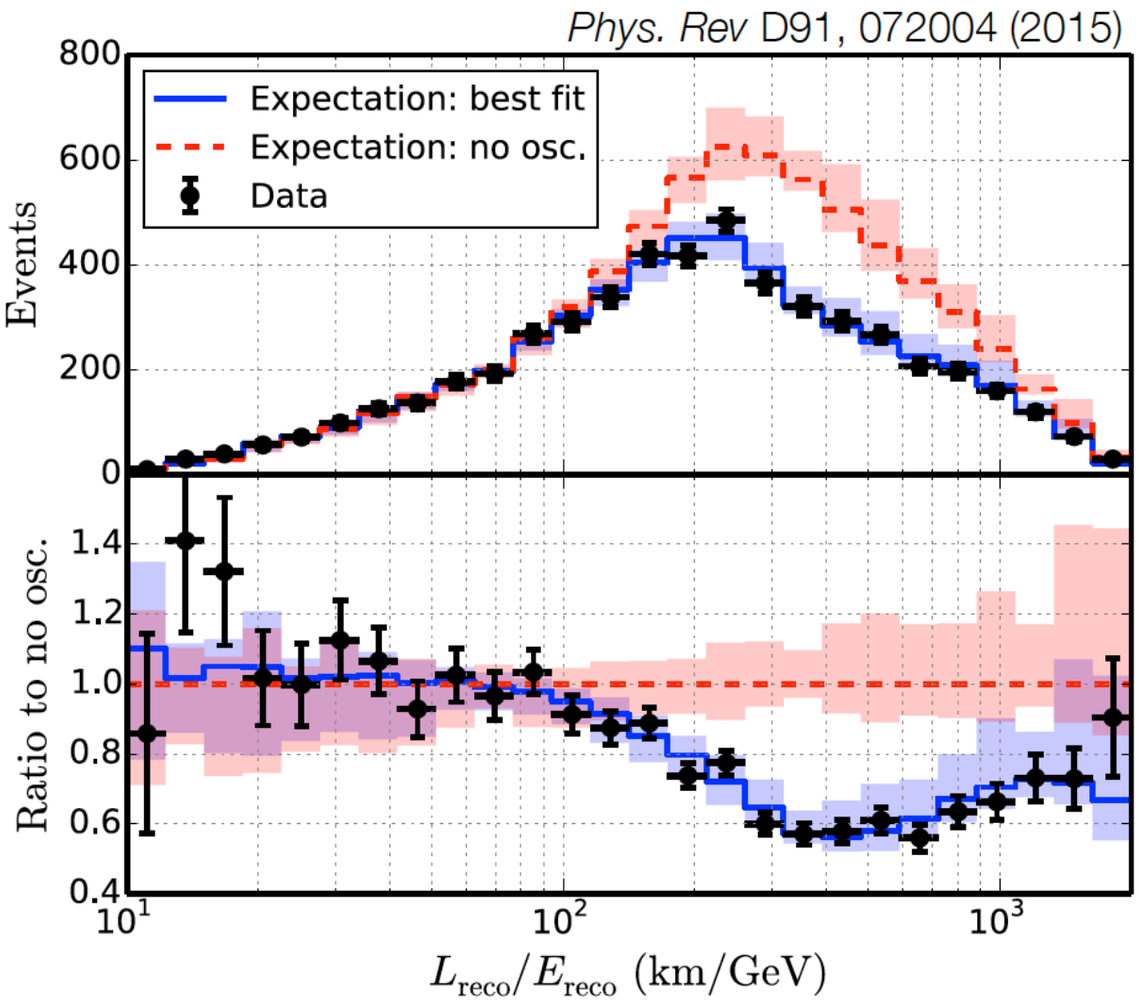


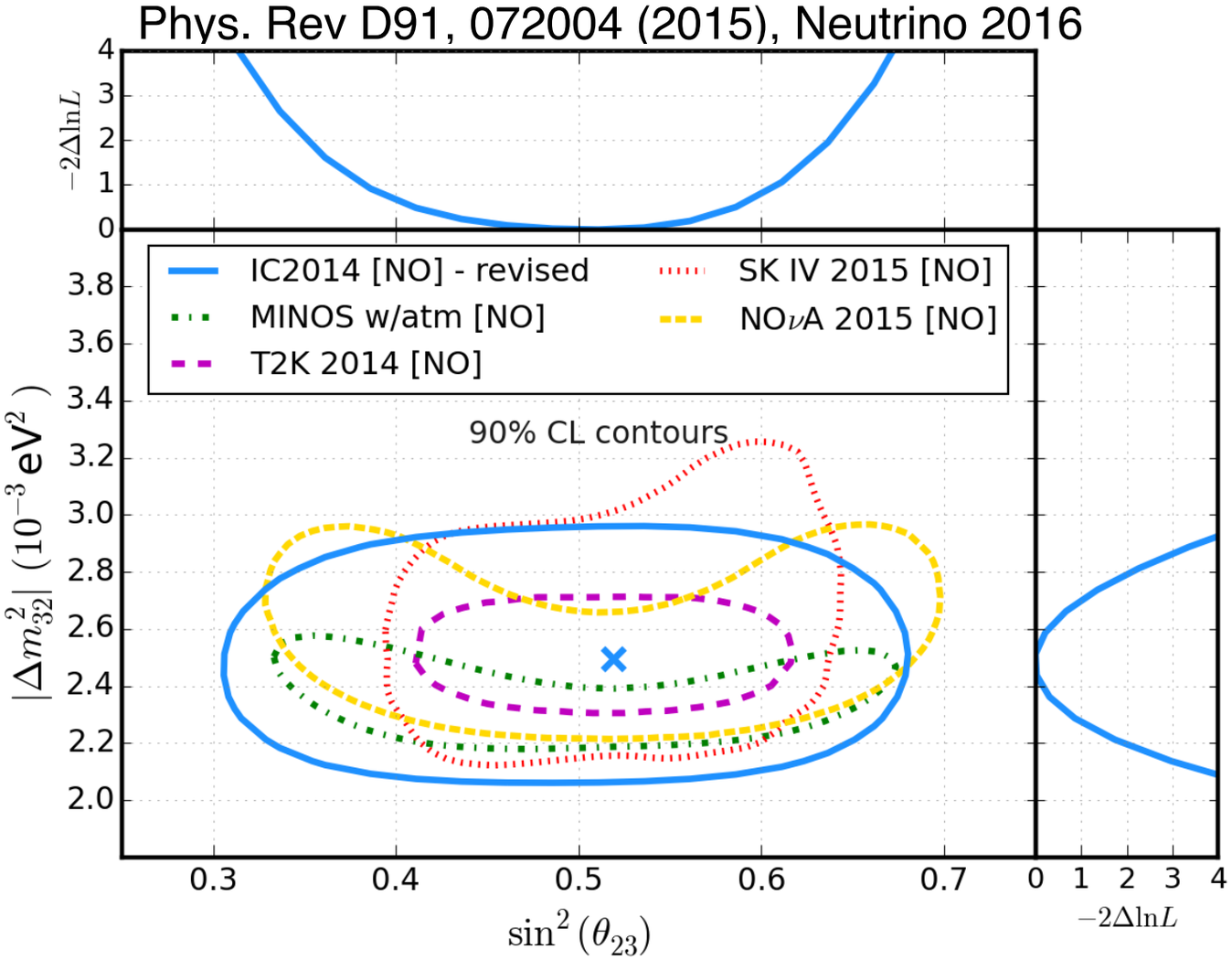


Compared to 40 string geometry

1. One less drilling/deployment season
2. Fewer holes drilled (saves fuel)
3. Fewer holes and one less season means we can refurbish drill rather than build a new one.

Neutrino Oscillation Physics with IceCube/DeepCore

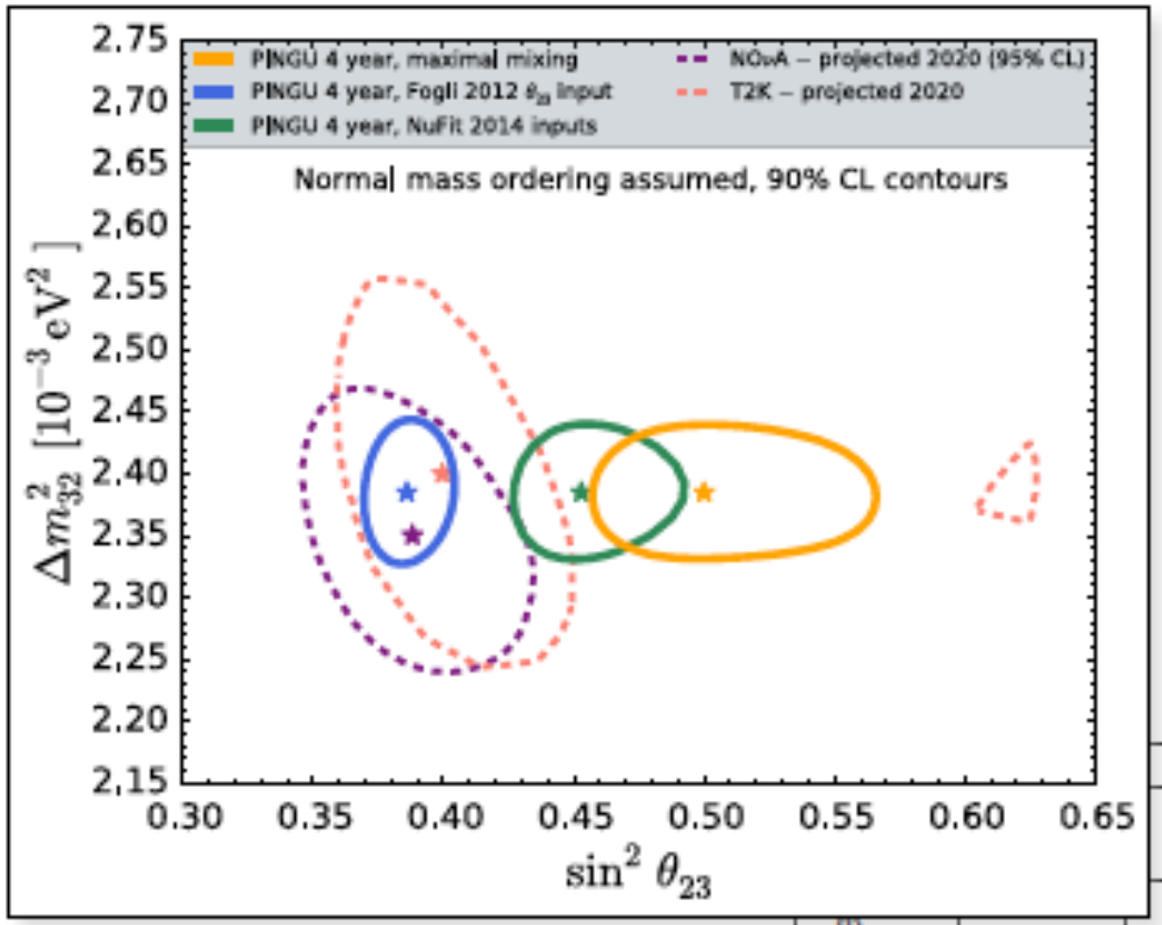




3 years of DeepCore data

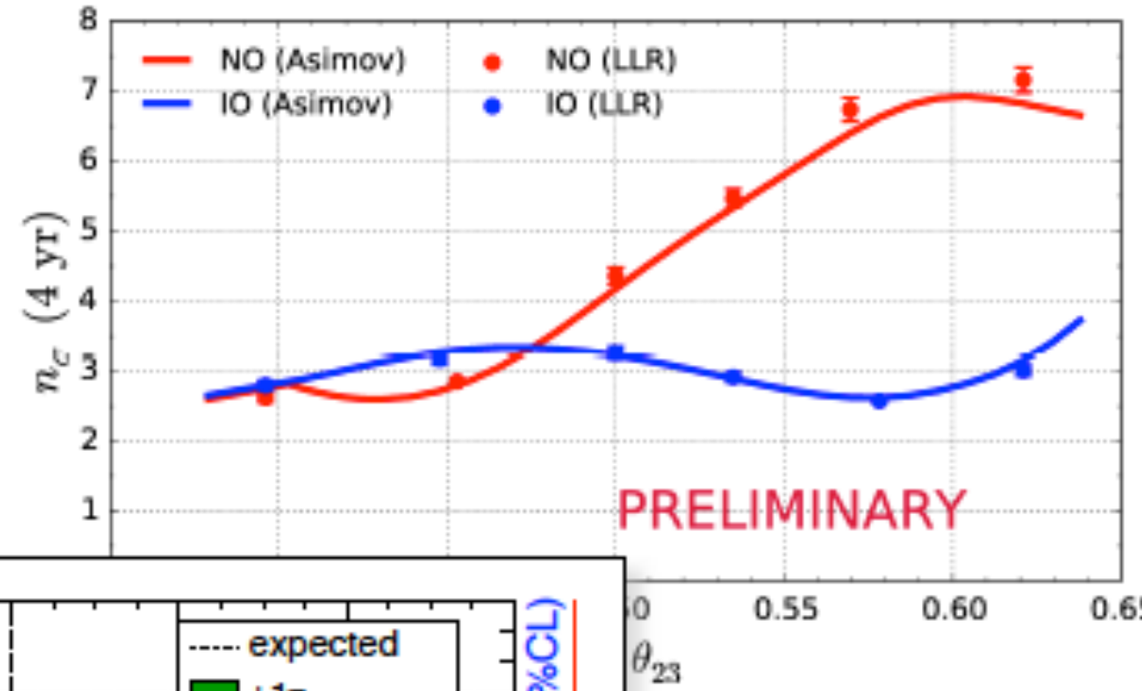
- 5174 events observed 6830 expected for no oscillation
- perform 2D fit in *E* and *cos(θ)*

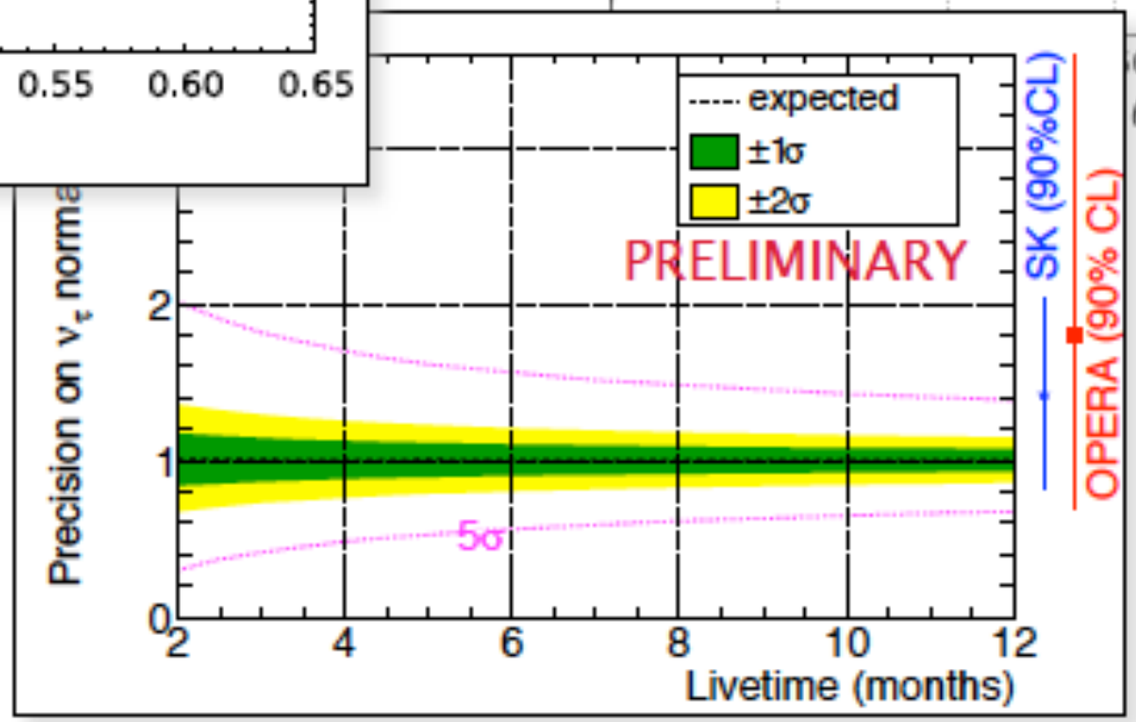
Neutrino Physics with PINGU



Measurement of mixing parameters with different method/energy range – Excellent sensitivity to octant of θ_{23}

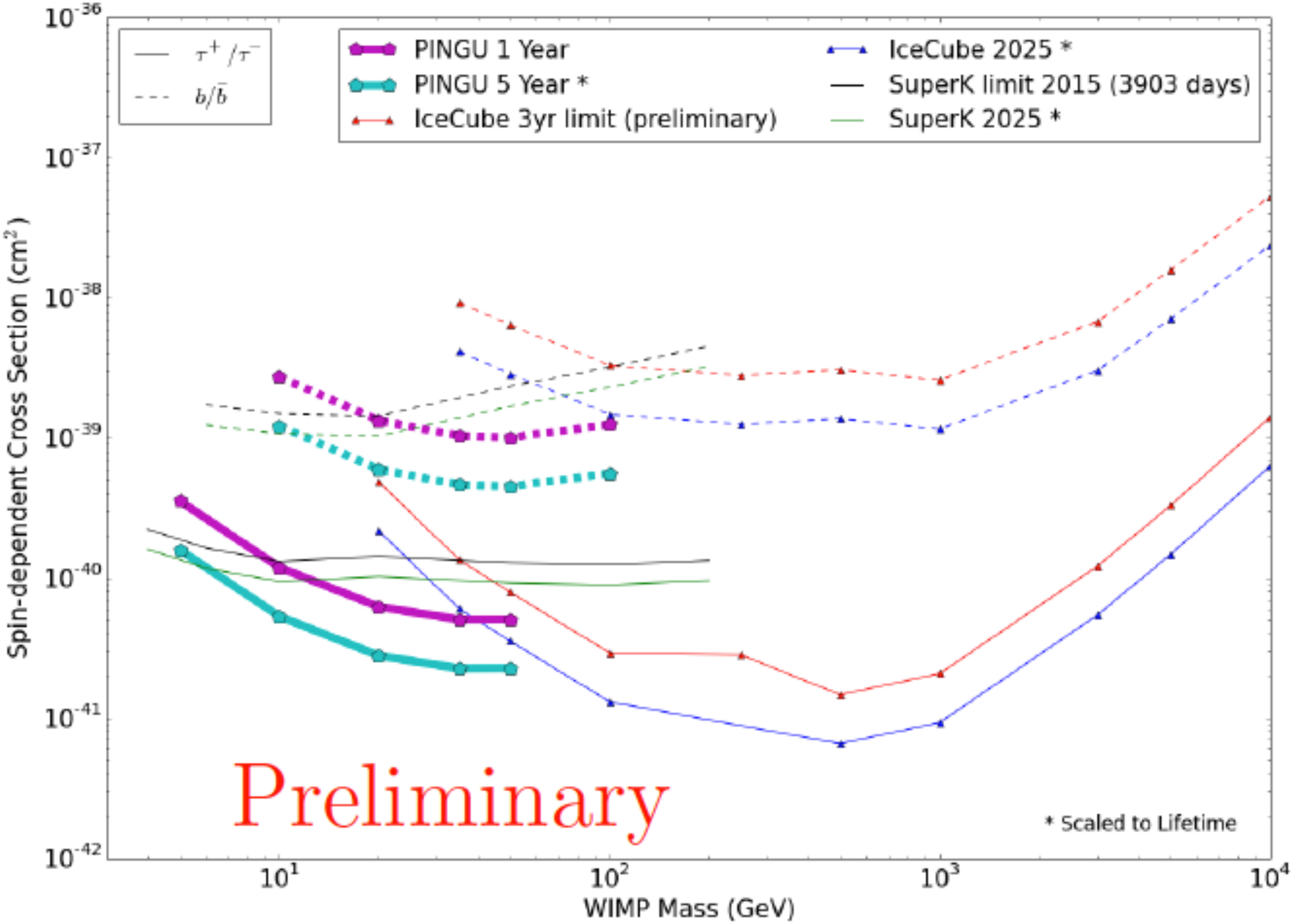
Determination of the neutrino mass ordering



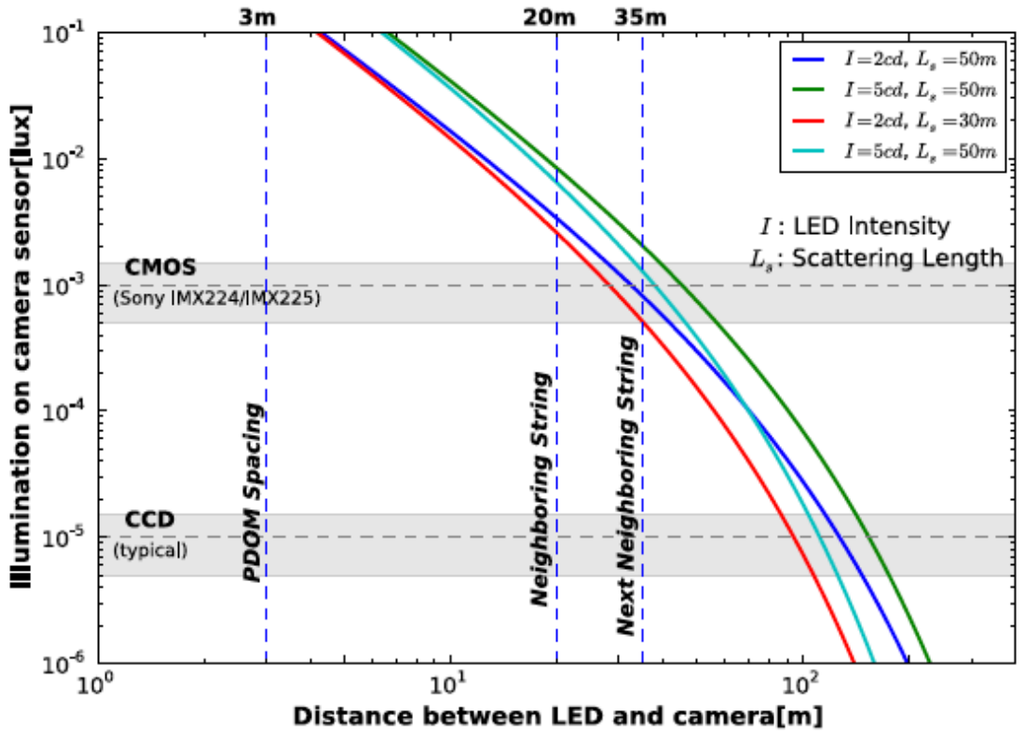


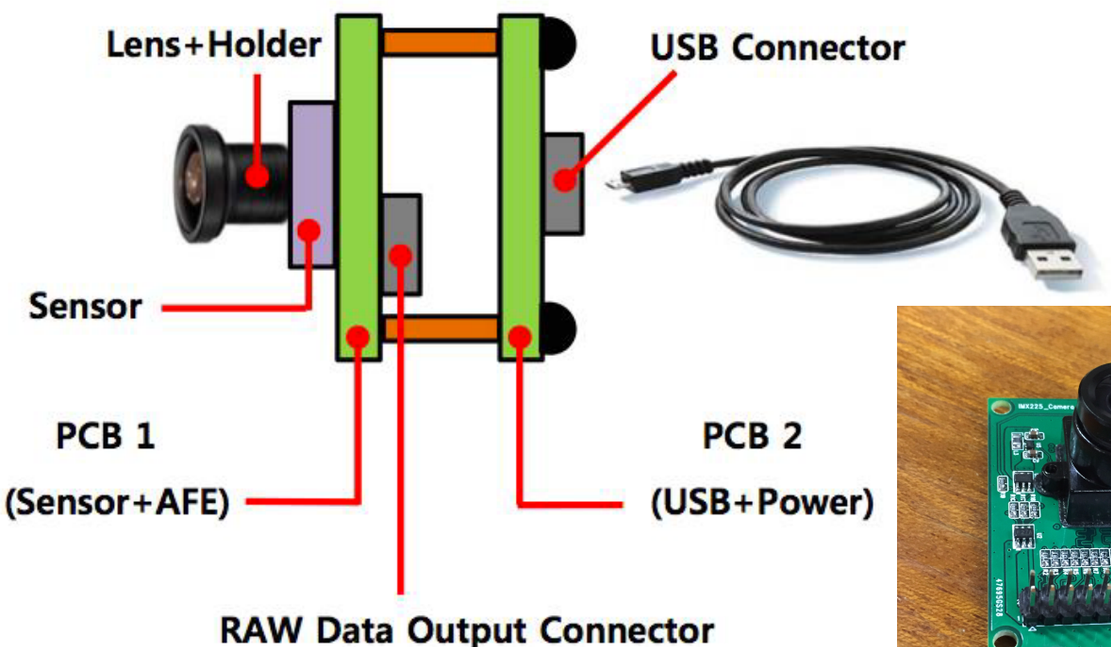
Precision measurement of ν_τ appearance – probe unitarity of PMNS matrix

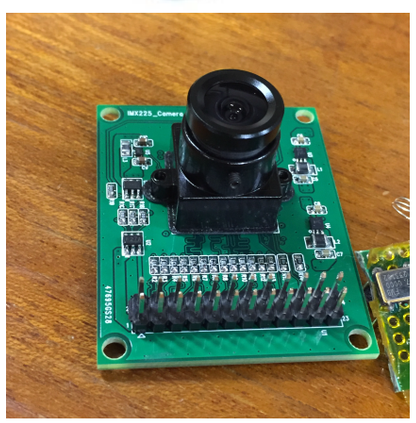
PINGU DM Sensitivity



Ice Camera System





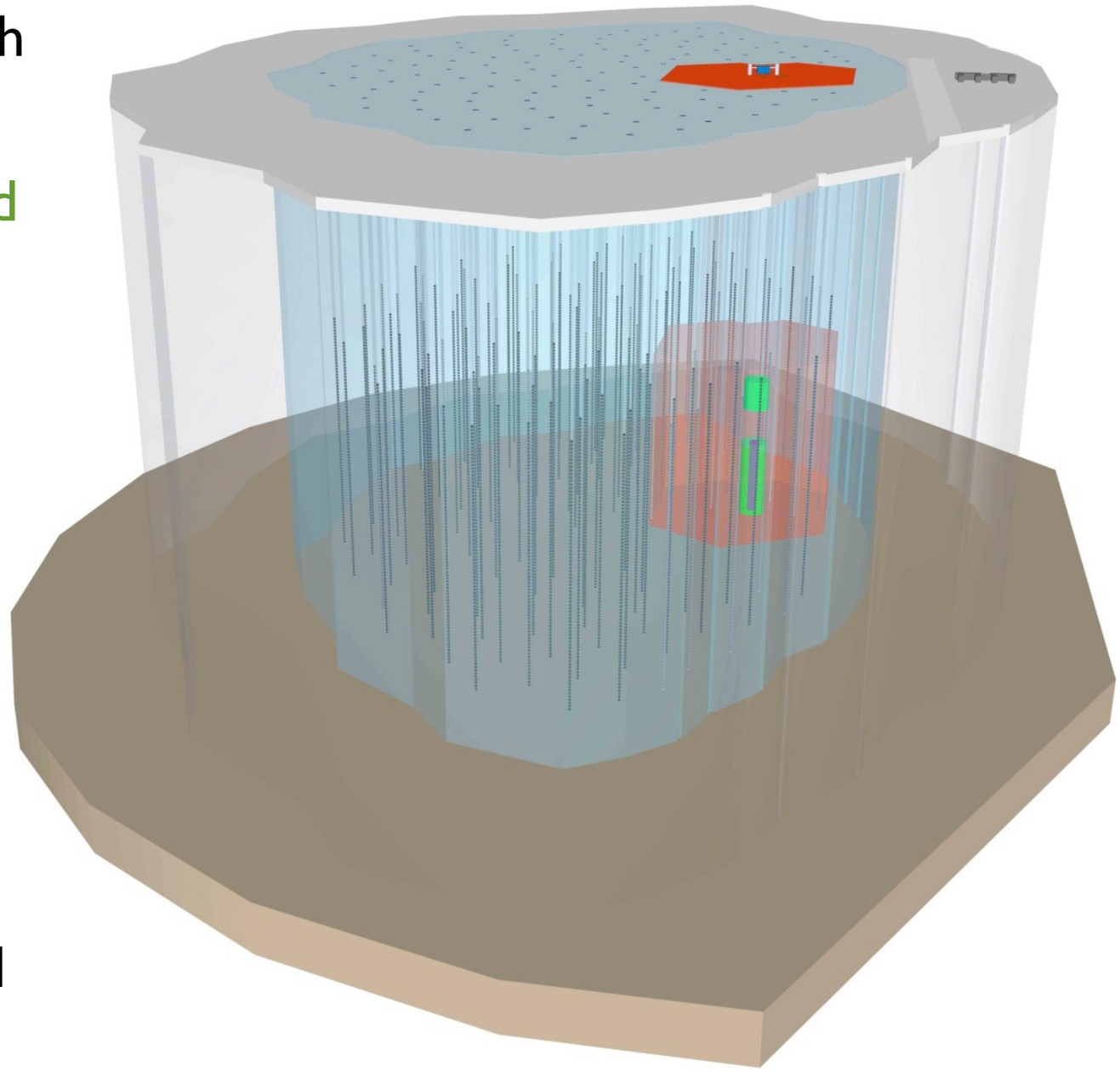


- Ice properties dominant source of sys. uncertainties for most analyses
- Low cost camera system
  - Monitor freeze in
  - Hole ice studies
    - Local ice environment
    - Position of the sensor in the hole
    - Geometry calibration
    - Survey capability

"Camera System to Study Properties of the Antarctic Ice"
D. Bose, M. Jeong, W. Kang, J. Kim, M. Kim, C. Rott.
ICRC Proceeding 2015 arXiv:1510.05228 [astro-ph.IM]

Conclusions

- Striking DM signatures might provide high discovery potential for indirect searches
- Models motivated by positron excess and gamma-ray observations can and have been tested with neutrino telescopes
- Lifetimes of heavy decaying dark matter can be constrained to 10^28s using neutrino signals
- Neutrino Telescopes provide world best limits on SD WIMP-Proton scattering cross section
- Neutrinos extremely sensitive to test low-mass WIMP scenarios at current and future detectors
- Efforts underway to expand searches beyond WIMP hypothesis …



Thanks !